

Photometry and Spectrum measurements of the Laser Guide Star beam emission at Observatorio del Roque de Los Muchachos with the Gran Telescopio CANARIAS

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ABSTRACT

In the past years ESO has built together with industry a transportable Laser Guide Star Unit (WLGSU) to test the laser technology while supporting some strategic experiments for the LGS-AO technologies. Since July 2016, the WLGSU has been installed at the Observatorio del Roque de Los Muchachos (ORM) in La Palma, next to the William Herschel Telescope (WHT), to perform different field experiments and studies related to the LGS technologies for Adaptive Optics. To study the photometric and spectral contamination from the scattering near the laser beam along the uplink at different heights, including the LGS itself, four observing runs, each of four hours, have been granted under Program GTC3-16ITP using the OSIRIS spectro-imager at the 10.4m Gran Telescopio CANARIAS (GTC). We report preliminary results of the first two runs on the photometric and spectroscopic measurements and on the observed Raman emission of atmospheric molecules along the uplink beam.

Keywords: Laser Guide Star, Adaptive Optics, Laser Tomography Adaptive Optics, Atmospheric Characterization

1. INTRODUCTION

ESO has built together with industry a transportable Laser Guide Star (LGS) Unit, i.e. a compact system, in which an experimental 20W CW laser at 589 nm based on the ESO patented Fiber Raman Amplifier technology, is attached directly to a 30 cm launch telescope mounted on high precision pointing motors ([1]). The so-called “Wendelstein Laser Guide Star Unit” (WLGSU) has been designed, assembled, tested in the ESO laboratory and finally commissioned successfully in the summer of 2011, in Bavaria, Germany ([2]). The goal of the WLGSU is to field test the laser technology while supporting some strategic experiments for the LGS-AO technologies.

The use of a 589nm laser for LGS at the Canary Islands is ongoing since January 2015, when the ESO Wendelstein Laser Guide Star Unit has been installed in Tenerife at the Observatorio del Teide to carry on return flux experiments ([3], [4], [5], [6], [7]). Since July 2016, the WLGSU has been installed at the Observatorio del Roque de Los Muchachos (ORM) in La Palma, next to the William Herschel Telescope (WHT), to perform different field experiments and studies related to the LGS technologies for Adaptive Optics ([8]).

The photometry and the emission spectrum measured on the uplink beam are important for the operations aspects of the astronomical observatories with several telescopes on site. Software tools are used to avoid that a telescope at the observatory crosses the laser uplink beam while pointing and tracking a science target.

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Figure 1. (left) the Gran Telescopio CANARIAS; (right) the ESO Wendelstein Laser Guide Star Unit at ORM.

Furthermore, it is important to avoid that photometric and spectral contamination from the scattering in the vicinity of the laser beam disturbs the observations. At which distance from the laser beam should the software prevent observations/laser propagation? The measurement of the scattering from the uplink beam is thus needed.

In order to investigate the emission spectrum and the photometry along the uplink laser beam at different heights, including the LGS itself, four observing runs, each of four hours, have been granted under Program GTC3-16ITP using the OSIRIS spectro-imager at the 10.4m Gran Telescopio CANARIAS (GTC). We report preliminary results of the first two runs, performed on 17 and 18 September 2016, on the photometric and spectroscopic measurements and on the observed Raman emission of atmospheric molecules along the uplink beam.

2. SETUP AND OBSERVING STRATEGY

OSIRIS is a spectro-imager for the optical wavelength range, located in the Nasmyth B focus of GTC. It provides standard Broad Band imaging and Long Slit Spectroscopy, and additional capability such as the Narrow Band Tunable Filters imaging, charge-shuffling and Multi Object Spectroscopy ([9], [10]). OSIRIS covers the wavelength range from 0.365 to 1.05 μm with a total field of view of 7.8 x 8.5 arcmin (7.8 x 7.8 arcmin unvignetted), and 7.5 x 6.0 arcmin, for direct imaging and MOS respectively.

For this experiment, we have used OSIRIS with two instrument modes:

1. **IMAGING:** to obtain the photometry of the laser uplink in order to measure the contamination introduced by the beam scattering.
 OSIRIS SETUP --> Sloan_r filter
 binning 2x2 - 0.254 arcsec/pixel
2. **SPECTROSCOPY:** to investigate the spectrum emission from the laser uplink beam.
 OSIRIS SETUP --> R2500R + GR + slit0.6 arcsec
 binning 1x1 - 0.508 Å/pixel

Figure 2 shows the relative positions of the GTC and the WLSU at the Observatorio del Roque de Los Muchachos, while Table 1 reports the coordinates for both. The WLSU and the GTC are about 1113 m away in SW-NE direction, with a difference in elevation (telescope pupil) of about 31 m (Figure 3).

The observations have been performed under clear or photometric conditions. We have pointed the GTC telescope toward the laser plume and the uplink beam at different heights (Figure 3). Because of the strong parallax due to the distance between the GTC and the laser launch telescope, parallax-corrected pointing coordinates for the GTC have been calculated using a Python script. We could intercept both the plume and the uplink beam at the desired height 100% of the time.

After every pointing, the laser mount and the telescope tracking have been stopped in order to acquire the data and avoid the rotation due to the parallax. A strong defocus of the GTC M2 has been necessary in order to focus the laser plume and the uplink.

Table 1. Coordinates of the GTC and the WLGSU at ORM.

	Longitude	Latitude	Elevation
GTC	17°53'30".8W	28°45'23".8N	2300 m.a.s.l.
WLGSU	17°52'54".1W	28°45'38".1N	2331 m.a.s.l.

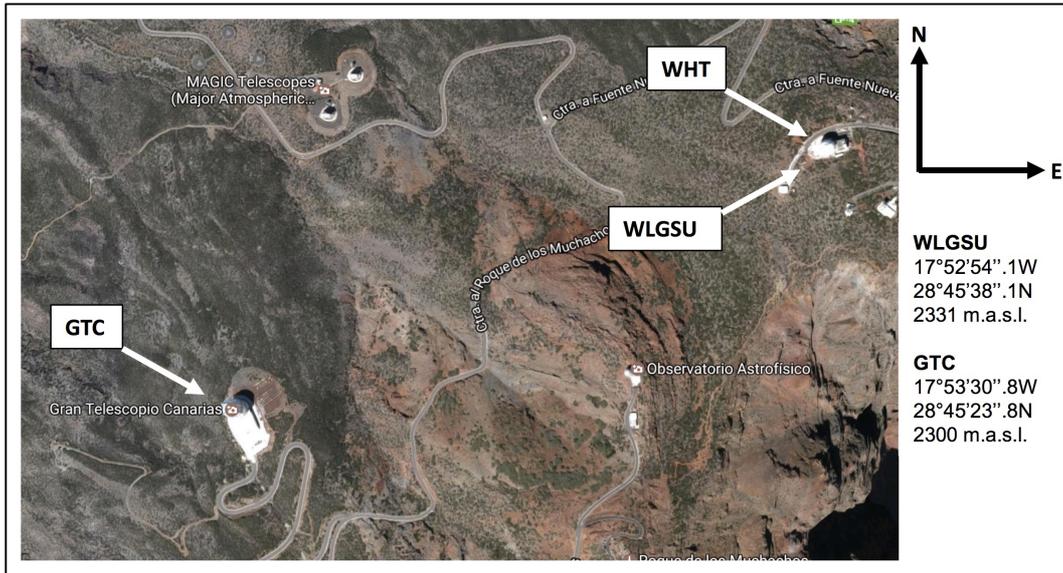


Figure 2. Relative positions of the GTC and the WLGSU at ORM.

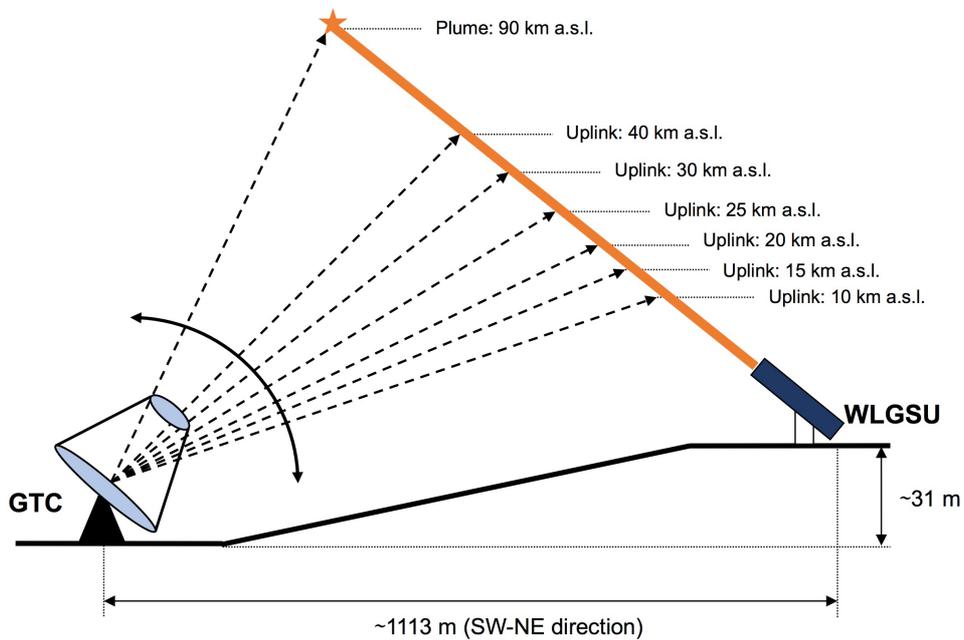


Figure 3. Schematic representation of the strategy used for the observations of the laser uplink beam and the LGS with the GTC.

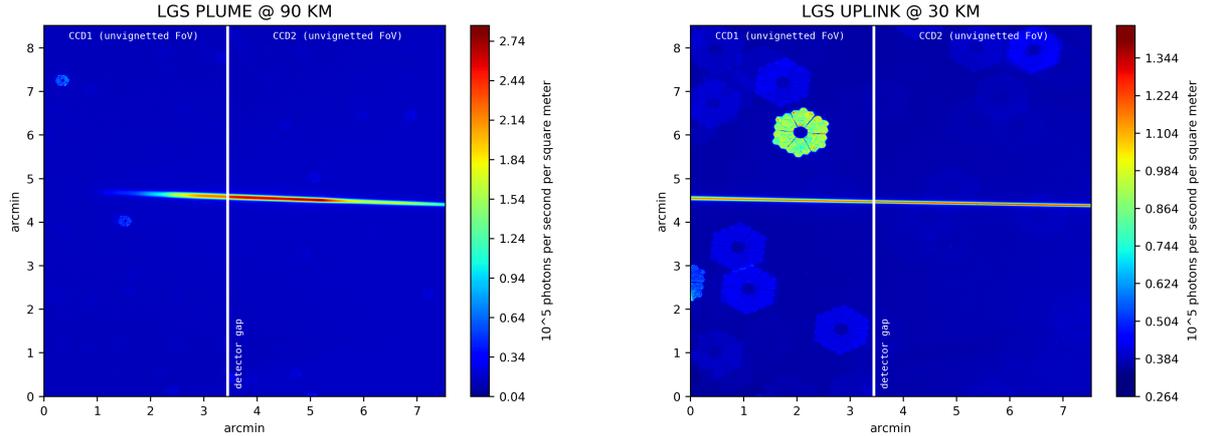


Figure 4. Fluxes for the LGS plume and the uplink beam at 30 km. The original OSIRIS calibrated images have been cut, so that the presented figures correspond to the unvignetted Field of View for both detector CCDs

3. PHOTOMETRY

For data reduction, OSIRIS Sloan_r images have been debiased, flat-fielded and corrected for distortion. The Photometric Standard SA95-193 has been acquired in the same nights and used for flux calibration. It is important to note that for these first two run, we have not acquired the sky with laser OFF, so no sky subtraction has been applied to the reduced images.

The atmospheric extinction for a wavelength of 589 nm at ORM is 0.09 mag per airmass ([11]). The measured OSIRIS magnitude for SA95-193 is 14.9795, while we expect to find 13.9491 at the ground at the star airmass. The above gives an instrument extinction of 1.03 mag, that correspond to an overall throughput of 39%.

The OSIRIS detector QE is 78% around (589 ± 200) nm, while the Sloan_r filter transmissivity at 589 nm is 90.4%, furthermore the GTC has an effective collecting area of 73 m^2 . We have converted the SA95-193 total ADU counts into the flux in photons per second per square meter. We have scaled the fluxes taking into account the difference in spectral behavior of the system between the Sloan r central wavelength (658 nm) and 589 nm.

Figure 4 show the fluxes for the LGS plume and the uplink beam at 30 km. The original OSIRIS calibrated images have been cut, so that the presented figures correspond to the unvignetted Field of View for both detector CCDs.

It is important to note that these are partial and preliminary results, since the 3rd and 4th runs of this experiment will be performed in July 2017. Furthermore, no sky subtraction has been applied to these results.

From these first results it looks that the scattering of the uplink beam might be considered negligible in an area of less than 0.5 arcmin for the beam, while for the LGS plume, the scattering extends up to about 0.5 arcmin around the plume.

4. SPECTROSCOPY

We have investigated the spectrum emission from the laser uplink beam at different heights. Data have been acquired using the R2500R grism and a slit of 0.6 arcsec. The acquisition procedure is as for the Photometry case and for the same reasons, calculating the parallax-corrected coordinates and stopping the telescope tracking after every pointing.

The OSIRIS instrument rotator has been positioned to have the slit parallel to the beam, so that the beam was superposed all along the slit (see Figure 5). For every height, we have acquired the spectra with the laser ON and OFF to retrieve the sky and subtract it from the uplink spectra in order to isolate the laser emission lines. The spectra have been flux-calibrated using the Spectrophotometric Standard g191-B2B observed during the same night.

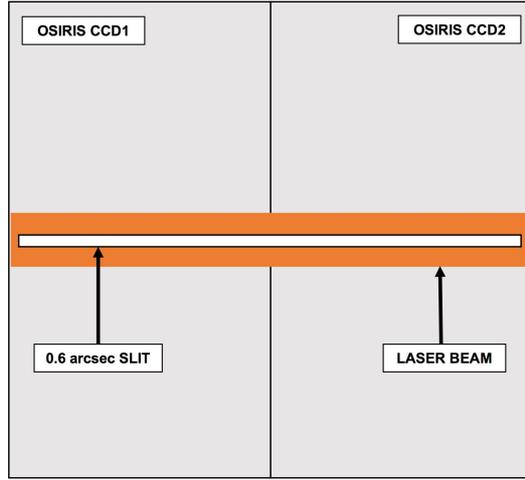


Figure 5. Schematic representation: for the Spectroscopy observations, the OSIRIS instrument rotator has been positioned to have the slit parallel to the beam, so that the beam was superposed all along the slit.

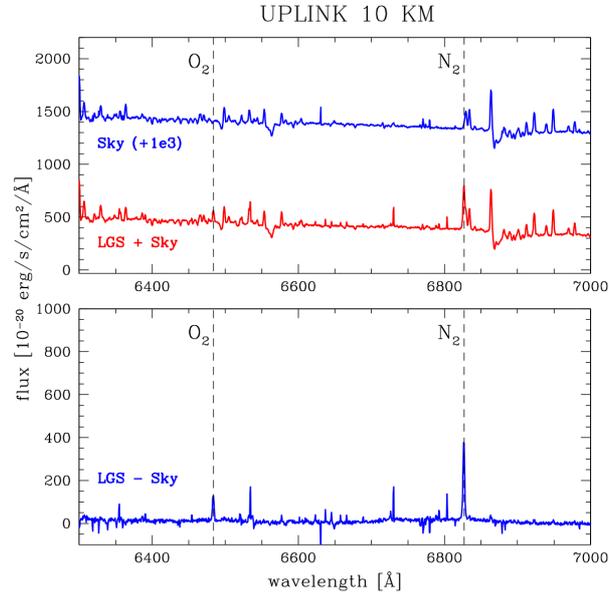


Figure 6. Spectra of the sky and the uplink at 10 km as well as the laser emission after sky subtraction. Raman emissions for O_2 and N_2 are clearly visible. Some spurious emissions from a star passing through the slit during the exposure must be ignored.

Figure 6 shows an example of spectra of the sky and the uplink at 10 km as well as the laser emission after sky subtraction. Raman emissions for O_2 and N_2 are clearly visible. Some spurious emissions from a star passing through the slit during the exposure must be ignored.

Figure 7 show the detail of the O_2 and N_2 Raman emissions for every height. The peak of the emission is at 25 km, and disappear above, while the situation is unclear for the O_2 emission at 20 km (a star passing through the slit?). The integrated fluxes are reported in Table 2.

It is important to note that these are partial and preliminary results, since the 3rd and 4th runs of this experiment have been performed in July 2017 and data have not been analyzed yet. In any case, it is interesting to report that similar results with the VLT at Paranal Observatory (at 25km only) have been reported by Vogt et al. ([12]).

Table 2. Integrated flux for LGS, O₂ and N₂ lines.

Integrated Flux [10^{-20} erg/s/cm ² /W]			
Height (a.s.l.)	10 km	15 km	20 km
LGS (5889.96 Å)	2.1×10^6	1.9×10^6	1.5×10^6
O₂ (6484.39 Å)	3.2×10^1	2.7×10^1	- - -
N₂ (6827.17 Å)	9.8×10^1	8.9×10^1	5.5×10^1

Height (a.s.l.)	25 km	30 km	40 km
LGS (5889.96 Å)	2.7×10^6	0.7×10^6	0.5×10^6
O₂ (6484.39 Å)	3.9×10^1	1.0×10^1	$<1.0 \times 10^1$
N₂ (6827.17 Å)	1.2×10^2	2.6×10^1	2.4×10^1

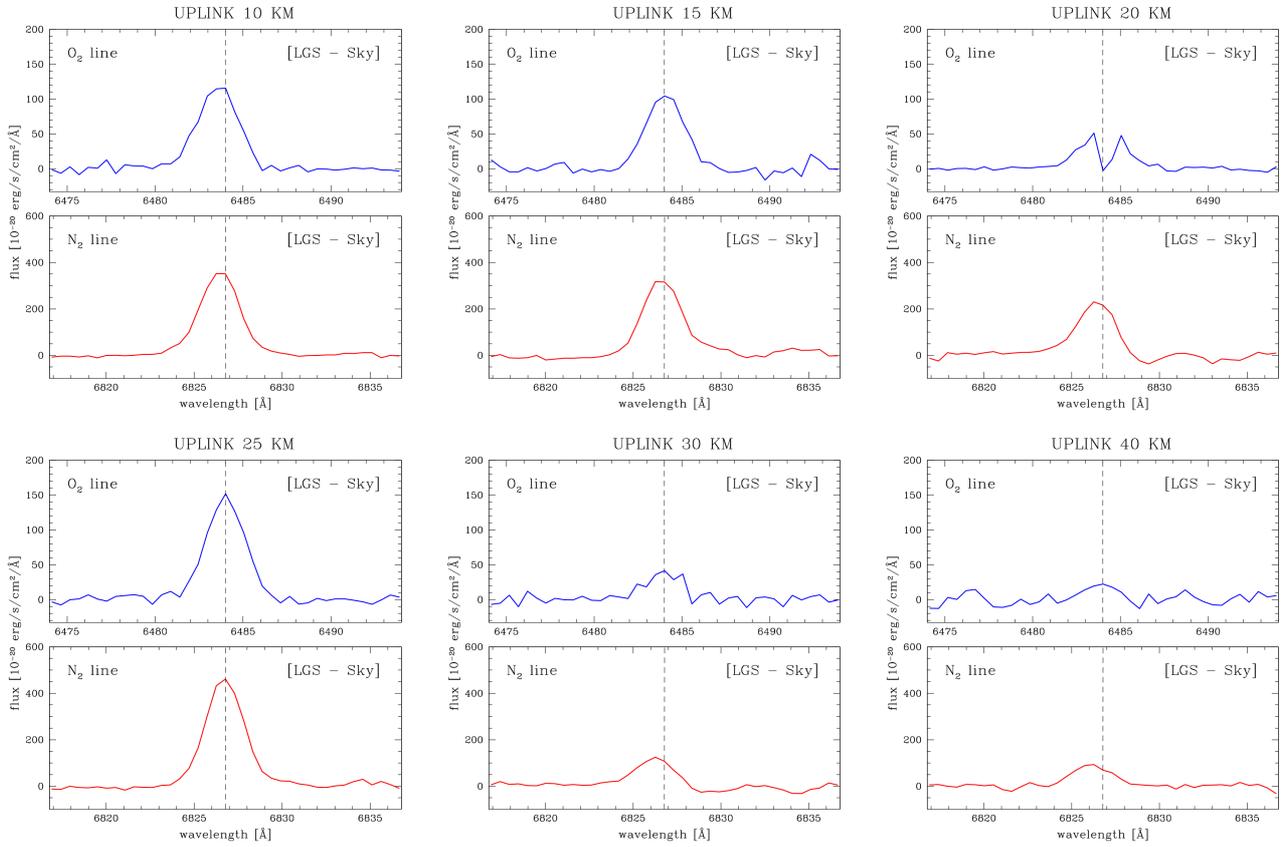


Figure 7. O₂ and N₂ Raman emissions for every height. The peak of the emission is at 25 km, and disappear above, while the situation is unclear for the O₂ emission at 20 km (a start passing through the slit?).

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