

MEASUREMENT OF LOW SIGNAL-TO-NOISE-RATIO SOLAR P MODES IN SPATIALLY-RESOLVED HELIOSEISMIC DATA

D. SALABERT^{1,*}, J. LEIBACHER^{1,2}, T. APPOURCHAUX², AND F. HILL¹

Draft version February 17, 2009

ABSTRACT

We present an adaptation of the rotation-corrected, m -averaged spectrum technique designed to observe low signal-to-noise-ratio, low-frequency solar p modes. The frequency shift of each of the $2l + 1$ m spectra of a given (n, l) multiplet is chosen that maximizes the likelihood of the m -averaged spectrum. A high signal-to-noise ratio can result from combining individual low signal-to-noise-ratio, individual- m spectra, none of which would yield a strong enough peak to measure. We apply the technique to GONG and MDI data and show that it allows us to measure modes with lower frequencies than those obtained with classic peak-fitting analysis of the individual- m spectra. We measure their central frequencies, splittings, asymmetries, lifetimes, and amplitudes. The low-frequency, low- and intermediate-angular degrees rendered accessible by this new method correspond to modes that are sensitive to the deep solar interior down to the core ($l \leq 3$) and to the radiative interior ($4 \leq l \leq 35$). Moreover, the low-frequency modes have deeper upper turning points, and are thus less sensitive to the turbulence and magnetic fields of the outer layers, as well as uncertainties in the nature of the external boundary condition. As a result of their longer lifetimes (narrower linewidths) at the same signal-to-noise ratio the determination of the frequencies of lower-frequency modes is more accurate, and the resulting inversions should be more precise.

Subject headings: methods: data analysis, Sun: interior, Sun: oscillations

1. INTRODUCTION

Our knowledge of the structure and dynamics of the solar interior has been considerably improved by the use of measurements of the properties of the normal modes of oscillation of the Sun. However, the Sun's interior is far from being fully understood, and better measurements of the mode parameters will also help to better understand the mode excitation and damping mechanisms as well as the physical properties of the outer layers by better constraining the turbulence models. A large number of predicted acoustic oscillation modes, defined by their radial orders (n) and their angular degrees (l), are not yet observed in the low-frequency range (i.e., approximately below 1800 μ Hz), because the amplitude of the acoustic modes decreases as the mode inertia increases as the frequency decreases, while the solar noise from incoherent, convective motions increases: thus the signal-to-noise ratio (SNR) of those modes is progressively reduced. Moreover, these low-frequency p modes have very long lifetimes, as much as several years, which results in very narrow linewidths, hence precise frequency measurements. Thanks to the long-duration helioseismic observations collected by the spaced-based instruments Michelson Doppler Imager (MDI) (Scherrer et al. 1995) and Global Oscillations at Low Frequencies (GOLF) (Gabriel et al. 1995) onboard the Solar and Heliospheric Observatory (SOHO) spacecraft, and by the ground-based, multi-site Global Oscillation Network Group (GONG) (Harvey et al. 1996) and Birmingham Solar Oscillations Network

(BiSON) (Chaplin et al. 1996), the frequency resolution is continuously improving and the observation of lower radial-order solar p modes is becoming possible. Their precise mode parameter determination is of great interest for improving our resolution throughout the solar interior because they cover a broad range of horizontal phase velocity, and thus a broad range of depths of penetration. Moreover, these low-frequency modes have lower reflection points in the outer part of the Sun, which make them less sensitive to the turbulence and the magnetic fields in the outer layers, where the physics is poorly understood.

The usual mode-fitting analysis consists of fitting the $2l + 1$ individual- m spectra of a given multiplet (n, l) , either individually or simultaneously. Such fitting methods fail to obtain reliable estimates of the mode parameters when the SNR of the individual- m spectra is low. Instead, various pattern-recognition techniques have been developed in an effort to reveal the presence of modes in the low-frequency range (see e.g., Schou 1998; Appourchaux et al. 2000; Chaplin et al. 2002; Broomhall et al. 2007, and references therein). In the case of spatially-resolved helioseismic data (such as GONG and MDI observations), m -averaged spectra appeared to be a powerful tool, since for a given multiplet (n, l) , there exist $2l+1$ individual- m spectra, which can result in an average spectrum with a SNR $\gg 1$ once the individual- m spectra are corrected for the rotation- and structure-induced frequency shifts. The m -averaged spectra were employed early in the development of helioseismology by Brown (1985), but were replaced by fitting the m spectra individually as the quality and the SNR of the data improved. However, years later, in order to fully take advantage of the long-duration helioseismic GONG and MDI instruments and reach lower frequencies in the solar oscillation spectrum, Schou (1998, 2002, 2004) and Appourchaux et al. (2000) used the m -averaged spectra corrected by the

¹ National Solar Observatory, 950 North Cherry Avenue, Tucson, AZ 85719

* Now at: Instituto de Astrofísica de Canarias, C/ Vía Láctea s/n, 38205 La Laguna, Tenerife, Spain, salabert@iac.es

² Institut d'Astrophysique Spatiale, CNRS-Université Paris XI UMR 8617, 91405 Orsay Cedex, France

TABLE 1
DETAILS OF THE LONG GONG AND MDI ANALYZED TIME SERIES

Series	Start Date	End Date	Fill (%)
3960-day GONG	1995 May 7	2006 Mar 9	84.6
2088-day GONG	1996 May 1	2002 Jan 17	83.7
2088-day MDI	1996 May 1	2002 Jan 17	88.9

modeled solar rotation to detect new low radial-order p modes and to set upper limits on the detectability of the g modes. These authors demonstrated the potential advantage of such rotation-corrected, m -averaged spectra.

We present here an adaptation of the m -averaged spectrum technique in which the m -dependent shift parameters are determined by maximizing the quality of the resulting average spectrum. The analysis is performed on long-duration time series of the spatially-resolved helioseismic GONG and MDI observations of the low- and medium-angular degrees ($1 \leq l \leq 35$). This range of oscillation multiplets samples the radiative interior down to the solar core. In Sec. 2, we introduce the different datasets used in this analysis. In Sec. 3, we describe this new technique in order to observe low signal-to-noise-ratio, low-frequency p modes, explaining the different steps of the analysis from the mode detection to peak-fitting. In Sec. 4, we demonstrate that this method allows us to successfully measure lower-frequency modes than those obtained from classic peak-fitting analysis of the individual- m spectra by comparing with other measurements obtained from coeval datasets. In Sec. 5, we present the mode parameters of these long-lived, low-frequency acoustic modes down to $\approx 850 \mu\text{Hz}$ extracted from the analysis of 3960 days of GONG observations using the m -averaged spectrum technique. Finally, Sec. 6 summarizes our conclusions.

2. OBSERVATIONS

Details of the spatially-resolved helioseismic observations collected by both GONG and MDI used for this work (the starting and ending dates, and their corresponding duty cycles) are given in Table 1. Coeval 2088-day observations of GONG and MDI were analyzed for oscillation multiplets with angular degrees from $l = 1$ to $l = 35$, and are then directly compared to those of Korzenik (2005) for $l \leq 25$ measurements of the same datasets. We also applied the analysis to 3960 days of GONG data ($1 \leq l \leq 35$), which constitutes so far the longest time series (≈ 11 years, spanning most of solar cycle 23) of spatially-resolved observations analyzed.

3. METHOD

An m -averaged spectrum corresponds to the average of the $2l + 1$ individual- m components of an oscillation multiplet (n, l) , thus reducing the non-coherent noise. Before averaging, each m spectrum of a given mode (n, l) is shifted by a frequency that compensates for the effect of differential rotation and structural effects on the frequencies. The m -averaged spectrum concentrates, for a given multiplet (n, l) , all of the $2l + 1$ m components, as it would be if the Sun were a purely-spherical, non-rotating object. Thus, the average of the $2l + 1$ individual- m

spectra considerably improves the SNR of the resulting m -averaged spectrum.

3.1. Determination of the shifts

The m -averaged spectrum is obtained by finding the estimates of the splitting coefficients, commonly called a -coefficients, which maximizes the likelihood of the m -averaged spectrum. The a -coefficients are individually estimated through an iterative process, with the initial values taken from a model. Thus, for a given mode (n, l) , the frequency shift $\delta\nu_{nlm}$ is parameterized by a set of coefficients, as:

$$\delta\nu_{nlm} = \sum_{i=1}^{i_{max}} a_i(n, l) P_{l,m}^i, \quad (1)$$

where $a_i(n, l)$ are the splitting coefficients, and $P_{l,m}^i$ corresponds to the Clebsch-Gordan polynomial expansion as defined by Ritzwoller & Lavelly (1991). In this definition, the odd orders of the a -coefficients describe the effects of solar rotation, while the even orders correspond to departures from spherical symmetry in the solar structure as well as to quadratic effects of rotation. Each a_i is chosen to maximize the likelihood of the m -averaged spectrum. This is performed through an iterative procedure. For a particular order i of the coefficients a_i , a range of values is scanned around its initial value, while the other $a_{i' \neq i}$ s are kept fixed to their previously estimated values.

For each scanned value of a_i , the individual- m spectra are shifted by the corresponding Clebsch-Gordan polynomials, and the mean of these $2l + 1$ shifted spectra is taken. The mean power spectrum is then fitted using a Maximum-Likelihood Estimator (MLE) minimization as described in Sec 3.2 and its likelihood determined.

For a Monte-Carlo simulation, the left panel of Fig. 1 shows the variation of the likelihood from the MLE minimization as a function of the first splitting coefficient a_1 , showing a well defined minimum which represents the best value of a_1 . The artificial power spectra were simulated following the methodology described in Fierry Fraillon et al. (1998). We also examined the sensitivity of the mode linewidth and the entropy as criteria for determining the best shifts. In our case, the entropy (Shannon 1948) can be seen as a measure of randomness in the m -averaged spectrum, \mathcal{S} , and is defined as $-\sum[\mathcal{S} \times \ln \mathcal{S}]$.

Both linewidth and entropy show well defined minima around the input value of a_1 (middle and right panels of Fig. 1 respectively). Indeed, the m -averaged spectrum gets narrower as a_1 converges to its input values of $0.4 \mu\text{Hz}$ and $a_1 = 400 \text{ nHz}$. Similar variations are obtained for all the a_i s. As detailed in Appendix A, the use of these different criteria to determine the best estimates of the a -coefficients returns consistent results.

The iteration is performed until the difference between two iterations in each of the computed a_i coefficients falls below a given threshold (such as 0.25σ in the case of a_1). Also, in order to remove any outliers, some quality checks are performed after each measure of an a_i which needs to fall within a constrained range of values. For example, a $\pm 15\%$ window around its theoretical expectation is used for a_1 . Here, we fitted only the six first a_i in the Clebsch-Gordan expansion, even though the quality of the data supports the determination of higher-order coefficients.

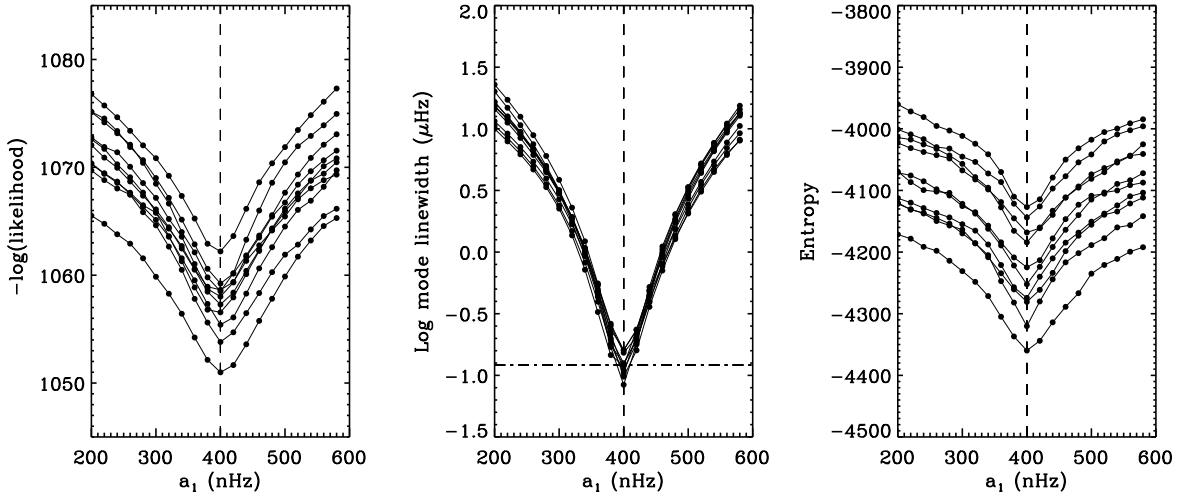


FIG. 1.— $-\log(\text{likelihood})$ (left panel), natural logarithm of the mode linewidth (middle panel), and entropy (right panel) of the m -averaged spectrum as a function of the shift coefficient a_1 obtained from ten simulated $l = 15$ spectra. The vertical dashed lines represented the introduced value of a_1 (400 nHz) in the simulations. The horizontal dotted-dashed line in the center panel represents the mode linewidth of the simulated spectra, $0.4 \mu\text{Hz}$.

Finally, low SNR peaks in the m -averaged spectrum (after adjustment) are tested against the H_0 hypothesis. In the framework of that hypothesis, the resulting spectra are tested against a statistics pertaining to pure noise (χ^2 with $2(2l+1)$ d.o.f.). This test has been widely applied to helioseismic observations in the search for long-lived, low radial-order p modes and g modes (see e.g., Appourchaux et al. 2000). In the present analysis, we rejected peaks that have a greater than 10% chance of being due to noise in the 238 analyzed windows, each containing 288 frequency bins. Here the fixed number of bins was chosen because we know that the range of theoretical frequency lie within $1.5 \mu\text{Hz}$ or so. Figures 2 and 3 illustrate the advantage of using the m -averaged spectrum technique in the case of two oscillation multiplets for 2088 days of GONG data, where the m -averaged spectra before and after the correction for the splitting coefficients are shown. These examples show the m -averaged spectra of the modes $l = 3$, $n = 5$ at $\approx 1015.0 \mu\text{Hz}$ (Fig. 2), and $l = 16$, $n = 4$ at $\approx 1293.8 \mu\text{Hz}$ (Fig. 3), as well as the corresponding $m - \nu$ diagrams. These two examples were chosen to demonstrate the performance at different SNR levels. The corresponding 10% probability levels are given. The $m - \nu$ diagrams in the case of the mode $l = 3$, $n = 5$ (right panels on Fig. 2) do not show any high SNR structure before or after correction. However, the m -averaged spectrum after correction clearly shows the target mode (lower left-panel on Fig. 2), with an unambiguous detection level. The mode $l = 16$, $n = 4$ presents a higher SNR (Fig. 3) and its $m - \nu$ diagram shows that the individual- m spectra line up after correction (lower-right panel on Fig. 3). The estimated splitting coefficients of the low-frequency modes with $1 \leq l \leq 35$ measured in the 3960-day GONG dataset are shown in Fig. 4 as a function of frequency and ν/L (with $L = \sqrt{l(l+1)}$), which is approximately proportional to the sound speed at the mode's inner turning point. Modes with selected ranges of radial orders are represented with different colors and symbols.

3.2. Extraction of the mode parameters

For a given mode (n, l) , the best estimates of the splitting coefficients determined as discussed in Sec 3.1 are used to calculate its m -averaged spectrum. When N independent power spectra are averaged together, the statistics of the mean power spectrum corresponds to a χ^2 with $2 \times N$ degree-of-freedom (d.o.f.) statistics. Appourchaux (2003) demonstrated that the mean of $2l + 1$ independent power densities, which has a χ^2 with more than 2 d.o.f. statistics, can be correctly fitted with a Maximum-Likelihood Estimator (MLE) minimization code developed for spectra following a χ^2 with 2 d.o.f. statistics. The asymmetric Lorentzian model of Nigam & Kosovichev (1998) was used to describe the m -averaged spectrum, as:

$$P_{n,l}(\nu) = H_{n,l} \frac{(1 + \alpha_{n,l} x_{n,l})^2 + \alpha_{n,l}^2}{1 + x_{n,l}^2} + B_{n,l}, \quad (2)$$

where

$$x_{n,l} = \frac{2(\nu - \nu_{n,l})}{\Gamma_{n,l}}. \quad (3)$$

Then, for a given mode (n, l) , the central frequency, the Full-Width-at-Half-Maximum (FWHM), and the power height of the spectral density are respectively $\nu_{n,l}$, $\Gamma_{n,l}$, and $H_{n,l}$. The peak asymmetry is described by the parameter $\alpha_{n,l}$, while $B_{n,l}$ represents an additive, constant background level in the fitted window. The first spatial leaks ($\delta l = 0$, $\delta m = \pm 2$), commonly called m -leaks, are also included in the fitting model and added to Eq. 2. The frequencies of the m -leaks are set from the central frequency of the target mode using the previously measured splitting coefficients (Sec. 3.1). Their peak asymmetries are assumed to be the same as that of the target mode, while their FWHMs are a free parameter of the fit and different from the target mode. The amplitude of

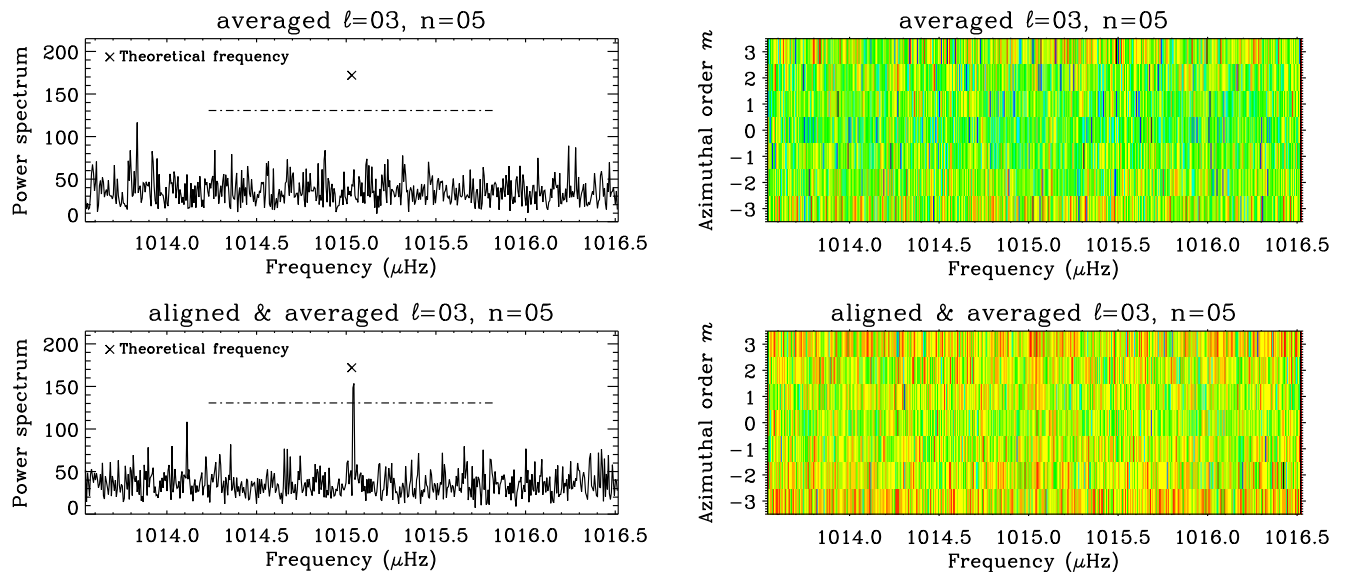


FIG. 2.— Examples of m -averaged power spectra (*left panels*) before (*top*) and after (*bottom*) correcting for the shift coefficients in the case of the oscillation multiplet $l = 3, n = 5$ observed in 2088 days of GONG data. The corresponding $m - \nu$ diagrams are also shown (*right panels*). The crosses indicate the position of the corresponding theoretical central frequency calculated from Christensen-Dalsgaard’s model S (Christensen-Dalsgaard et al. 1996). The dot-dashed lines on the left hand-side panels give the 10% probability limit that a peak is due to noise in the 238 windows, $1.5 \mu\text{Hz}$ wide. The illustrated spectral window contains the $2l + 1$ components of the represented multiplet.

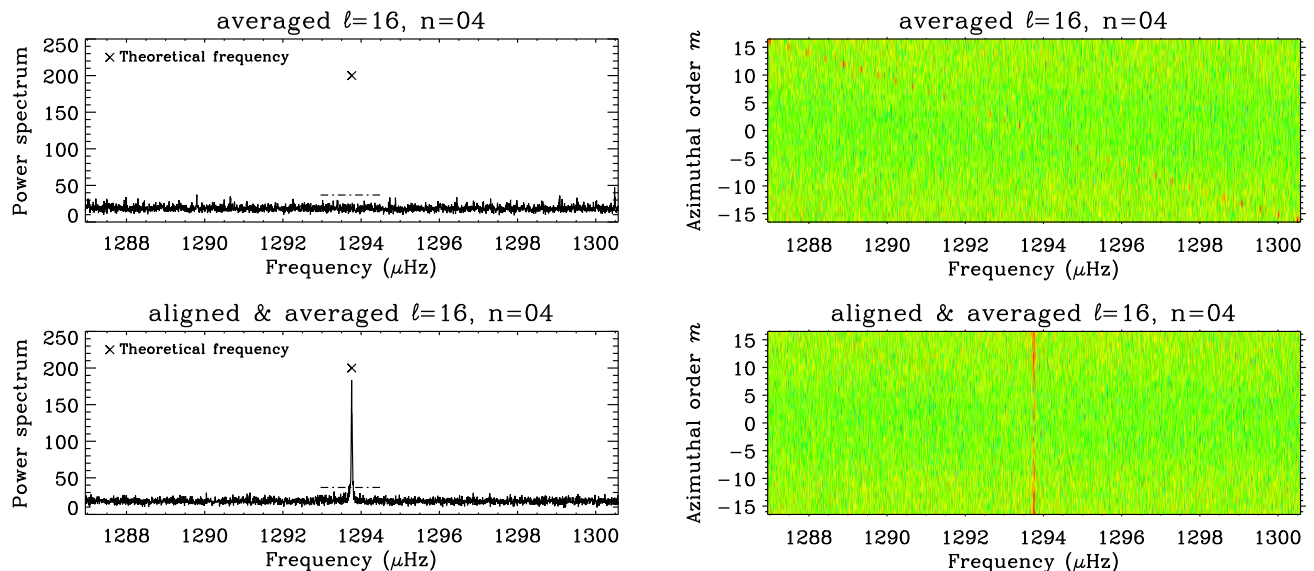


FIG. 3.— Examples of m -averaged power spectra (*left panels*) before (*top*) and after (*bottom*) correcting for the shift coefficients in the case of the oscillation multiplet $l = 16, n = 4$ observed in 2088 days of GONG data. The corresponding $m - \nu$ diagrams are also shown (*right panels*). The crosses indicate the position of the corresponding theoretical central frequency calculated from Christensen-Dalsgaard’s model S (Christensen-Dalsgaard et al. 1996). The dot-dashed lines on the left hand-side panels give the 10% probability limit that a peak is due to noise in the 238 windows, $1.5 \mu\text{Hz}$ wide. The illustrated spectral window contains the $2l + 1$ components of the represented multiplet.

the m -leaks is specified to be a fixed fraction of the central peak, which was estimated from the leakage matrix developed especially for the GONG (Hill & Howe 1998) and MDI (J. Schou, private communication) data. The first spatial leaks in the m -averaged spectrum were determined by averaging for a given multiplet (n, l) the $\delta m \pm 2$ leaks over the entire $2l + 1$ spectra.

The size of the fitting window, Ω_ν , is proportional to the first estimates of the mode width, $\Gamma_{n,l}$, and centered

around the frequency of the target mode. It is defined as:

$$\Omega_\nu = 20 \sqrt{\Gamma_{n,l}^2 + \Delta\nu_r^2} + \Delta\delta m, \quad (4)$$

where $\Delta\nu_r$ is the frequency resolution of the power spectrum. The first spatial leaks are always included in the fitting range by adding the offset $\Delta\delta m = 800$ nHz. The multiplicative factor 20 ensures a good sampling of the mode profile in the low-frequency range. A compar-

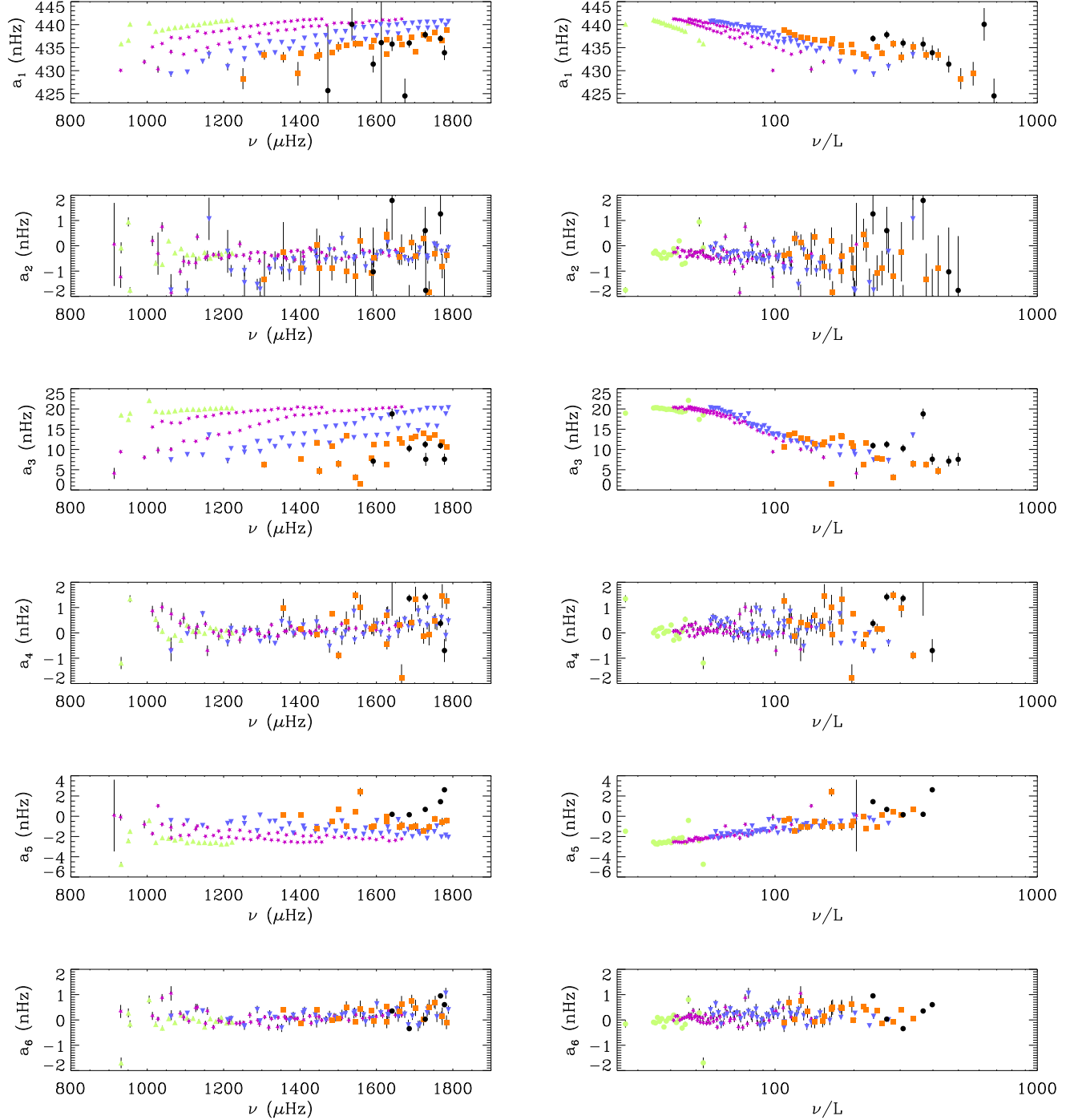


FIG. 4.— First six a_i splitting coefficients of the low-frequency p modes with $1 \leq l \leq 35$ measured in 3960 days of GONG data. The a -coefficients are represented as a function of the mode frequency ν (left column), and ν/L , with $L = \sqrt{l(l+1)}$ (right column). The different colors and symbols correspond to selected ranges of radial orders n : green triangles, modes with $n = 1, 2$; purple stars, $n = 3, 4$; blue upside-down triangles, $n = 5, 6$; orange squares, $n = 7, 8$; and black dots, $n \geq 9$.

ble definition of the fitting window was adopted by Korzennik (2005). Bad fits were removed based on a set of quality criteria based on the fitted mode parameters and associated uncertainties, such as, (1) the error of the mode frequency must be less than its mode width; (2) the SNR must be larger than 1; and (3) the mode width must be larger than the frequency resolution. A dis-

cussion on the impact of the fitting model (asymmetry, spatial leaks) on the extracted mode parameters used to describe the m -averaged spectrum can also be found in the Appendix B.

Figure 5 shows examples of the m -averaged power spectra for four different radial orders n of the multiplet $l = 17$, and the corresponding best MLE fits, which in-

cluded the mode asymmetry and the $\delta m \pm 2$ spatial leaks. The blending of the first m leaks is particularly clear as the linewidths increase with increasing frequency.

3.3. Mode parameter and a -coefficient uncertainties

The mode parameter uncertainties are established in the usual manner by the inverse of the covariance matrix. However, because the m -averaged spectrum is fitted using a MLE minimization and, as explained in Appourchaux (2003), the formal uncertainties must be normalized by the square root of the number of averaged spectra, i.e., in our case, by $\sqrt{2l+1}$. But this a-posteriori error normalization is correct only if the $2l+1$ spectra of a given (n, l) mode have the same variance (or SNR). Since the condition of equal SNR among the m spectra within a multiplet is not satisfied in our case, the uncertainties of the mode parameters have to be taken as a first approximation only. However, Monte-Carlo simulations show that this error normalization holds even in the case of m -dependent SNR (see Sec. 3.3.2).

It can also be derived that the errors on the a -coefficients can be estimated as follows:

$$\sigma_{a_i}^{-2} = \frac{l^2}{2l+1} \left(\sum_m [P_{l,m}^i(m/l)]^2 \right) \sigma_{\nu_0}^{-2}, \quad (5)$$

where i is the a -coefficient order and $P_{l,m}^i$ the associated Clebsch-Gordan polynomials. The derivation of Eq. 5 is detailed in Appendix C.

3.3.1. m/l dependence of the signal-to-noise ratio

Figure 6 shows the dependence in m/l of the SNR in the GONG data. This was obtained with modes observed in the 3960-day GONG dataset below 2000 μHz and of angular degree up to $l = 35$. Both mode amplitude and background noise depends on the azimuthal order m and can be described by polynomials with only even terms, the polynomials being different for both parameters. Note that any frequency dependence of the m/l dependence is averaged out in Fig. 6.

The m/l dependence of the SNR implies that the a -coefficients are not exactly orthogonal and that their errors are correlated (see Appendix C). However, as a first order approximation, the errors on the a -coefficients can be estimated by using Eq. 5 (see Sec. 3.3.2).

3.3.2. Validation of the error estimates: Monte-Carlo simulations

The formal uncertainties of the mode parameters and of the a -coefficients were verified through Monte-Carlo simulations. The artificial power spectra were simulated following the methodology described in Fiery Fraillon et al. (1998). In a first series of simulated spectra, the m dependence in amplitude within a given multiplet (n, l) was introduced, the SNR being symmetric in $|m|$ around the $m = 0$ spectrum. In a second series, no m dependence was introduced, i.e., a constant SNR over m . The mean values of the formal errors returned by the MLE minimization were compared to the RMS value of the corresponding fitted parameter. The Monte-Carlo simulations showed that in both cases the formal uncertainties of the m -averaged spectra determined as in Sec 3.3 using a MLE minimization are a very good approximation of the errors.

4. COMPARISON WITH OTHER MEASUREMENTS

4.1. Comparison with spatially-resolved observations ($l \leq 25$)

GONG and MDI use two independent peak-finding approaches to extract the mode parameters. Developed in the early 1990s, and mostly unchanged since, they provide mode parameters on a routine basis. Time series of 108 days are used by the GONG project (Anderson et al. 1990), while the MDI project uses 72-day time series (Schou 1999). Recently, Korzennik (2005) developed a new and independent peak-finding method of the individual- m spectra, optimized to take advantage of the long, spatially-resolved, helioseismic time series available today from both projects.

4.1.1. Mode detection: $l - \nu$ diagram

Korzennik (2005) applied his peak-fitting to extract the low- and medium-degree ($l \leq 25$) mode parameters from both GONG and MDI observations using one 2088-day long time series, as well as using five overlapping segments of 728 days. In order to compare our results obtained with the m -averaged spectrum technique, we applied the procedure described in Sec. 3 to the same 2088 days of GONG and MDI observations (Table 1). Figure 7 shows the $l - \nu$ diagrams of the low-frequency modes measured with the two different analyses in the case of the 2088-day GONG (*left panel*) and MDI (*right panel*) datasets. The modes measured by Korzennik (2005) with a classic peak-fitting method of the individual- m spectra are represented by the open circles. We considered that a given mode (n, l) from Korzennik (2005) was detected when at least two of the $2l+1$ m spectra were successfully fitted, which is enough to obtain estimates of the corresponding central frequency and first splitting coefficient a_1 . The red dots represent modes measured with the m -averaged spectrum technique which were not observed by Korzennik (2005). A significantly larger number of low-frequency modes in the 2088-day GONG and MDI datasets (respectively 45 and 14 new modes) down to ≈ 900 μHz can be measured using the m -averaged spectrum technique.

4.1.2. Mode parameter and uncertainty comparisons

In order to check the accuracy of the technique and to identify any potential bias in our analysis, we compare the central frequencies and splitting coefficients obtained by the two methods. The individual- m frequencies of Korzennik (2005) were fitted using a Clebsch-Gordan polynomial expansion (Ritzwoller & Lavelly 1991) in order to estimate the corresponding central frequencies and a -coefficients of each (n, l) multiplet. The formal uncertainties of the individual- m frequencies were used as fitting weights. The left panel on Fig. 8 shows the distribution of the differences in central frequencies below ≈ 1800 μHz of the common modes between the 2088-day GONG estimates measured using the m -averaged spectrum technique and from Korzennik (2005) (as represented on Fig. 7), demonstrating that there is no frequency dependence over the analyzed low-frequency range. The distribution was fitted by a Gaussian function, and its associated parameters (mean, standard deviation) are indicated on Fig. 8. While, on average,

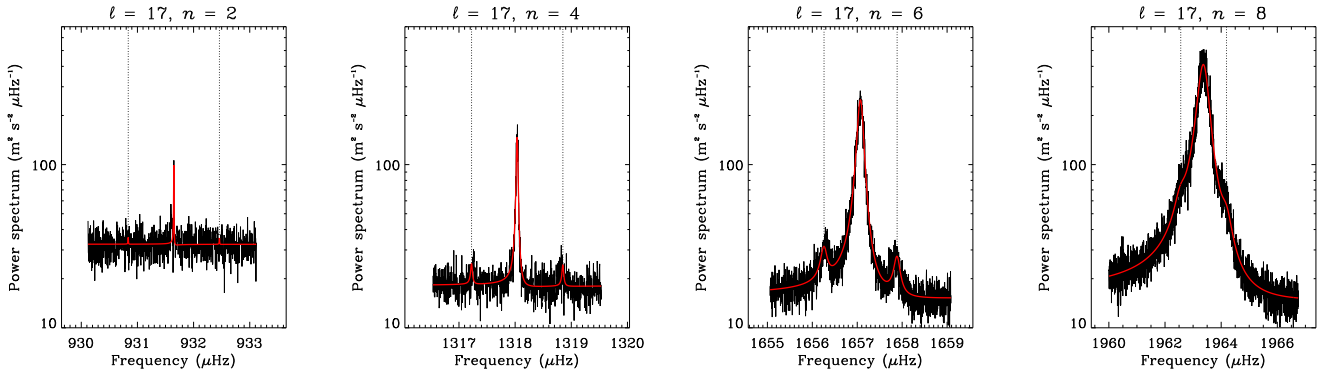


FIG. 5.— Examples of m -averaged power spectra in the case of the $l = 17$ multiplet for four different radial orders $n = 2, 4, 6,$ and 8 . The red lines represent the best MLE fits (Eq. 2), including the closest $\delta m = \pm 2$ spatial leaks, whose positions are illustrated by the dotted lines.

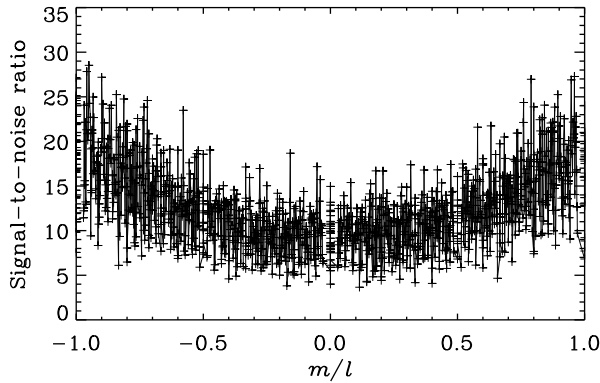


FIG. 6.— Signal-to-noise ratio for modes up to $l = 35$ as a function of m/l , obtained with the 3960-day GONG dataset.

the GONG central frequencies obtained using the m -averaged spectrum technique are less than 1 nHz smaller than Korzennik (2005)’s estimates, this offset is not significant — the corresponding standard deviation being about 5 times larger. The MDI frequencies estimated with the m -averaged spectrum technique give comparable, insignificant mean differences with Korzennik (2005). Similar results are obtained with the splitting coefficients.

We also compared the low frequencies ($\nu \lesssim 1800 \mu\text{Hz}$, see Fig. 7) estimated in both the 2088-day GONG and MDI datasets using the m -averaged spectrum technique. The right panel of Fig. 8 represents the distribution of the frequency differences of the common modes, in the sense GONG minus MDI. The mean difference value is of -0.17 ± 1.99 nHz, i.e. the GONG and MDI low-frequency modes are essentially the same. The mean difference in Korzennik (2005)’s central frequencies between the 2088-day GONG and MDI datasets for modes below $1800 \mu\text{Hz}$ is of 0.35 ± 5.40 nHz. The splitting coefficient estimates are also consistent between the two datasets with in the case of the a_1 coefficient a mean difference of -0.04 ± 0.31 nHz.

The left panel of Fig. 9 shows the formal 1σ uncertainties of the central frequencies of the measured common modes between the 2088-day GONG dataset analyzed in the present analysis and the coval 2088-day GONG dataset from Korzennik (2005) up to $\approx 1800 \mu\text{Hz}$. Our estimates of the frequency uncertainties are much smaller than those quoted by Korzennik (2005). However, Fig. 10 in Korzennik (2005) suggests that the errors are overestimated, and that his results “might be too conservative”. Korzennik (2008) reported that in the case of a 2088-day long time series, as a first estimate, a multiplicative factor of 0.75 needs to be applied to the frequency uncertainties reported in Korzennik (2005). However, despite these uncertainty scaling issues, while the Korzennik (2005)’s uncertainties show an increase with decreasing frequency from $\approx 1500 \mu\text{Hz}$, the uncertainties returned from the m -averaged spectrum technique do not show this increase, thanks to the higher SNR of the m -averaged spectrum than for the individual- m spectra.

The uncertainties on the a -coefficients returned by the m -averaged spectrum technique are also smaller than the ones obtained by fitting the individual- m spectra, as shown on the right panel of Fig. 9 in the case of the a_1 coefficients. As for the frequencies, the a -coefficients of the modes below $\approx 1500 \mu\text{Hz}$ are better constrained using the m -averaged spectrum technique.

4.2. Comparison with Sun-as-a-star observations ($l \leq 3$)

The spatially-resolved GONG and MDI instruments are not optimized to observe low-degree solar p modes below $l \leq 3$, unlike the Sun-as-a-star, integrated-light instruments such as the space-based instrument GOLF onboard SOHO and the ground-based, multi-site BiSON network. The low-degree modes are of particular interest as they reach the very deep interior of the Sun. However, the spatially-resolved observations are still able to observe such low-degree oscillations.

Low-degree ($l \leq 3$) modes down to $\approx 1000 \mu\text{Hz}$ are observed in both GONG and MDI data with the m -averaged spectrum technique as illustrated in Figs. 7 and 12. In order to test the capability and the precision of the m -averaged spectrum technique to observe low-degree, low-frequency modes in spatially-resolved data,

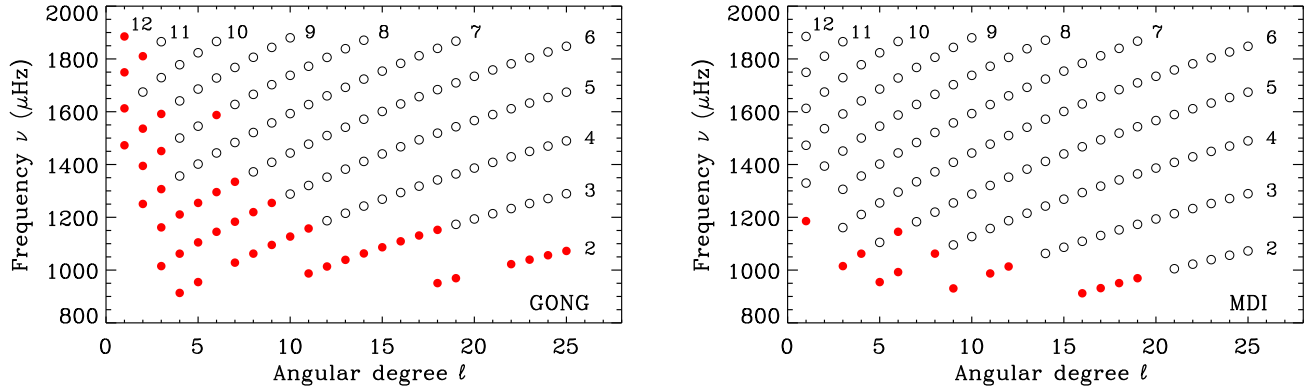


FIG. 7.— $l - \nu$ diagrams of the low-frequency modes with $1 \leq l \leq 25$ measured in 2088 days of GONG (left panel) and MDI data (right panel). The open circles represent the modes measured by Korzennik (2005) using a classic peak-fitting method of the individual- m spectra, while the red dots correspond to the additional modes measured using the m -averaged spectrum technique that were not observed by Korzennik (2005). The ridges of same radial order are also indicated from $n = 2$ to $n = 12$.

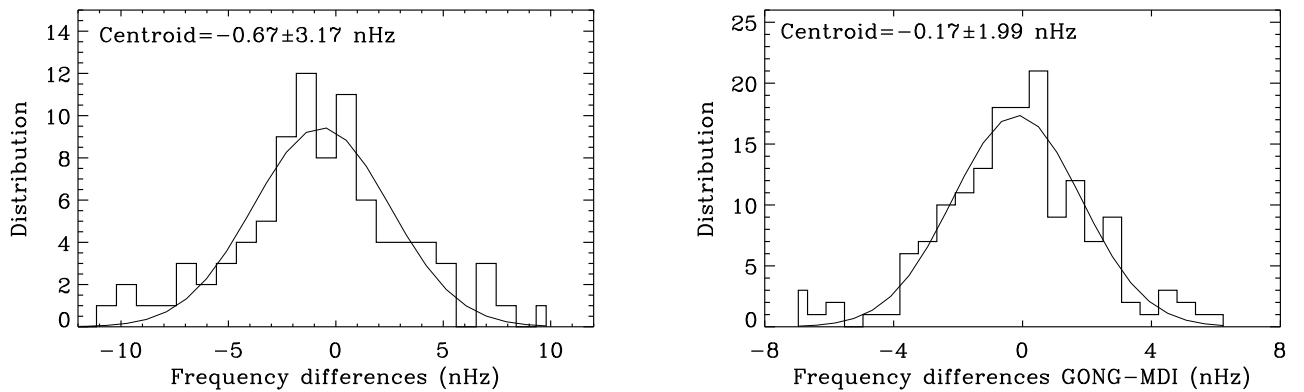


FIG. 8.— Left panel - Histogram of the differences (in nHz) in the estimated central frequencies between the 2088-day GONG dataset using the m -averaged spectrum technique and the coeval 2088-day GONG Korzennik (2005)'s estimates. Right panel - Histogram of the differences (in nHz) in the estimated frequencies between the 2088-day GONG and MDI datasets using the m -averaged spectrum technique in the sense GONG minus MDI. The corresponding Gaussian function fits and their associated mean values and standard deviations are also indicated.

measurements obtained for ≈ 11 years of the Sun-as-a-star GOLF and BiSON instruments were compared with the 3960-day GONG dataset and the 2088-day GONG and MDI datasets. The GOLF data were independently analyzed by two mode-fitting algorithms (R. A. García, private communication⁴, and P. Boumier, private communication⁵). The BiSON observations come from a combination of frequencies obtained with different long time series analyzed in order to measure low-frequency modes (W. J. Chaplin, private communication. See also Chaplin et al. 2002; Broomhall et al. 2007). However, while imaging instruments give us access to all of the $2l + 1$ individual- m components, only $l + 1$ components can be clearly observed with integrated-sunlight technique. These various components have different spatial structure over the solar surface, which can lead to differences in the extracted central frequencies of the multiplet. Chaplin et al. (2004) and Appourchaux & Chaplin (2007) derived expressions to predict the differ-

ences between the low-degree frequencies extracted from spatially-resolved (as MDI and GONG) and Sun-as-a-star (as GOLF and BiSON) observations. However, in the following, as a first approximation, we compared directly the extracted mode parameters.

The comparisons of the estimated mode frequencies and a_1 rotational splittings between the common low-degree ($1 \leq l \leq 3$), low-frequency modes in the two types of observation are shown on the left and right panels respectively of Fig. 10. The three different datasets and analysis methods give consistent results, for both the frequency and the splitting coefficient a_1 . Of course, this is only assuming that the different subsets of observed multiplets from both types of observational technique “see” the same central frequencies.

Thanks to decade-long available datasets, the low-degree, low-frequency modes are today measured lower than $1200 \mu\text{Hz}$ with high precision, demonstrated by the consistency in the extracted parameters from different instruments using distinct and independent analysis. Figure 10 also demonstrates that spatially-resolved observations can provide as accurate measurements of

⁴ From 1996 April 11 to 2006 April 18.

⁵ From 1996 April 11 to 2006 May 23.

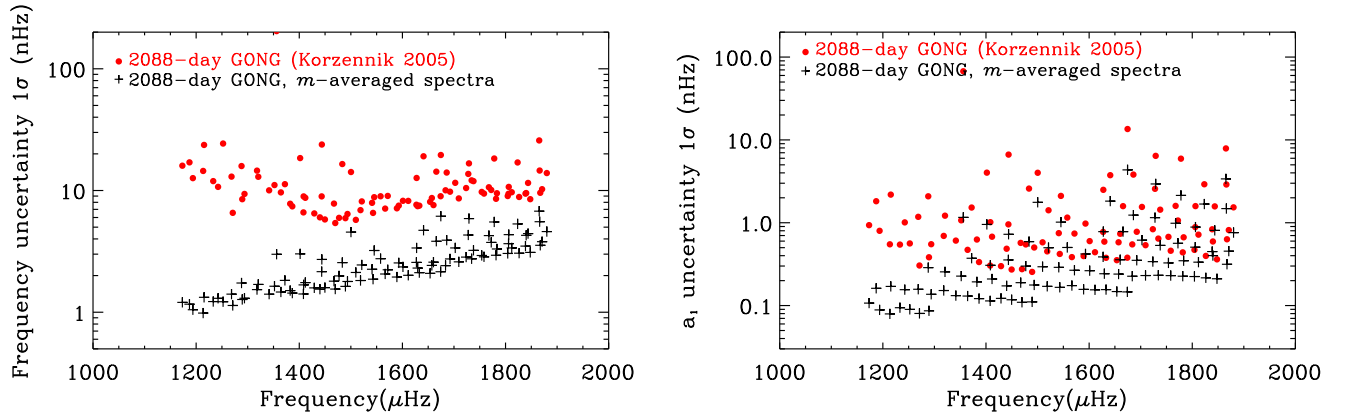


FIG. 9.— Formal uncertainties (1σ) in nHz of the central frequencies (*left panel*) and of the a_1 coefficients (*right panel*) as a function of frequency of the common modes measured in the 2088-day GONG dataset by Korzennik (2005) fitting the individual- m spectra (red dots) and by using the m -averaged spectrum technique (black plus signs).

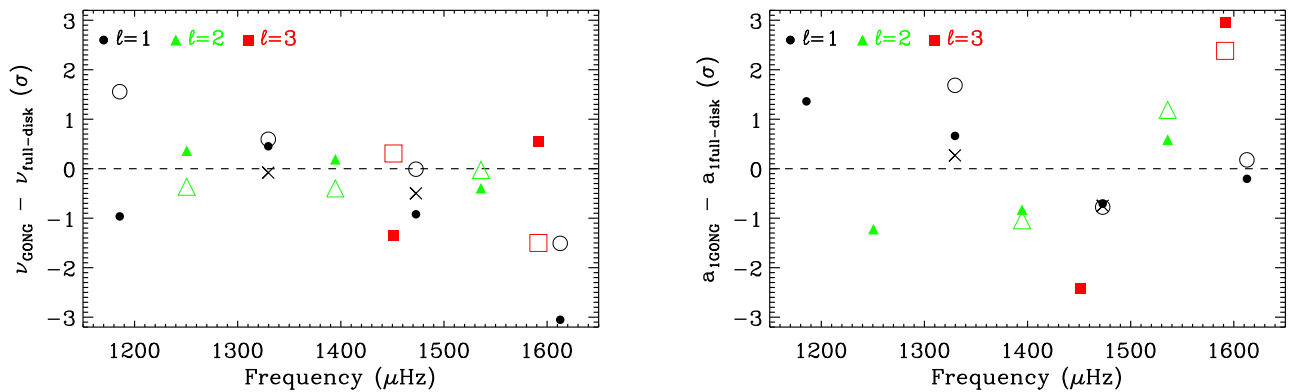


FIG. 10.— *Left panel* - Frequency differences as a function of frequency of the common low-degree ($1 \leq l \leq 3$), low-frequency p modes observed in spatially-resolved GONG and MDI data using the m -averaged spectrum technique and in two Sun-as-a-star, full-disk observations, GOLF and BiSON, over comparable periods of time. The GOLF data were analyzed by two independent peak-fitting algorithms (R. A García, filled symbols; P. Boumier, crosses). The comparison with BiSON is illustrated with the open symbols. The frequency differences were normalized by the combined errors. *Right panel* - Same as on the left panel, but for the splitting coefficient a_1 .

the low-degree modes as the Sun-as-a-star instruments do. Moreover, the m -averaged spectrum technique allows the observation of lower radial-order $l = 3$ modes than the integrated-light GOLF and BiSON observations, for commensurate observation lengths, thanks to the observations of the $2l + 1$ components (Fig. 11).

5. MODE PARAMETERS OF THE LOW-FREQUENCY OSCILLATIONS

The m -averaged spectrum technique has been applied to 3960 days of GONG observations (see Sec. 2), spanning most of the 11 years of solar cycle 23. The analysis covered low-frequency modes with angular degrees from $l = 1$ to $l = 35$. Oscillation multiplets well below $1000 \mu\text{Hz}$ were detected with good precision, such as the modes $l = 4$, $n = 4$ at $\approx 913.5 \mu\text{Hz}$; $l = 9$, $n = 3$ at $\approx 930.5 \mu\text{Hz}$; $l = 16$, $n = 2$ at $\approx 912.1 \mu\text{Hz}$; or $l = 31$, $n = 1$ at $\approx 907.5 \mu\text{Hz}$. Some examples are illustrated in Figs. 2 and 3. These low horizontal-phase-velocity modes do not penetrate deeply into the Sun, but their very high inertias afford higher precision frequencies for the inversions. It is clear from Sec. 4 that this method allows us to observe modes that are otherwise lost in the background

of each individual- m spectrum of a given multiplet (n, l), and thus unobservable with a classic peak-fitting analysis. The $l - \nu$ diagram of the observed low-frequency modes ($1 \leq l \leq 35$) down to $n = 1$ and $\approx 850 \mu\text{Hz}$ in the 3960-day GONG dataset and 2088-day GONG and MDI datasets with the m -averaged spectrum technique is shown on Fig. 12.

5.1. Mode linewidths, heights, and background levels

Figure 13 shows the fitted mode FWHMs $\Gamma_{n,l}$ (*upper-left panel*) and mode heights $H_{n,l}$ (*upper-right panel*) of the measured low-frequency oscillations. The fitted background level is also represented on the right panel. The FWHMs and heights are extremely valuable tests of models of the physical processes responsible for the mode damping and excitation by the turbulent convective motions in the outer layers of the Sun: the mode damping is inversely related to the FWHM of the mode, and the mode excitation is proportional to the mode height times the mode FWHM squared (for a detailed description, see, e.g., Salabert & Jiménez-Reyes 2006). The leveling off of the mode widths observed below $\approx 1100 \mu\text{Hz}$, despite their

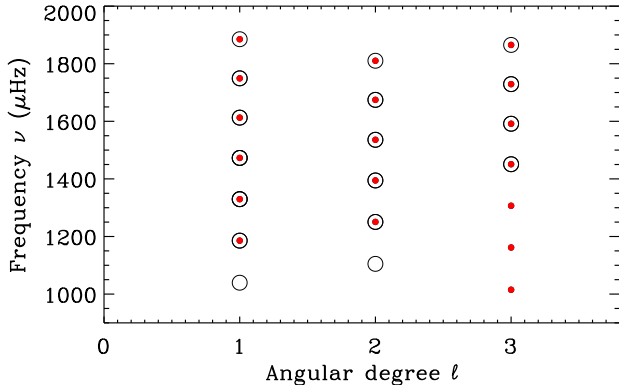


FIG. 11.— $l - \nu$ diagram of the low-frequency, low-degree ($1 \leq l \leq 3$) modes measured with the integrated-light instruments GOLF and BiSON (open circles) by R. A. García, P. Boumier, and W. J. Chaplin (private communications) and with the spatially-resolved instruments GONG and MDI using the m -averaged spectrum technique (red dots), over comparable periods of time.

dispersion becoming larger, could be a resolution effect, the peaks being then so narrow that the limiting resolution of the spectrum becomes an issue. Moreover, Schou (2004) did not observe such behaviour at low frequency in MDI data with a 2952-day time series.

As indicated by different colors and symbols on Fig. 13, the fitted mode widths follow ridges for equal radial orders n . This dependence on angular degree (l) is directly related to the mode inertia (I) in terms of a power law, as illustrated on the lower-left panel of Fig. 13. The l dependence in the mode FWHMs is removed when represented as a function of the mode inertia I .

5.2. Mode asymmetry

The mode parameters extracted through the routine GONG and MDI peak-fitting pipelines are obtained by use of symmetric Lorentzian profiles (Anderson et al. (1990) and Schou (1992) respectively). However, it was demonstrated that ignoring the peak asymmetry in the description of the acoustic modes leads to biases in the estimated mode parameters (see Appendix B and Thiery et al. 2000). Today, most of the estimates of the mode asymmetries have been restricted to low degrees ($l \leq 3$) only, from Sun-as-a-star, integrated-sunlight observations. However, Korzennik (2005) used asymmetric profiles and presented estimates of the peak asymmetry for modes with angular degrees $1 \leq l \leq 25$, obtained with GONG and MDI observations. Recently, Larson & Schou (2008) are planning to reprocess all the MDI medium- l data including a set of corrections and improvements (such as the mode asymmetry) in the MDI pipeline algorithm itself.

The asymmetry parameter ($\alpha_{n,l}$) in the low-frequency range, obtained by fitting the 3960-day GONG m -averaged spectrum ($1 \leq l \leq 35$), is shown on the lower-right panel of Fig. 13. The extracted peak asymmetry is well constrained down to $\approx 1400 \mu\text{Hz}$, with a mean value of about -0.044 ± 0.002 , and no discernable l dependence. The average asymmetry observed in the m -averaged spectrum is consistent with other measurements. For instance, the mean value observed by Korzen-

nik (2005) was about -0.04 for modes below $2000 \mu\text{Hz}$ and $l \leq 25$, once his estimates are transformed back into the Nigam & Kosovichev (1998)'s definition of the peak asymmetry. A comparable mean value is also observed at the lowest frequencies for which asymmetries were reported in Sun-as-a-star, integrated-sunlight observations (e.g., Thiery et al. 2000).

5.3. Mode frequencies

Figure 14 shows the frequency differences (in μHz) between the fitted low-frequency modes observed in the 3960-day GONG dataset using the m -averaged spectrum technique and the corresponding theoretical values calculated from Christensen-Dalsgaard's model S (Christensen-Dalsgaard et al. 1996). The corresponding frequency uncertainties were multiplied by 20 to render them visible. These comparisons are represented as a function of the angular degree (*left panel*), of the frequency (*middle panel*), and of the inner turning point (*right panel*). Modes of equal radial orders are connected. As these differences between observed and theoretical frequencies show, there is still room to improve the model of solar internal structure. Note that the right panel on Fig. 14 illustrates also the wide range of depths of penetration that these low-frequency modes cover.

6. CONCLUSION AND DISCUSSION

We presented here an adaptation of the rotation-corrected, m -averaged spectrum technique to observe low signal-to-noise-ratio, low-frequency solar p modes in spatially-resolved helioseismic data. For a given multiplet (n, l), the shift coefficients describing the differential rotation- and structural-induced effects are chosen to maximize the likelihood of the m -averaged spectra. The average of the $2l + 1$ individual- m spectra can result in a high signal-to-noise ratio when the individual- m spectra have a too low signal-to-noise ratio to be successfully fitted. This technique was applied to long time series of the spatially-resolved GONG and MDI observations for low-frequency modes (i.e., approximately below $1800 \mu\text{Hz}$) with low- and intermediate-angular degrees ($1 \leq l \leq 35$). We demonstrated that it allows us to measure lower frequency modes than with classic peak-fitting analysis of the individual- m spectra. Figure 15 shows the new low-frequency solar p modes observed in spatially-resolved data using the m -averaged spectrum technique in long time series of both GONG and MDI observations. Their central frequencies and splitting a_1 coefficients, as well as their associated uncertainties are indicated in Table 2. These normal modes of oscillation were predicted but were not measured previously. The potential of the m -averaged spectrum technique for increasing our knowledge of the solar interior is clearly illustrated on Fig. 15, p modes well below $1000 \mu\text{Hz}$ being measured with a high accuracy thanks to their longer lifetimes. We also demonstrated that the m -averaged spectrum technique returns unbiased results with no systematic differences with other long-duration measurements, which also include the asymmetry in the mode profile description.

The oscillation parameters of these low signal-to-noise-ratio, low-frequency modes, such as their central frequencies, splittings, asymmetries, lifetimes, and heights were measured. These low-frequency p modes contribute to improve our resolution throughout the solar interior

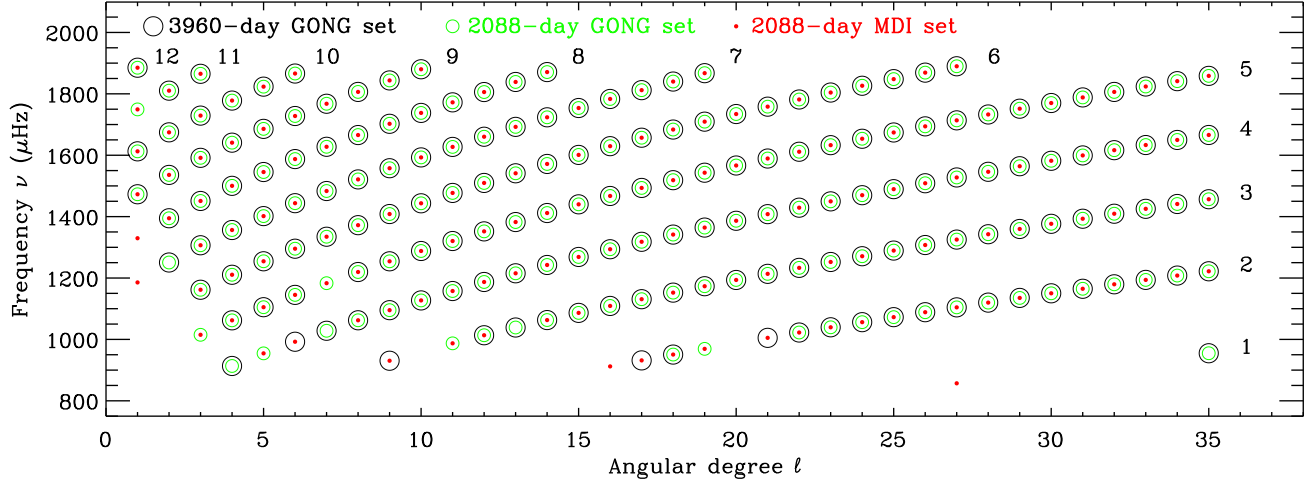


FIG. 12.— $l-\nu$ diagram of the low-frequency solar p modes from $l = 1$ up to $l = 35$ observed with the m -averaged spectrum technique in 3960 days of GONG observations (large black circles), and 2088 coeval days of GONG (medium green circles) and MDI (red dots) observations. The ridges of same radial order are also indicated from $n = 1$ to $n = 12$.

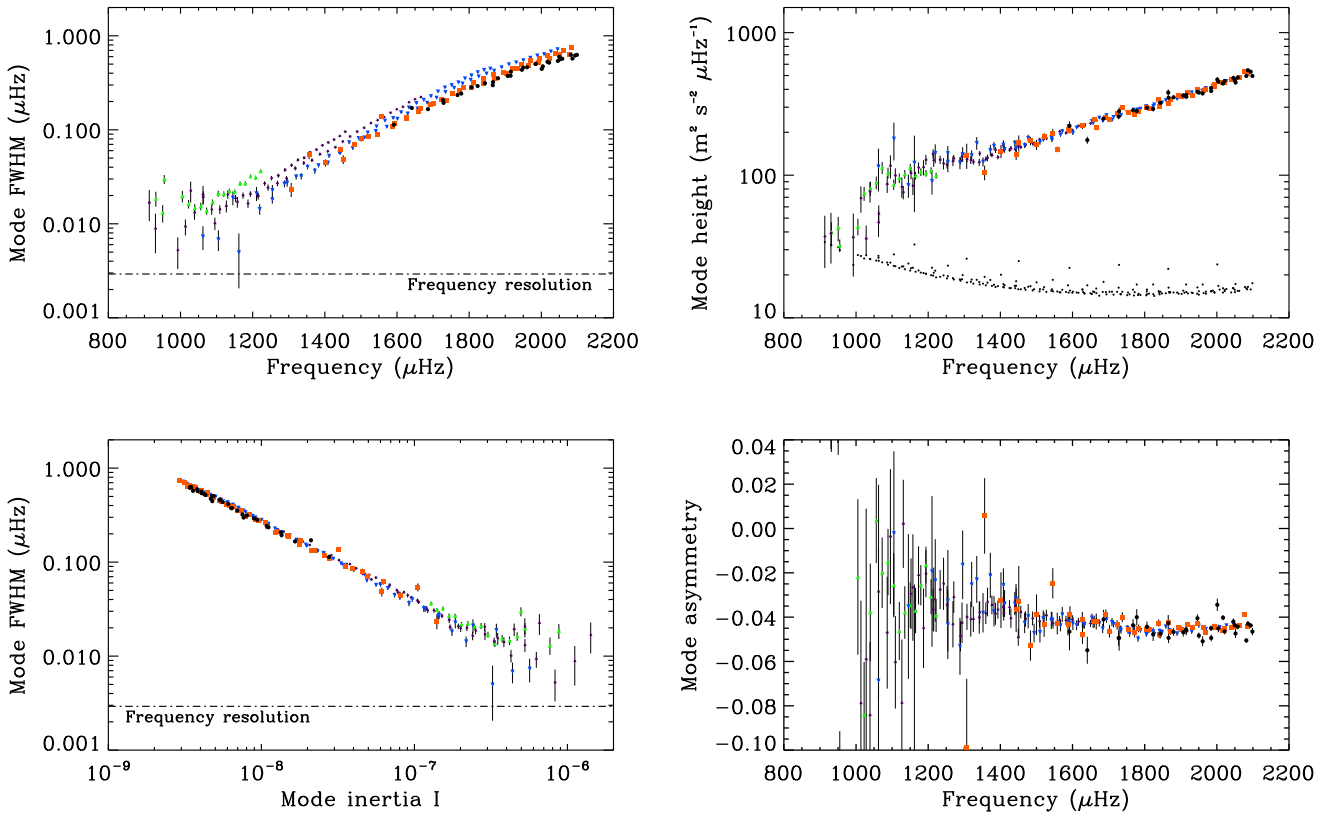


FIG. 13.— Mode FWHM $\Gamma_{n,l}$ (upper-left panel) and height $H_{n,l}$ (upper-right panel) as a function of frequency for the low-frequency p modes $l \leq 35$ extracted from 3960 days of GONG data with the m -averaged spectrum technique. The different colors and symbols correspond to selected ranges of radial orders n : green triangles, modes with $n = 1, 2$; purple stars, $n = 3, 4$; blue upside-down triangles, $n = 5, 6$; orange squares, $n = 7, 8$; and black dots, $n \geq 9$. The frequency resolution of the analyzed dataset is indicated on the upper-left hand-side plot, and the fitted background level ($B_{n,l}$) is also represented on the upper-right hand-side plot (small black dots). Lower-left panel - Same as above, but for the mode FWHM $\Gamma_{n,l}$ as a function of the mode inertia (I), calculated from Christensen-Dalsgaard model S (Christensen-Dalsgaard et al. 1996). The frequency resolution is also indicated. Lower-right panel - Same as above, but for the mode asymmetry ($\alpha_{n,l}$).

since they sample a large range of penetration depths.

Moreover, because these modes have lower upper turning

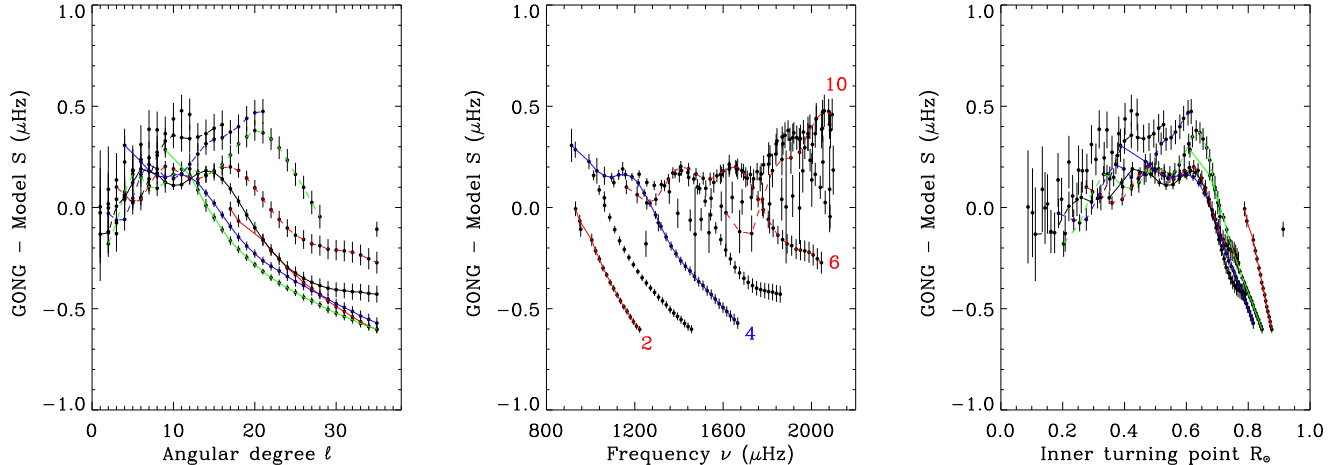


FIG. 14.— Differences between the low-frequency central frequencies observed in 3960 days of GONG data ($1 \leq l \leq 35$) using the m -averaged spectrum technique and the associated theoretical values, taken from Christensen-Dalsgaard’s model S (Christensen-Dalsgaard et al. 1996), as a function of the angular degree (*left panel*), of the frequency (*middle panel*), and of the inner turning point (*right panel*). The uncertainties multiplied by a factor 20 are also presented. Ridges of same radial orders are connected, and for guidance some of the corresponding radial orders are indicated in the middle panel.

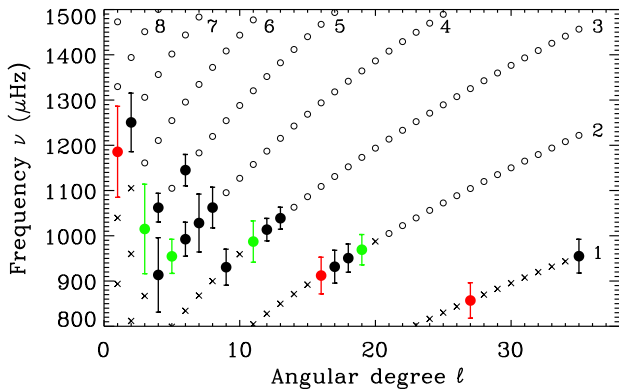


FIG. 15.— $l - \nu$ diagram of the new low-frequency p modes observed in spatially-resolved data in the range of angular degrees $1 \leq l \leq 35$ (black dots: observed in the 3960-day GONG dataset / green dots: observed in the 2088-day GONG dataset / red dots: modes observed in the 2088-day MDI dataset). The corresponding frequency uncertainties were multiplied by 2×10^4 . The already known modes are represented by the open circles, and the predicted modes by the crosses. The ridges of same radial order are also indicated from $n = 1$ to $n = 8$.

points in the outer part of the Sun, they are less sensitive to the turbulence and magnetic fields in the outer layers, which should make them extremely valuable for the study of the physical processes responsible for the oscillation excitation and damping by the turbulent convection.

We would like to recall that Schou (1992)’s peak-finding approach consists in fitting the individual- m spectra simultaneously by using a model in which the shift coefficients are introduced, while in the present technique, the best shifts are determined first, based on the calculation of figure-of-merits (Sec. 3.1 and Appendix A), and then the rotation-corrected, m -averaged spectrum is fitted (Sec. 3.2).

TABLE 2
SET OF THE NEW LOW-FREQUENCY SOLAR P MODES OBSERVED IN THE GONG AND MDI DATASETS WITH THE m -AVERAGED SPECTRUM TECHNIQUE IN THE RANGE $1 \leq l \leq 35$.

l	n	Frequency (μHz)	a_1 -coefficient (nHz)
1	7	1185.599 ± 0.005	431.491 ± 6.161
2	7	1250.555 ± 0.003	428.263 ± 2.286
3	5	1015.046 ± 0.005	430.154 ± 2.471
4	4	913.477 ± 0.004	420.055 ± 1.594
4	5	1062.140 ± 0.002	429.275 ± 0.614
5	4	954.560 ± 0.002	430.712 ± 0.596
6	4	992.412 ± 0.002	431.906 ± 0.502
6	5	1145.074 ± 0.002	432.019 ± 0.464
7	4	1028.156 ± 0.003	430.292 ± 0.740
8	4	1062.338 ± 0.002	434.087 ± 0.459
9	3	930.540 ± 0.002	430.045 ± 0.363
11	3	987.206 ± 0.002	436.639 ± 0.344
12	3	1013.572 ± 0.001	435.158 ± 0.172
13	3	1038.795 ± 0.001	435.900 ± 0.156
16	2	912.080 ± 0.002	436.331 ± 0.213
17	2	931.609 ± 0.002	435.855 ± 0.180
18	2	950.625 ± 0.002	436.627 ± 0.146
19	2	969.222 ± 0.002	438.012 ± 0.149
27	1	856.964 ± 0.002	437.741 ± 0.123
35	1	954.940 ± 0.002	440.118 ± 0.091

The development of the m -averaged spectrum technique towards both higher frequencies and larger angular degrees is one of the next step to be addressed, as also the analysis of shorter datasets, such as the canonical 108- and 72-day time series.

This work utilizes data obtained by the Global Oscillation Network Group (GONG) program, managed by the National Solar Observatory, which is operated by AURA, Inc. under a cooperative agreement with the National Science Foundation. The data were acquired by

instruments operated by the Big Bear Solar Observatory, High Altitude Observatory, Learmonth Solar Observatory, Udaipur Solar Observatory, Instituto de Astrofísica de Canarias, and Cerro Tololo Interamerican Observatory. The GOLF and MDI instruments onboard SOHO are cooperative efforts to whom we are indebted. SOHO is a project of international collaboration between ESA and NASA. BiSON is funded by the Science Technology and Facilities Council (STFC). We thank the members of the BiSON team, and colleagues at the host institutes at each of the BiSON sites. The authors are particularly grateful to S. G. Korzennik for providing us with the

2088-day MDI dataset, and to J. Schou for the MDI leakage matrix. The authors thank R. A. García, P. Boumier, and W. J. Chaplin for providing estimates of GOLF and BiSON mode frequencies and a_1 rotational splittings observed in decade-long time series. The authors also thank S. J. Jiménez-Reyes and J. Schou for their useful comments on the manuscript, and S. G. Korzennik for helpful discussions during the various stages of this work. D. S. acknowledges the support of the NASA SEC GIP grant NAG5-11703. This work has been partially funded by the grant PNAyA2007-62650 of the Spanish National Research Plan.

APPENDIX

FIGURES-OF-MERIT AND DETERMINATION OF THE A -COEFFICIENTS

The best estimates of the splitting a -coefficients are obtained by maximizing the likelihood of the m -averaged spectrum (see Sec. 3.1) through the calculation of a figure-of-merit (FOM). However, as shown on Fig. 1, other criteria to define a FOM can be used, such as the narrowest peak (i.e., the minimum mode linewidth), or the minimum entropy of the resulting m -averaged spectrum. In order to compare the actual mode parameters and associated uncertainties obtained with two different definitions of the FOM, we applied the m -averaged spectrum technique to the 3960-day GONG dataset by using both the maximum likelihood and the narrowest peak in the m -averaged spectrum as FOM. Figure A16 shows the corresponding central frequencies, and the odd a_1 , a_3 , and a_5 splitting coefficients of the common, measured low-frequency p modes. The associated formal uncertainties are also represented. The two FOMs return consistent mode parameters within the error bars, the difference between the two being within the 3σ limit for all of the mode parameters.

IMPACT OF THE FITTING MODEL USED

As a test of the dependence of the measured frequencies on the fitting model used to describe the m -averaged spectra, we fitted the m -averaged spectra using three different models: an asymmetric Lorentzian profile (Eq. 2) including the closest $\delta m \pm 2$ spatial leaks (hereafter called $A2$ and used as the reference model); a symmetric Lorentzian profile including the $\delta m \pm 2$ spatial leaks (hereafter $S2$); and an asymmetric Lorentzian profile (Eq. 2) but omitting the neighbouring $\delta m \pm 2$ spatial leaks (hereafter A). Figure B17 shows the differences as a function of frequency in the 2088-day GONG low frequencies estimated using the m -averaged technique between $S2$ and $A2$ (red dots), and between A and $A2$ (black plus signs), in both cases $A2$ being the reference model. Ignoring the peak asymmetry in the fitting model leads to a systematic underestimation of the mode frequency as the frequency increases, the effect becoming particularly large above $\approx 1400 \mu\text{Hz}$ (red dots). The differences become much larger than 3σ , for example, at $\approx 1800 \mu\text{Hz}$, the fitted frequencies between $S2$ and $A2$ are about 20σ apart. These results obtained for modes below $2000 \mu\text{Hz}$ confirm previous observations, e.g. Thiery et al. (2000) who analyzed low-degree modes above $2000 \mu\text{Hz}$ in 805 days of GOLF data.

On the other hand, omitting the spatial leaks has no effect below $\approx 1600 \mu\text{Hz}$, as they become well separated from the main peak because the corresponding mode linewidths are much smaller than their frequency separation. As the frequency increases, the mode linewidths increase, and ignoring the spatial leaks in the fitting model of the m -averaged spectrum between about 1600 and $2000 \mu\text{Hz}$ leads to an underestimation of the target mode frequency, the maximum difference occurring around $1800 \mu\text{Hz}$. The frequency separation between the target mode and the m -leaks then becomes comparable to their linewidths and the lines blend together in the m -averaged spectrum. Above $2000 \mu\text{Hz}$, this underestimation seems to vanish. Indeed, at that frequency range, the mode linewidths are much larger than the frequency separation, and the first spatial leaks (at least) are totally blended into the target mode in the m -averaged spectrum, having a much lower impact on the frequency determination. However, the effect of ignoring the peak asymmetry is much larger than that from ignoring the m leaks even in the frequency range where the m leaks have the strongest impact. For instance, at $1800 \mu\text{Hz}$, the effect on the frequency underestimation by ignoring the mode asymmetry is about seven times larger than by ignoring the m leaks.

As an example of the other mode parameters, the right panel of Fig. B17 shows the differences in the extracted mode linewidths between the different fitting models. The color code is the same as for the differences in frequency represented on the left panel of Fig. B17. Ignoring the presence of the m leaks in the fitting model leads to a 35% overestimation at most of the extracted linewidths in the low-frequency range showing a maximum mismatch around $1900 \mu\text{Hz}$. Interestingly, if the m leaks are omitted, the linewidths are underestimated below $\approx 1600 \mu\text{Hz}$, showing a maximum 10% underestimation around $1500 \mu\text{Hz}$. On the other hand, ignoring the peak asymmetry has a very small influence on the fitted linewidths in the low-frequency range. However, above $\approx 1800 \mu\text{Hz}$, the linewidths extracted using an asymmetric profile are systematically larger than the ones returned using a symmetric profile.

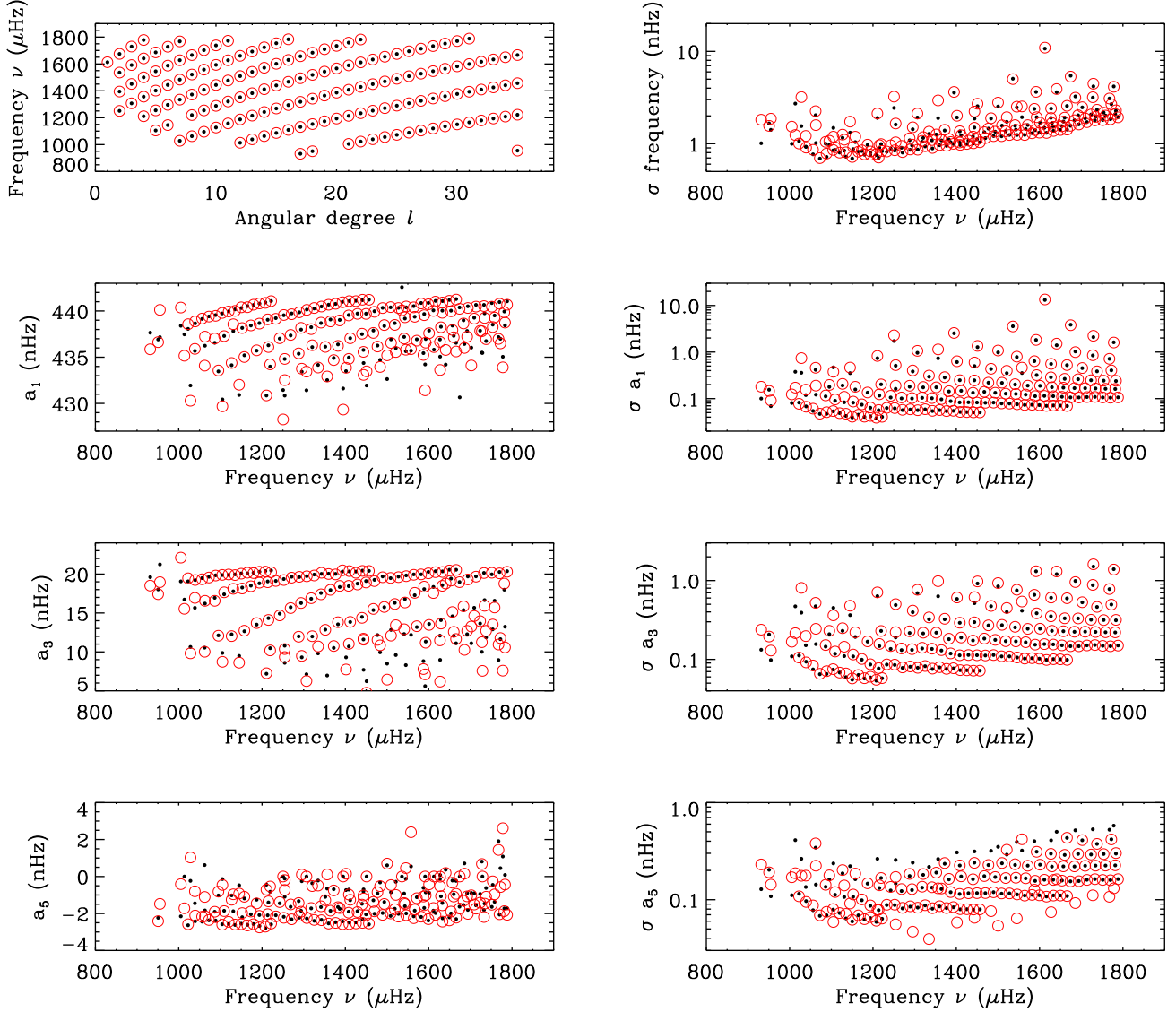


FIG. A16.— *Left column* - l - ν diagram, and a_1 , a_3 , and a_5 splitting coefficients for the common, low-frequency p modes measured in the 3960-day GONG dataset obtained by using the maximum likelihood (circles) and the narrowest peak (dots) of the m -averaged spectrum as figure-of-merit to determine the best estimates of the a -coefficients (see Sec. 3.1 and Fig. 1). *Right column* - Associated 1σ formal uncertainties (nHz) of the mode central frequencies, and of the a_1 , a_3 , and a_5 splitting coefficients.

DERIVATION OF ERRORS FROM AN M -AVERAGED SPECTRUM

The derivation of the errors of the mode central frequencies and of the a -coefficients measured from the m -averaged spectrum technique is detailed here.

Approximation of the statistics of the m -averaged spectrum

The m -averaged spectrum is obtained from the summation of $2l + 1$ spectra assumed to be with χ^2 with 2 d.o.f statistics each having a *different* mean or signal-to-noise ratio. All of the individual m -spectra are independent from each other. The solar background noise is assumed to depend on m with a polynomial with only even terms (0, 2, etc...). The amplitude of the modes is assumed to depend on m with a different polynomial also with even terms (0, 2, etc...). In a first step, the a_i -coefficients are calculated to maximize the likelihood of the resulting m -averaged spectrum.

Using Appourchaux (2003), we can derive an approximation of the statistics of the summation of the $2l + 1$ spectra. The statistics of the m -averaged spectrum \mathcal{S} can be approximated by a Gamma law given by:

$$p(\mathcal{S}) = \frac{\lambda^{\nu_1}}{\Gamma(\nu_1)} \mathcal{S}^{\nu_1-1} e^{-\lambda \mathcal{S}} \quad (\text{C1})$$

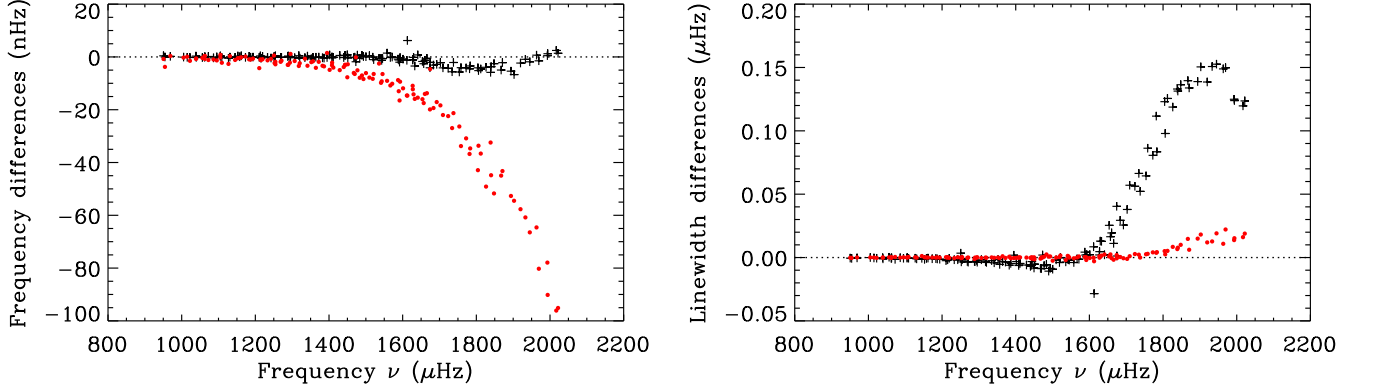


FIG. B17.— Effect of asymmetry and spatial leaks on the fits. *Left panel* - Differences (in nHz) in the 2088-day GONG frequencies estimated using different models to describe the m -averaged spectrum: symmetric Lorentzian profile minus asymmetric Lorentzian profile (Eq. 2) both models including the first $\delta m \pm 2$ spatial leaks, i.e., $S2$ minus $A2$ (red dots); both asymmetric Lorentzian profiles (Eq. 2) but ignoring the closest $\delta m \pm 2$ spatial leaks for one of them, i.e., A minus $A2$ (black plus signs). The differences in the extracted frequencies using different fitting profiles are represented as a function of frequency. *Right panel* - Same as the left panel, but for the differences (in μHz) in the fitted mode linewidths.

The mean and σ are given by:

$$E[S] = \frac{\nu_1}{\lambda} \text{ and } \sigma^2 = \frac{\nu_1}{\lambda^2} \quad (\text{C2})$$

λ and ν_1 are then derived from the mean and σ as:

$$\lambda = \frac{E[S]}{\sigma^2} \text{ and } \nu_1 = \frac{E[S]^2}{\sigma^2} \quad (\text{C3})$$

In our case, the mean $E[S]$ and σ are given by:

$$E[S] = \sum_{m=-l}^{m=l} f_m \text{ and } \sigma = \sqrt{\sum_{m=-l}^{m=l} f_m^2} \quad (\text{C4})$$

where f_m is the power spectrum for azimuthal order m which can be expressed as:

$$f_m(\nu, \nu_0, a_i) = B_m(\nu) + A_m(\nu, \nu_0, a_i) \quad (\text{C5})$$

where ν is the frequency, ν_0 the central frequency, a_i are the usual Ritzwoller-Lavely coefficients, B_m is the background noise, A_m is the profile of the mode (the linewidth and amplitude have been omitted for simplifying the notation). We can write the noise as:

$$B_m(\nu) = \mathcal{B}(\nu)(1 + g_B(m)) \quad (\text{C6})$$

where the m -dependence is assumed to be independent of frequency. $g_B(m)$ is such that:

$$E[B_m(\nu)] = (2l + 1)\mathcal{B}(\nu) \quad (\text{C7})$$

If the correction of the a_i has been done properly, to the first order the m -averaged spectrum is independent of the a_i . We can write the mode amplitude as:

$$A_m(\nu) = \mathcal{A}(\nu)(1 + h_A(m)) \quad (\text{C8})$$

where the m -dependence is assumed to be independent of frequency. $h_A(m)$ is such that

$$E[A_m(\nu)] = (2l + 1)\mathcal{A}(\nu) \quad (\text{C9})$$

Then we find:

$$E[S] = (2l + 1)(\mathcal{A}(\nu) + \mathcal{B}(\nu)) \quad (\text{C10})$$

and

$$\sigma^2 = (2l + 1) [\mathcal{A}^2(\nu)(1 + \alpha) + \mathcal{B}^2(\nu)(1 + \beta) + 2\mathcal{A}\mathcal{B}(\nu)(1 + \rho)] \quad (\text{C11})$$

with

$$\alpha = \frac{1}{2l + 1} \sum_{m=-l}^{m=l} h_A^2(m) \quad (\text{C12})$$

and

$$\beta = \frac{1}{2l+1} \sum_{m=-l}^{m=l} g_B^2(m) \quad (\text{C13})$$

$$\rho = \frac{1}{2l+1} \sum_{m=-l}^{m=l} h_A(m)g_B(m) \quad (\text{C14})$$

we finally get for λ and ν_1 the following:

$$\lambda = \frac{\mathcal{A}(\nu) + \mathcal{B}(\nu)}{\mathcal{A}^2(\nu)(1+\alpha) + \mathcal{B}^2(\nu)(1+\beta) + 2\mathcal{A}\mathcal{B}(\nu)(1+\rho)} \quad (\text{C15})$$

and ν_1 as:

$$\nu_1 = \frac{(2l+1)(\mathcal{A}(\nu) + \mathcal{B}(\nu))^2}{\mathcal{A}^2(\nu)(1+\alpha) + \mathcal{B}^2(\nu)(1+\beta) + 2\mathcal{A}(\nu)\mathcal{B}(\nu)(1+\rho)} \quad (\text{C16})$$

After simplification we get:

$$\lambda = \frac{\mathcal{A}(\nu) + \mathcal{B}(\nu)}{(\mathcal{A}(\nu) + \mathcal{B}(\nu))^2 + \alpha\mathcal{A}^2(\nu) + \beta\mathcal{B}^2(\nu) + 2\rho\mathcal{A}(\nu)\mathcal{B}(\nu)} \quad (\text{C17})$$

and

$$\nu_1 = \frac{(2l+1)(\mathcal{A}(\nu) + \mathcal{B}(\nu))^2}{(\mathcal{A}(\nu) + \mathcal{B}(\nu))^2 + \alpha\mathcal{A}^2(\nu) + \beta\mathcal{B}^2(\nu) + 2\rho\mathcal{A}(\nu)\mathcal{B}(\nu)} \quad (\text{C18})$$

Using the dependence observed in the GONG data, we have $\alpha \approx 0.17$, $\beta \approx 0.035$ and $\rho \approx 0.075$. They are sufficiently small such that we have:

$$\lambda \approx \frac{1}{\mathcal{A}(\nu) + \mathcal{B}(\nu)} \quad (\text{C19})$$

and

$$\nu_1 \approx (2l+1) \quad (\text{C20})$$

then we find the following statistics for the m -averaged spectrum:

$$p(\mathcal{S}) = \frac{1}{\Gamma(2l+1)} \frac{\mathcal{S}^{2l}}{(\mathcal{A}(\nu) + \mathcal{B}(\nu))^{2l+1}} e^{-\frac{\mathcal{S}}{\mathcal{A}(\nu) + \mathcal{B}(\nu)}} \quad (\text{C21})$$

After a change of variable $u = \mathcal{S}/(2l+1)$, we have

$$p(u) \propto \frac{1}{(\mathcal{A}(\nu) + \mathcal{B}(\nu))^{2l+1}} e^{-\frac{(2l+1)u}{\mathcal{A}(\nu) + \mathcal{B}(\nu)}} \quad (\text{C22})$$

When we use MLE, we minimize the following

$$\mathcal{L}(\nu, \nu_0, a_i) = -\ln p(u) = -(2l+1) \left[\ln(\mathcal{A}(\nu) + \mathcal{B}(\nu)) + \frac{u}{(\mathcal{A}(\nu) + \mathcal{B}(\nu))} \right] + \dots \quad (\text{C23})$$

which shows that using the MLE applied to a χ^2 with 2 d.o.f as prescribed by Appourchaux (2003) is in the case of the m -averaged spectrum a good approximation. It is not an approximation when averaging several power spectra of identical mean (or variance), i.e. when $\alpha = \beta = \rho = 0$. Note that what we minimize is the sum over a range of frequency that can be approximated as:

$$L(\nu_0, a_i) = \int \mathcal{L}(\nu, \nu_0, a_i) d\nu \quad (\text{C24})$$

Error bars on the central frequencies

Error bars for frequency are derived from the inverse of the Hessian (second derivative of L) as:

$$\sigma_{\nu_0}^{-2} = \frac{\partial^2 L}{\partial \nu_0^2}. \quad (\text{C25})$$

Toutain & Appourchaux (1994) showed that we could express the error bars as a function of the mode profile \mathcal{P} ($=\mathcal{A} + \mathcal{B}$) as:

$$\sigma_{\nu_0}^{-2} = T(2l+1) \int \frac{1}{\mathcal{P}^2(\nu)} \left(\frac{\partial \mathcal{P}}{\partial \nu_0} \right)^2 d\nu, \quad (\text{C26})$$

where T is the observation time. The $2l+1$ factor is due to the fact that the likelihood is $2l+1$ times larger than the likelihood of Toutain & Appourchaux (1994) (cf Eq. 23). Eq. (C26) shows that the error bars on the frequencies in the m -averaged spectrum will be $\sqrt{2l+1}$ smaller than for the *mean* of the individual modes. In deriving Eq. (C26), we assumed that $\langle u \rangle = \mathcal{P}$. This is an approximation good enough for getting the error bars on the frequency but not on the a_i .

Error bars on the a-coefficients

The error bars on the a -coefficients are derived from the inverse of the Hessian (second derivative of L) as:

$$h_{ij} = \frac{\partial^2 L}{\partial a_i \partial a_j}. \quad (\text{C27})$$

As shown by Toutain & Appourchaux (1994), these coefficients can be related to the mode profile as using Eq. (C23):

$$h_{ij} = T \sum_m \int \frac{1}{\mathcal{P}_m^2(\nu)} \frac{\partial \mathcal{P}_m}{\partial a_i} \frac{\partial \mathcal{P}_m}{\partial a_j} d\nu, \quad (\text{C28})$$

using the following property:

$$\frac{\partial \mathcal{P}_m}{\partial a_i} = \frac{\partial \mathcal{P}_m}{\partial \nu_0} l P_{l,m}^i(m/l), \quad (\text{C29})$$

where $P_{l,m}^i(m/l)$ are the Ritzwoller-Lavely polynomials. And we finally get:

$$h_{ij} = T l^2 \sum_m P_{l,m}^i(m/l) P_{l,m}^j(m/l) \int \frac{1}{\mathcal{P}_m^2(\nu)} \left(\frac{\partial \mathcal{P}_m}{\partial \nu_0} \right)^2 d\nu. \quad (\text{C30})$$

We recognize the error bars of the frequency for the m spectrum depending on the inverse of the signal-to-noise ratio β_m (as in Libbrecht 1992):

$$\sigma_m^{-2} = T \int \frac{1}{\mathcal{P}_m^2(\nu)} \left(\frac{\partial \mathcal{P}_m}{\partial \nu_0} \right)^2 d\nu. \quad (\text{C31})$$

Finally Eq. (C30) can be written as:

$$h_{ij} = l^2 \left(\sum_m P_{l,m}^i(m/l) P_{l,m}^j(m/l) \right) \sigma_m^{-2}. \quad (\text{C32})$$

The errors σ_{a_i} scale like $l^{-\frac{3}{2}}$ (as in Veitzer et al. 1993). If the SNR is the same for all m , then we have $\sigma_{\nu_0}^{-2} = (2l+1)\sigma_m^{-2}$. Thus, by simply using the orthogonality property of the $P_{l,m}^i$ polynomials, and as given in Sec. 3.3 (Eq. 5), we obtain the following expression to calculate the error bars of the a -coefficients in the m -averaged spectrum:

$$\sigma_{a_i}^{-2} = \frac{l^2}{2l+1} \left(\sum_m [P_{l,m}^i(m/l)]^2 \right) \sigma_{\nu_0}^{-2}. \quad (\text{C33})$$

All terms off of the diagonal are zero. Of course, when the SNR varies with m , the off-diagonal terms are non-zero and correlations appear.

REFERENCES

- Anderson, E. R., Duvall, T. L., Jr., & Jefferies, S. M. 1990, *ApJ*, 364, 699
- Appourchaux, T. 2003, *A&A*, 412, 903
- Appourchaux, T., & Chaplin, W. J. 2007, *A&A*, 469, 1151
- Appourchaux, T., et al. 2000, *ApJ*, 538, 401
- Broomhall, A. M., Chaplin, W. J., Elsworth, Y., & Appourchaux, T. 2007, *MNRAS*, 379, 2
- Brown, T. M. 1985, *Nature*, 317, 591
- Chaplin, W. J., Appourchaux, T., Elsworth, Y., Isaak, G. R., Miller, B. A., & New, R. 2004, *A&A*, 424, 713
- Chaplin, W. J., Elsworth, Y., Isaak, G. R., Marchenkov, K. I., Miller, B. A., New, R., Pinter, B., & Appourchaux, T. 2002, *MNRAS*, 336, 979
- Chaplin, W. J., et al. 1996, *Sol. Phys.*, 168, 1
- Christensen-Dalsgaard, J., et al. 1996, *Science*, 272, 1286
- Fierry Fraillon, D., Gelly, B., Schmider, F. X., Hill, F., Fossat, E., & Pantel, A. 1998, *A&A*, 333, 362
- Gabriel, A. H., et al. 1995, *Sol. Phys.*, 162, 61
- Harvey, J. W., et al. 1996, *Science*, 272, 1284
- Hill, F., & Howe, R. 1998, *Structure and Dynamics of the Interior of the Sun and Sun-like Stars*, 418, 225
- Larson, T. P., & Schou, J. 2008, in L. Gizon (ed.), *Helioseismology, Asteroseismology, and MHD Connections*, *Journal of Physics* (in press)
- Libbrecht, K. G. 1992, *ApJ*, 387, 712
- Nigam, R., & Kosovichev, A. G. 1998, *ApJ*, 505, L51
- Korzennik, S. G. 2008, *Astron. Nachr.*, 329, 453
- Korzennik, S. G. 2005, *ApJ*, 626, 585
- Ritzwoller, M. H., & Lavely, E. M. 1991, *ApJ*, 369, 557
- Salabert, D., & Jiménez-Reyes, S. J. 2006, *ApJ*, 650, 451
- Scherrer, P. H., et al. 1995, *Sol. Phys.*, 162, 129
- Schou, J. 2004, in *SOHO 14/GONG 2004 Workshop, Helio- and Asteroseismology: Towards a Golden Future*, ed. D. Danesy (ESA SP-559), 134
- Schou, J. 2002, in *SOHO 11 Symposium, From Solar Min to Max: Half a Solar Cycle with SOHO*, ed. A. Wilson (ESA SP-508), 99
- Schou, J. 1999, *ApJ*, 523, L181
- Schou, J. 1998, in *SOHO 6/GONG 98 Workshop, Structure and Dynamics of the Interior of the Sun and Sun-like Stars* (ESA SP-418), 341
- Schou, J. 1992, Ph.D. Thesis, Aarhus Univ.
- Shannon, C. E. 1948, *Bell System Technical Journal*, vol. 27, pp. 379-423, 623-656
- Thiery, S., et al. 2000, *A&A*, 355, 743
- Toutain, T., & Appourchaux, T. 1994, *A&A*, 289, 649
- Veitzer, S. A., Tomczyk, S., & Schou, J. 1993, *GONG 1992. Seismic Investigation of the Sun and Stars*, 42, 465