

Catching dual AGN activity and kiloparsec-scale outflows in IRAS 20210+1121

F. G. Saturni^{1,2}, G. Vietri^{1,3}, E. Piconcelli¹, C. Vignali^{4,5}, M. Bischetti⁶, A. Bongiorno¹, S. Cazzoli⁷, C. Feruglio⁶,
F. Fiore⁶, B. Husemann⁸, and C. Ramos Almeida^{9,10}

¹ INAF – Osservatorio Astronomico di Roma, Via Frascati 33, I-00078 Monte Porzio Catone (RM), Italy.
e-mail: francesco.saturni@inaf.it

² ASI – Space Science Data Center, Via del Politecnico snc, I-00133 Roma, Italy.

³ INAF – Istituto di Astrofisica Spaziale e Fisica Cosmica di Milano, Via A. Corti 12, I-20133 Milano, Italy.

⁴ Università di Bologna, Dip. di Fisica e Astronomia “A. Righi”, Via P. Gobetti 93/2, I-40129 Bologna, Italy.

⁵ INAF – Osservatorio di Astrofisica e Scienza dello Spazio di Bologna, Via P. Gobetti 93/3, I-40129 Bologna, Italy.

⁶ INAF – Osservatorio Astronomico di Trieste, Via G. B. Tiepolo 11, I-34143 Trieste, Italy.

⁷ CSIC – Instituto de Astrofísica de Andalucía, Dep.to de Astronomía Extragaláctica, Glorieta de la Astronomía s/n, E-18008 Granada, Spain.

⁸ Max-Planck Institut für Astronomie, Königstuhl 17, D-69117 Heidelberg, Germany.

⁹ Instituto de Astrofísica de Canarias, C/ Vía Láctea s/n, E-38205 La Laguna (Tenerife), Spain.

¹⁰ Universidad de La Laguna, Dep.to de Astrofísica, Av.da Astrofísico F. Sánchez s/n, E-38206 La Laguna (Tenerife), Spain.

Received XXX / Accepted YYY

ABSTRACT

The most accepted scenario for the evolution of massive galaxies across cosmic time predicts a regulation based on the interplay between AGN feedback, which injects large amounts of energy in the host environment, and galaxy mergers, being able to trigger massive star formation events and accretion onto the supermassive black holes. Interacting systems hosting AGN are useful laboratories to get key insights into both phenomena. In this context, we present the analysis of the optical spectral properties of IRAS 20210+1121 (I20210), a merging system at $z = 0.056$. According to X-ray data, this object comprises two interacting galaxies, each hosting an obscured AGN. The optical spectra confirm the presence of AGN features in both galaxies. In particular, we are able to provide a Seyfert classification for I20210 North. The spectrum of I20210 South shows broad blueshifted components associated with the most intense emission lines that indicate the presence of an ionized outflow, for which we derive a maximum velocity of ~ 2000 km s⁻¹, an extension of ~ 2 kpc and a mass rate of $\sim 0.6 M_{\odot}$ yr⁻¹. We also report the existence of an ionized nebular component with $v \sim 1000$ km s⁻¹ at ~ 6.5 kpc Southwards of I20210 South, that can be interpreted as disrupted gas ejected from the host galaxy by the action of the outflow. I20210 therefore exhibits a double obscured AGN, with one of them showing evidence of ongoing events for AGN-powered outflows. Future spatially-resolved spectroscopy will allow to accurately map on the gas kinematics in this AGN pair and evaluate the impact of the outflow on both the interstellar medium and galaxy environment.

Key words. galaxies: active – galaxies: groups: general – galaxies: groups: individual: IRAS 20210+1121 – galaxies: Seyfert – quasars: emission lines – quasars: supermassive black holes

1. Introduction

The past history of formation and evolution of present-day massive galaxies is a key point to assess on the route to a full comprehension of the functioning of the Universe. In this context, the study of processes operating at galaxy-wide scales such as the presence of active galactic nuclei (AGN; e.g., Lynden-Bell 1969) or events of galaxy merging (e.g., Hernquist 1989) is crucial to understand the mechanisms that are able to boot, maintain, enhance and/or quench star formation in galaxies, thus shaping the whole environment.

It is now widely accepted that AGN activity and galaxy mergers are among the most effective phenomena that regulate star formation in massive galaxies at nearly all redshifts: the first through injection in their surrounding environment of large amounts of energy that are able to originate powerful gas winds (e.g., Di Matteo et al. 2005; Cattaneo et al. 2009; Fabian 2012), the second by triggering massive star formation and starburst events in molecular gas-rich clouds (e.g., Sanders & Mirabel

1996). Both simulations of the evolutionary history of the Universe in the framework of the Λ -CDM model (e.g., Davis et al. 1985; Springel et al. 2005; Croton et al. 2006) and observations of their action in shaping galactic environments also simultaneously (e.g., Sanders et al. 1988; Kormendy & Ho 2013; Ellison et al. 2019) have confirmed the major role that such processes play in galaxy formation and evolution.

Within this general picture, however, several quantities related to how exactly AGN energetics and mergers directly impact the star formation history of galaxies are still missing. It is for example unknown if radiation-powered gas outflows are ubiquitous to all AGN (e.g., Elvis 2000) or they affect only a fraction of the AGN lifetime (e.g., Farrah et al. 2007), and what is their effectiveness in altering the physical and dynamical status of gas reservoirs at several spatial scales (e.g., Scannapieco & Oh 2004; Cicone et al. 2018). In addition, the relative dominance of a process onto the other for moving large gas masses and triggering or quenching star formation has been found to be dependent on the details of the AGN emission mode, the galaxy surrounding

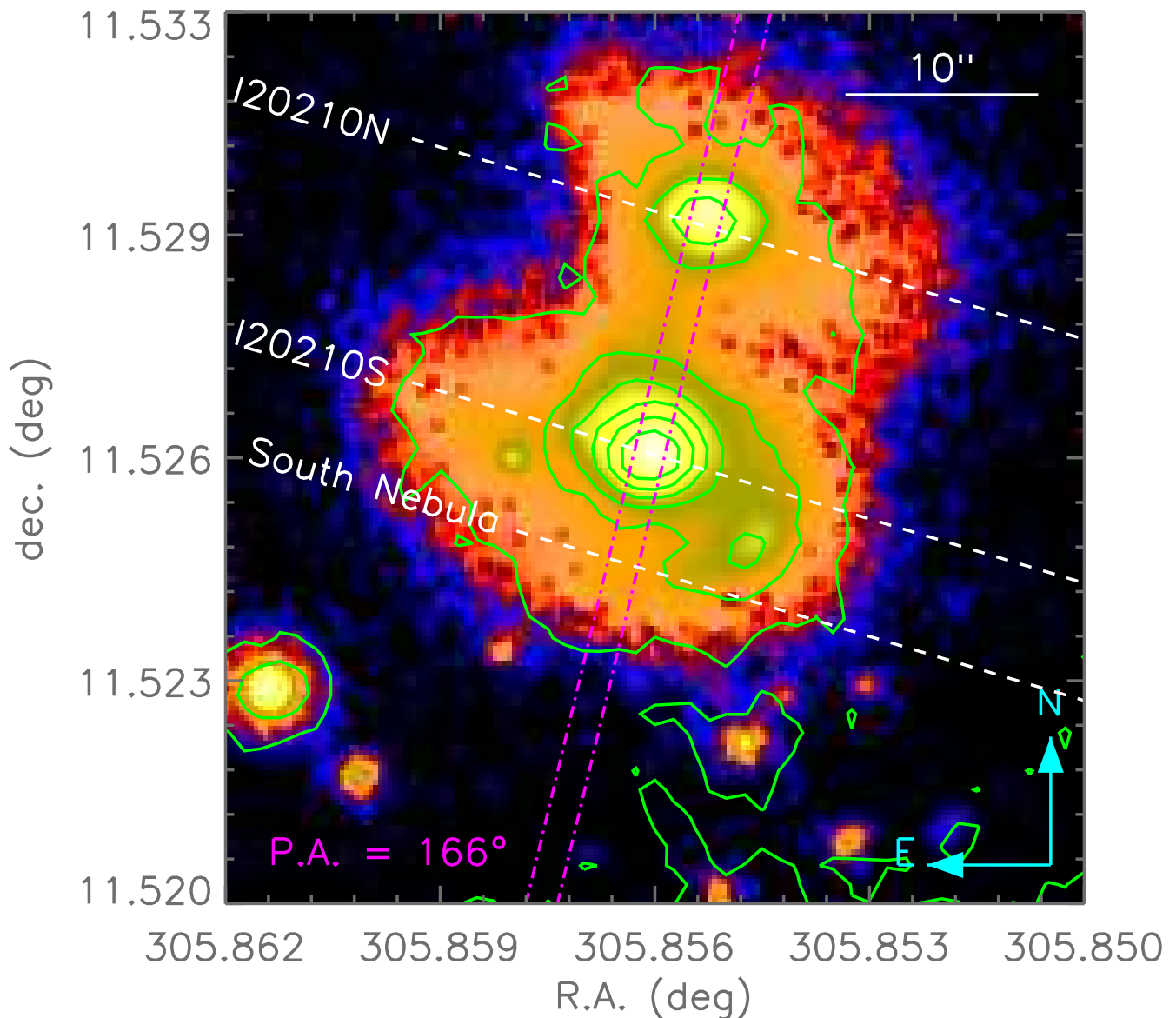


Fig. 1. Image in false colors of the I20210 system, obtained by combining the *grizy* exposures of the Pan-STARRS1 survey (PS1; Chambers & Pan-STARRS Team 2016) centered on the sky coordinates of I20210S ($\alpha_{J2000} = 20\ 23\ 25.4$, $\delta_{J2000} = +11\ 31\ 34.7$). The isophotes of the XMM-Newton Optical Monitor (OM) UVW1 mosaic exposure (*green solid lines*) taken simultaneously to the X-ray data analyzed by Piconcelli et al. (2010) – along with some of the associated CCD count levels – are drawn onto the PS1 image to highlight weak features. The TNG slit direction and position (*magenta dot-dashed lines*) is also indicated along with the positions and directions of the trace centers (*white dashed lines*) identified to extract the 1D spectrum of each object.

environment and its star formation history (e.g., Hopkins et al. 2006; Heckman & Best 2014). Therefore, the study of interacting galactic systems with the presence of multiple AGN (e.g., Veilleux et al. 2002) offers an extremely interesting possibility to understand the properties and links between such competing mechanisms. The most common objects of this kind are the dual AGN, in which an active nucleus is hosted in both members of a pair of interacting galaxies with separation on the scale of $5 \div 20$ kpc (see e.g. De Rosa et al. 2018, and refs. therein).

In this work, we present the results of the optical spectroscopic analysis of the $z = 0.056$ dual AGN IRAS 20210+1121

(I20210 hereafter; Perez et al. 1990, P90 hereafter), which is composed of two interacting galaxies oriented in the N–S direction and separated by $12''.2$ (i.e. ~ 13.3 kpc; Davies et al. 2002; Arribas et al. 2004). Considered at first composed by a Seyfert 2 with asymmetric emission lines (the Southern component) and a normal galaxy (the Northern component), X-ray observations performed with XMM-Newton revealed that this system is actually a merger between two obscured AGN hosts (Piconcelli et al. 2010), in which the Southern member is an ultra-luminous infrared galaxy (ULIRG; e.g., Sanders et al. 1988). Additionally, despite having also spectroscopic data in

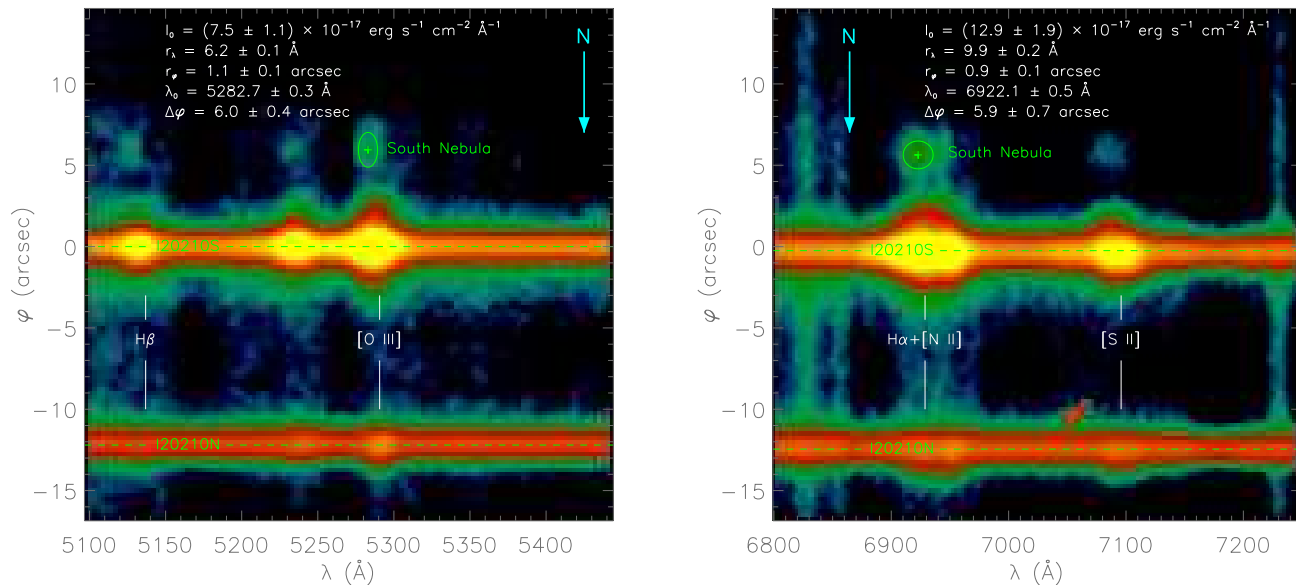


Fig. 2. *Left panel:* H β and [O III] region of the 2D spectrum of the I20210 system. *Right panel:* H α + [N II] and [S II] region. In both panels, the trace centers of the main components are identified for reference (green dashed lines). The elliptical fits to the [O III] and H α + [N II] emissions from the South Nebula are reported (green ellipses) along with the respective best-fit parameters and statistical uncertainties. In the right panel, the vertical features are sky lines, whereas the extrusion close to the [S II] emission of I20210N is a cluster of saturated pixels that is excluded from the IRAF extraction of the 1D spectrum.

the near-infrared (Burston et al. 2001), the optical spectrum of the Northern member was still unobserved, due to its faintness compared to the Southern galaxy (Heisler & Vader 1995). The image of the I20210 system, obtained by combining the *grizy* exposures from the Pan-STARRS1 survey (PS1; Chambers & Pan-STARRS Team 2016), is shown in Fig. 1. This image already gives an idea of the complex structure of the system, displaying a luminous bridge that connects the two galaxies.

The paper is organized as follows: we describe the observation and data-reduction process in Sect. 2; we characterize the extracted spectrum of the Northern galaxy in Sect. 3, and that of the Southern one in Sect. 4; we discuss the relevant physical properties of the structural components of the Southern I20210 member in Sect. 5 and 6; we estimate the supermassive black hole (SMBH) mass of both I20210 members in Sect. 7; finally, we summarize our findings in Sect. 8. For simplicity, we abbreviate the name of the two galaxies as I20210N (Northern member) and I20210S (Southern member) hereafter. Throughout the article, we adopt a Λ -CDM cosmology with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$ and $\Omega_\Lambda = 0.7$.

2. Observations and data reduction

Observations of the I20210 optical spectra were carried out on 2010 August 01 at the *Telescopio Nazionale Galileo* (TNG; Canary Islands, Spain). The spectra were simultaneously obtained with the *B*-band grism (wavelength range $\lambda\lambda 3000 - 8430 \text{ \AA}$, dispersion of 2.52 \AA px^{-1} , $\lambda/\Delta\lambda = 585$, implying a resolution of 9.8 \AA that corresponds to $\sim 510 \text{ km s}^{-1}$) of the DOLoRes instrument (point-spread function PSF $\sim 0''.85$), coupled to the $1''.5$ slit. To this end, the instrument configuration was rotated to a position angle of 166° in order to align the slit along the system axis connecting the two nuclei. The two exposures of 600 s each (total exposure time of 1200 s) were then reduced with standard IRAF

procedures to extract and calibrate the one-dimensional spectra. We show the slit position and orientation (P.A. = 166° East of North) along with the directions of the apertures used to extract the spectra of each object in Fig. 1, superimposed to the PS1 image of the system. The resulting spectra have signal-to-noise ratios $S/N \sim 23.6$ (I20210N) and ~ 33.2 (I20210S), respectively, as computed in line-free continuum regions (Rosales-Ortega et al. 2012).

During the extraction and calibration procedures of the 1D spectra, we found that a spectrum emitted from a third location was visible southwards of I20210S, at a projected distance of $\sim 6''$ (corresponding to $\sim 6.5 \text{ kpc}$ given the distance scale of $1.087 \text{ kpc}''$ at $z = 0.056$) from its trace center. This additional spectrum, already identified by P90 in their low-resolution data as extended emission in the I20210S host galaxy (South Nebula hereafter), is shown in 2D form in Fig. 2: it exhibits the main transitions detected in I20210S (H β $\lambda\lambda 4862$, [O III] $\lambda\lambda 4959, 5007$, H α + [N II] $\lambda\lambda 6548, 6583$ and [S II] $\lambda\lambda 6716, 6731$) detached by $\sim 6''$ from the I20210S nuclear spectrum, extending over $\sim 2''$ (i.e. $\sim 2.3 \text{ kpc}$) in the N–S direction and blueshifted by $\sim 450 \text{ km s}^{-1}$ with respect to the systemic rest frame. We thus extract and calibrate it in the same way as the spectra of the main components.

Since the I20210 system is viewed through the Galactic plane, we deredden the spectra with the Milky Way extinction curve by Pei (1992) and $A_V = 0.6$ (P90). The final spectra obtained in this way are shown in Fig. 3. A visual inspection of the (so-far undetected) I20210N optical spectrum reveals prominent [O II] $\lambda 3727$, [O III] $\lambda\lambda 4959, 5007$, H α + [N II] $\lambda\lambda 6548, 6583$ and [S II] $\lambda\lambda 6716, 6731$ emission lines, as well as the lack of the H β $\lambda 4862$ feature. Such a spectrum shows strong similarities with those of typical Seyfert 2 galaxies, such as NGC 1667 (Ho et al. 1993, 1995; Jones et al. 2009) and Mrk 1018 (Osterbrock 1981).

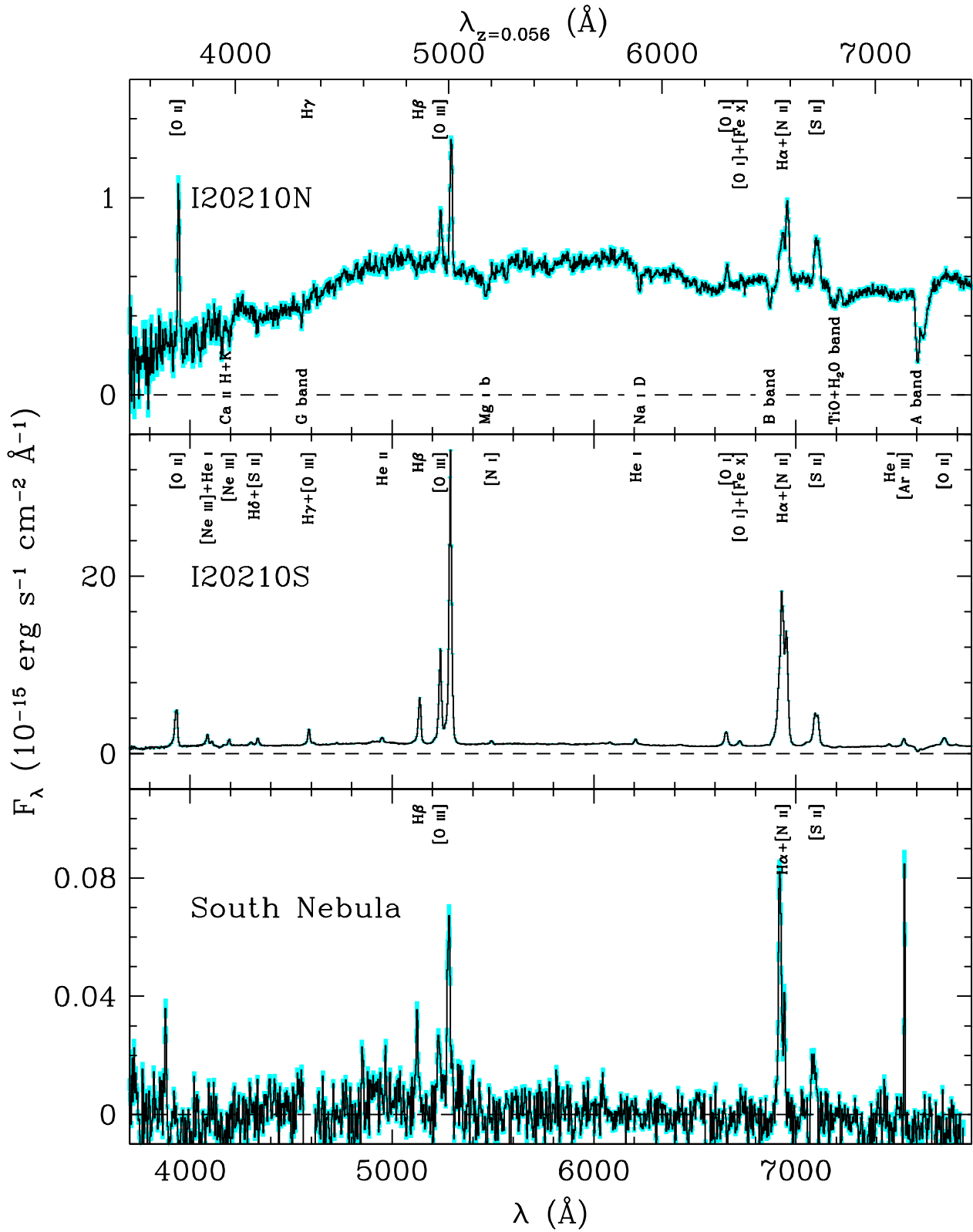


Fig. 3. *Top panel:* optical spectrum of I20210N. *Middle panel:* optical spectrum of I20210S. *Bottom panel:* optical spectrum of the spatially extended South Nebula. In all panels: (i) the detected signal is reported along with its rms uncertainty (cyan bands), (ii) the zero-flux level (dashed line) is indicated, and (iii) the positions of major emission (top) and absorption features (bottom) are labeled accordingly.

3. Characterization of I20210 North

We first proceed to estimate the amount of intrinsic dust extinction in each object. To this end, we decided to measure the reddening $E(B-V)$ of the AGN spectrum through the Balmer decrement $F_{H\alpha}/F_{H\beta}$ (e.g., Miller & Mathews 1972):

$$\frac{F_{H\alpha}}{F_{H\beta}} = \frac{I_{H\alpha}}{I_{H\beta}} \cdot 10^{-0.4E(B-V)(1+R_V)(\kappa_\alpha - \kappa_\beta)} \quad (1)$$

with the intrinsic ratio $I_{H\alpha}/I_{H\beta}$ depending on the physical conditions of the emitting gas only (see e.g. Gaskell & Ferland 1984, and refs. therein), and R_V and κ_λ determined by the adopted extinction model. Since no evidence for narrow lines associated with $H\beta$ is visible bluewards of the [O III] doublet in the I20210N spectrum, we first proceed to model the underlying continuum in order to recover the Balmer emission from the narrow-line region (NLR) of I20210N.

3.1. I20210N continuum and emission-line fitting

We model the I20210N continuum under the assumption of a negligible AGN contribution to the continuum emission. This is justified by the fact that the I20210N central engine ($L_X = 4.7 \times 10^{42} \text{ erg s}^{-1}$) is highly obscured by a column density $N_H \sim 5 \times 10^{23} \text{ cm}^{-2}$ (Picconcelli et al. 2010), and therefore no light from accretion activity is visible. We subtract the stellar continuum with absorption lines from the I20210N galaxy spectrum by using the penalized-pixel fitting public code pPXF (Cappellari & Emsellem 2004; Cappellari 2012, 2017). The spectrum is fitted with a linear combination of stellar spectra templates from the MILES library (Vazdekis et al. 2010), which contains single stellar population synthesis models covering the same wavelength range as the I20210N spectrum with a resolution FWHM = 2.54 Å. This procedure also yields information about the kinematics status of the stellar population in the galaxy through the stellar velocity dispersion σ_v^* .

We rebin the MILES templates ($\lambda/\Delta\lambda \sim 2.5 \text{ Å}$) to match the DOLoRes spectral resolution of $\sim 10 \text{ Å}$. We include low-order additive (4th-degree) and multiplicative (1st-degree) Legendre polynomials to adjust the continuum shape of the templates to the observed spectrum. During the fitting procedure, strong emission features are masked out and the spectra are shifted to the rest frame. The pPXF best-fit model is chosen through χ^2 minimization. To estimate the uncertainty on the velocity dispersion, we produce 10^3 realizations of the I20210N spectrum by adding noise to the pPXF best-fit model; this noise is drawn from a Gaussian distribution with dispersion equal to the rms of the input spectrum. We then iterate the pPXF fitting procedure over such mock spectra, and compute the error associated to σ_v^* as the standard deviation of the parameter posterior distribution. In doing this, we find a best-fit $\sigma_v^* = 390 \pm 50 \text{ km s}^{-1}$. The residual spectrum obtained by subtracting off the best-fit stellar model from the spectrum is then used to derive emission-line properties. This procedure allows us to recover the $H\beta$ narrow emission and therefore compute the Balmer decrement $F_{H\alpha}/F_{H\beta}$. Both the fitted starlight continuum and the residual emission-line spectrum of I20210N are shown in Fig. 4.

We note that the derived stellar-velocity dispersion value is very high compared to what is expected for typical galaxies: for example, a search in the catalogue of galactic dynamics by Forbes & Ponman (1999) yields only two elliptical/S0 objects with $\sigma_v^* > 300 \text{ km s}^{-1}$. Similarly, the stellar velocity dispersions measured by Falcón-Barroso et al. (2017) in a large sample of

galaxies from the CALIFA survey and by Perna et al. (2021) in a sample of nearby ULIRGs never exceed $\sim 200 \text{ km s}^{-1}$. Nevertheless, objects exhibiting exceptional values of σ_v^* exist: it is e.g. the case of NGC 6240 ($\sigma_v^* \sim 360 \text{ km s}^{-1}$; Doyon et al. 1994), which is indeed a final-state merging system. Therefore, the stellar velocity dispersion value found in I20210N may indicate that the internal kinematics of the galaxy is deeply altered by the gravitational interaction with I20210S.

We then fit the relevant emission lines with Gaussian profiles through the IDL minimization package MPFIT (Markwardt 2009). All the narrow components are simultaneously fitted considering them as emitted at the same distance from the AGN, i.e. with equal FWHM in the velocity space. In addition, we fix the intensities of the faint components of the [O III] and [N II] doublets to a ratio 1/3.06 with the respective dominant component (e.g., Osterbrock & Ferland 2006). In order to compute meaningful uncertainties of measurement for the free parameters, we iterate this process over 10^3 Monte-Carlo (MC) realizations of each line spectrum. Such realizations have fluxes at each wavelength altered by a random quantity extracted from a Gaussian distribution, which is centered at the specific flux value and wide as the corresponding 1σ rms flux error.

The best fit of the I20210N emission lines is shown in Fig. 4, along with the corresponding standardized residuals¹. The value of the Balmer decrement derived from this procedure is 2.97 ± 0.31 , compatible within errors to both the intrinsic ratio $I_{H\alpha}/I_{H\beta} \sim 2.85$ typical of [H II] region-like objects and the AGN ratio of 3.1 (Veilleux & Osterbrock 1987). Therefore, we can assume that the NLR of I20210N is viewed along an unreddened line of sight, with $E(B-V) \sim 0$. The best-fit parameters of the narrow emission lines are reported in Tab. 1, with FWHM corrected for the instrumental broadening $\Delta v_{\text{inst}} \sim 510 \text{ km s}^{-1}$ corresponding to the DOLoRes resolution of $\sim 10 \text{ Å}$:

$$\text{FWHM}_{\text{corr}} = \sqrt{\text{FWHM}_{\text{obs}}^2 - \Delta v_{\text{inst}}^2} \quad (2)$$

3.2. I20210N classification

To assess the nature of the AGN hosted in I20210N, we compute the [O III]/ $H\beta$, [N II]/ $H\alpha$, [S II]/ $H\alpha$, [O I]/ $H\alpha$ and [O II]/ $H\beta$ logarithmic line ratios from the best-fit emission line parameters. The derived values are presented in Tab. 1 along with their uncertainties. For completeness, we also report the value of the [O III]/[O II] ratio that is used to further classify active galaxies (e.g., Heckman 1980; Kewley et al. 2006).

The values derived for I20210N are plotted in the BPT diagrams shown in Fig. 5, superimposed to the values for SDSS-DR7 objects retrieved from the OSSY database (Oh et al. 2011, 2015). To discriminate between the different classes of emission-line galaxies (star-forming, Seyferts, LINERs), we adopt from the current literature the relations defining the boundaries between types of galactic activity:

- the extreme-starburst relation and the Seyfert-LINER boundaries by Kewley et al. (2001) in the original diagrams by Baldwin et al. (1981);
- the star-forming, Seyfert-LINER and mixed-region boundaries by Lamareille (2010) in the [O III]/ $H\beta$ -to-[O II]/ $H\beta$ diagram;

¹ The standardized residuals are computed everywhere as $[F_\lambda - F_\lambda^{(\text{BF})}]/\sigma_F$, where $F_\lambda^{(\text{BF})}$ is the best-fit model of the line flux and σ_F is the standard deviation of the dimensional residuals $F_\lambda - F_\lambda^{(\text{BF})}$ (e.g., Cook & Weisberg 1982).

Transition	I20210N ($\chi^2/\nu_{\text{d.o.f.}} = 427/405$)		I20210S ($\chi^2/\nu_{\text{d.o.f.}} = 442/405$)				South Nebula ($\chi^2/\nu_{\text{d.o.f.}} = 228/209$)		
	Flux (10^{-14} erg s^{-1} cm^{-2})	FWHM (km s^{-1})	Flux (10^{-14} erg s^{-1} cm^{-2})		FWHM (km s^{-1})	Blueshift (km s^{-1})	Flux (10^{-14} erg s^{-1} cm^{-2})	FWHM (km s^{-1})	Blueshift (km s^{-1})
			N	B					
[O II] $\lambda\lambda 3726, 3729$	1.060 ± 0.055	690 ± 40	10.4 ± 1.1	10.6 ± 4.4	530 ± 90	*	—	—	—
[O III] $\lambda 4363$	—	—	0.500 ± 0.054	0.42 ± 0.17	—	*	—	—	—
H β	0.214 ± 0.010	—	13.72 ± 0.94	6.3 ± 2.1	—	1960 ± 470	330 ± 180	0.0691 ± 0.0036	710 ± 330
[O III] $\lambda 5007$	1.058 ± 0.047	—	85.2 ± 5.4	32.1 ± 9.8	—	2080 ± 310	390 ± 120	0.1193 ± 0.0073	—
[O I] $\lambda 6300$	0.193 ± 0.010	—	4.13 ± 0.20	2.39 ± 0.57	—	1910 ± 350	230 ± 150	—	—
H α	0.623 ± 0.028	—	42.5 ± 2.0	19.5 ± 4.3	—	*	*	0.1418 ± 0.0071	—
[N II] $\lambda 6583$	0.863 ± 0.039	—	33.7 ± 1.6	8.0 ± 2.2	—	*	*	0.0334 ± 0.0047	—
[S II] $\lambda 6717$	0.436 ± 0.021	—	6.80 ± 0.31	2.98 ± 0.66	—	*	*	0.0371 ± 0.0012	—
[S II] $\lambda 6731$	0.366 ± 0.018	—	7.21 ± 0.32	5.0 ± 1.0	—	*	*	0.0273 ± 0.0010	—
$E(B - V)$	~0	—	0.271 ± 0.019	0.195 ± 0.091	—	—	—	~0	—

Ratio	I20210N		I20210S		Outflow		South Nebula	
	Value	Uncertainty	Value	Uncertainty	Value	Uncertainty	Value	Uncertainty
$\log([\text{O III}]/\text{H}\beta)$	0.686	0.040	0.798	0.061	0.71	0.28	0.237	0.049
$\log([\text{N II}]/\text{H}\alpha)$	0.142	0.039	-0.114	0.044	-0.39	0.21	-0.628	0.083
$\log([\text{S II}]/\text{H}\alpha)$	0.110	0.041	-0.489	0.044	-0.39	0.19	-0.343	0.037
$\log([\text{O I}]/\text{H}\alpha)$	-0.509	0.042	-0.933	0.052	-0.91	0.20	—	—
$\log([\text{O II}]/\text{H}\beta)$	0.695	0.043	-0.120	0.077	0.23	0.33	—	—
$\log([\text{O III}]/[\text{O II}])$	-0.001	0.042	0.913	0.073	0.48	0.31	—	—

— Transition not detected in the spectrum.

— Value fixed to be equal to the first non-null one upwards in the column.

* Value anchored to the best fit of the [O III] $\lambda 5007$ broad component.

Table 1. *Top:* best-fit parameters of the dereddened diagnostic emission lines in the optical spectrum of I20210N, I20210S NLR (“N” components) and outflowing emission (“B” components), and emission from the South Nebula. The integrated fluxes presented here have been dereddened by the indicated amount of $E(B - V)$ with the SMC extinction by Pei (1992). *Bottom:* values of the diagnostic line ratios with the corresponding uncertainties.

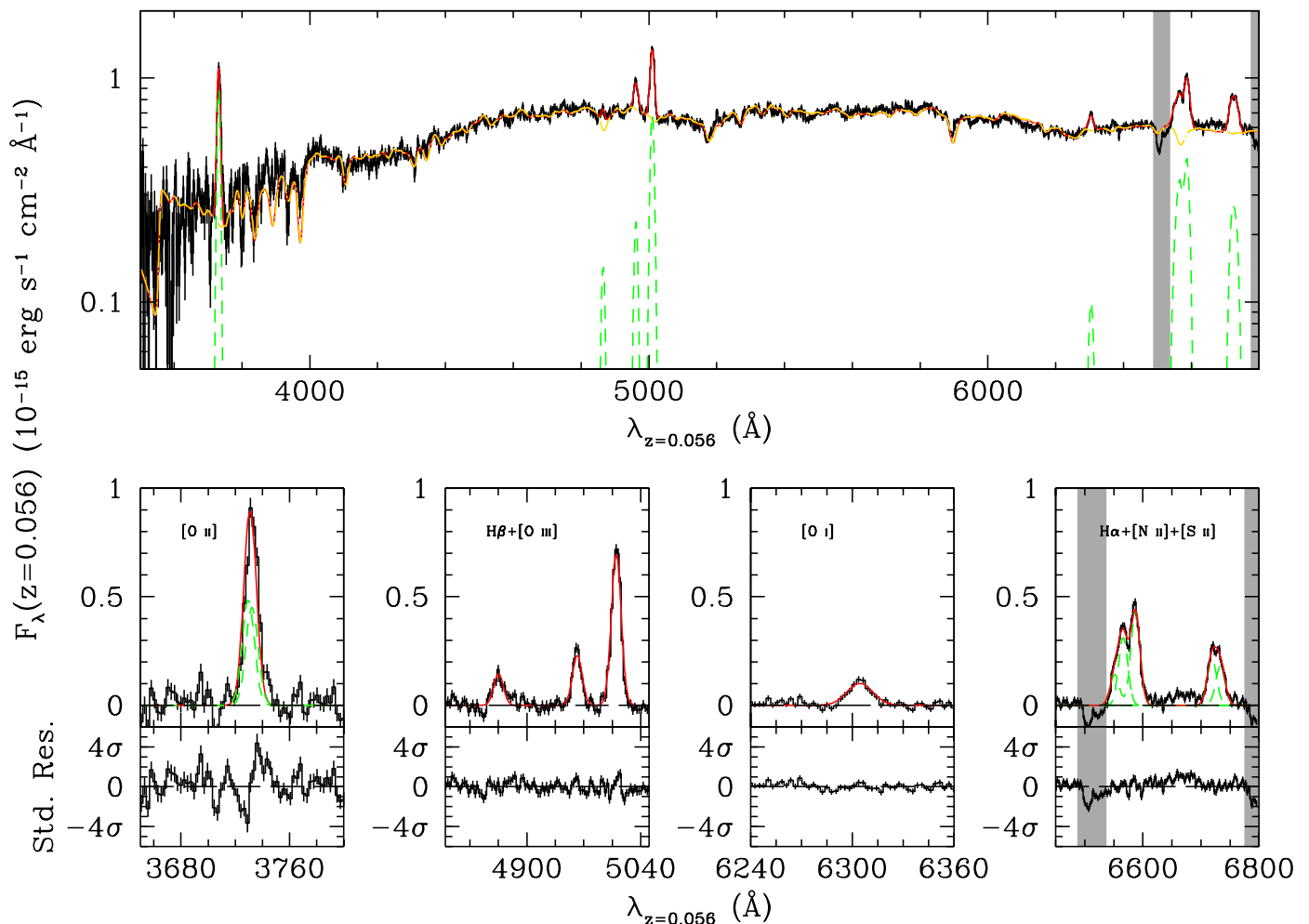


Fig. 4. *Upper panel:* rest-frame spectrum of I20210N (black solid line), along with the best-fit starlight model adopted for continuum subtraction (yellow dot-dashed line), the best-fit reddened emission profiles (green short-dashed lines), the global spectral model (red solid line) and the masks applied to the telluric absorption lines (grey bands) shown superimposed to the data. *Lower panels:* zoom on the continuum-subtracted emission lines (black solid line), shown along with the global best fit (red solid line) and the best-fit single components (green short-dashed lines) for the blended [O II] doublet, H α + [N II] and [S II] transitions. The standardized residuals after the best-fit subtraction are also shown in separate windows below each spectral region. In all panels, the zero-level flux (black long-dashed line) is indicated; in the panel with H α + [N II] and [S II], the masks applied to the telluric absorption lines (grey bands) are shown superimposed to the data.

– for the [O III]/H β -to-[N II]/H α diagram, the pure star-formation boundary by Kauffmann et al. (2003), the alternative Seyfert-LINER relation by Cid Fernandes et al. (2010) and the redshift-dependent star formation boundary by Kewley et al. (2013a, see also Kewley et al. 2013b) computed for $z = 0.128 \pm 0.044$, i.e. the average redshift value of the OSSY catalogue (Oh et al. 2015).

Fig. 5 shows that the I20210N line ratios are generally consistent with those found for Seyfert galaxies in the [O III]/H β -to-[N II]/H α and [O III]/H β -to-[O II]/H β diagrams, while being intermediate between a Seyfert and a LINER in the [O III]/H β -to-[S II]/H α diagram where they sit on the Seyfert-to-LINER boundary by Kewley et al. (2001). Due to their intermediate nature, Heckman (1980) calls this kind of objects as “transition galaxies” between Seyfert 1.9 (Osterbrock 1981) and LINERs. However, the severe unresolved blending that affects the H α + [N II] emission system may offer an explanation to this issue: in fact, decreasing intrinsic H α intensity would eventually

move the position of I20210N in both the [O III]/H β -to-[S II]/H α and [O III]/H β -to-[O I]/H α diagrams in the Seyfert region (neglecting the shift in the [O III]/H β -to-[N II]/H α diagram given by the consequent increase in the [N II] emission). In addition, the hard X-ray luminosity of $\sim 5 \times 10^{42}$ erg s $^{-1}$ measured for I20210N by Piconcelli et al. (2010) helps in breaking this uncertainty, pushing the classification of I20210N towards a Seyfert 2 galaxy.

4. Characterization of I20210 South

For both the nuclear spectrum and the South Nebula, we apply the same fitting procedure as done for I20210N. Therefore, we first proceed to model the I20210S continuum emission. We assume for I20210S an AGN-dominated object, i.e. exhibiting a power-law continuum reddened by foreground dust; such an assumption is well motivated by the quasar-like energetics of I20210S (e.g., $L_{\text{bol}} \sim 10^{45}$ erg s $^{-1}$; Piconcelli et al. 2010).

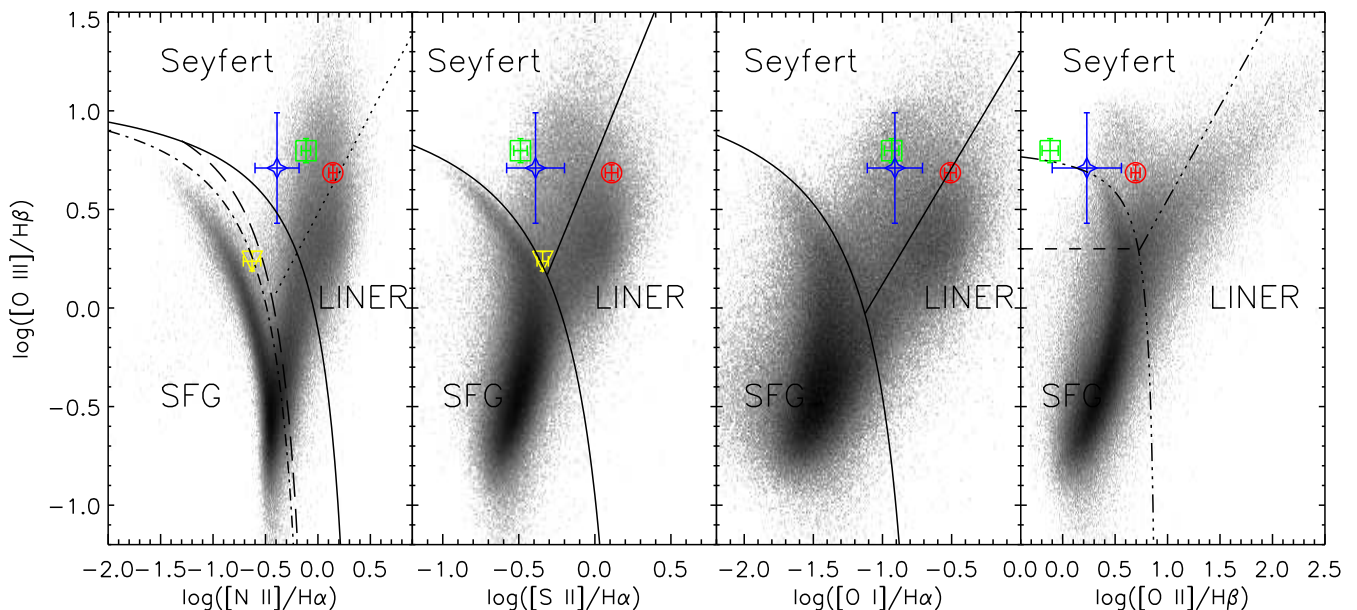


Fig. 5. BPT diagnostic diagrams showing the position of I20210N (red circle), I20210S (green square), its outflow (blue star) and the South Nebula (yellow triangle) with the relative uncertainties on top of the SDSS data from the OSSY database (Oh et al. 2015, grey dots). As a reference, in the first three panels the extreme-starburst and Seyfert-LINER classification boundaries by Kewley et al. (2001, solid lines) are indicated, along with the pure star-formation boundary by Kauffmann et al. (2003, long-dashed line), the alternative Seyfert-LINER relation by Cid Fernandes et al. (2010, dotted line) and the redshift-dependent relation at $z \sim 0.13$ by Kewley et al. (2013a, dot-dashed line) in the $[\text{O III}]/\text{H}\beta$ -to- $[\text{N II}]/\text{H}\alpha$ diagram. In the $[\text{O III}]/\text{H}\beta$ -to- $[\text{O II}]/\text{H}\beta$ diagram, the star-forming and Seyfert-LINER boundaries by Lamareille (2010, triple dot-dashed line) are indicated, along with the mixed-region boundary (short-dashed line).

4.1. I20210S nuclear continuum and emission-line fitting

The optical spectrum of I20210S is heavily reddened by intrinsic dust, which is particularly evident by the lack of a flux rise bluewards the $[\text{O II}]$ emission. However, due to the presence of a large number of emission and telluric lines that greatly reduce the intervals of featureless spectral regions, we decide not to perform a global continuum fit to be subtracted from the spectrum. Instead, we select sections of the I20210S spectrum free of major features and adjacent to the lines of interest, and interpolate them with local power laws in order to remove the underlying AGN emission (see Fig. 6).

Then, we use the MC fitting procedure to model the emission lines with Gaussian profiles, which are shown in Fig. 6 along with the corresponding standardized residuals. Differently from I20210N, we use two components for each transition in order to account for the total emission profile in the spectrum as also done by Arribas et al. (2014). The main narrow components obey the prescriptions presented in Sect. 3.1; the additional emission, as already found by P90, consists in a broad line blueshifted by $\sim 400 \text{ km s}^{-1}$ with $\text{FWHM} \sim 2000 \text{ km s}^{-1}$. The possibility that such features are an effect of the orientation of the I20210S narrow-line region (NLR) as described in Bisogni et al. (2017) is ruled out: in fact, for emission lines associated with permitted transitions the asymmetries would be redshifted with respect to the line center (see e.g. their figure 3). Therefore, we can conclude that I20210S exhibits evidence of an ionized gas outflow.

To account for this additional emission, we include the broad components in the fit of the I20210S line profiles anchoring the

blueshift and FWHM values of the transitions affected by severe blending – i.e. the $[\text{O II}] \lambda\lambda 3726, 3729$ doublet, the $\text{H}\gamma + [\text{O III}] \lambda 4363$, the $\text{H}\alpha + [\text{N II}]$ system and the $[\text{S II}]$ doublet – to those of the $[\text{O III}] \lambda 5007$ (see Tab. 1). This choice is motivated by the fact that the $[\text{O III}]$ emission has the highest S/N; in the cases in which the anchoring to its parameters is adopted, only the line amplitude is left free to vary. In addition, we estimate the amount of intrinsic dust extinction for the NLR and the outflow separately, since the two regions are in principle located at different distances from the central engine and can thus be affected by different amounts of reddening.

The Balmer ratio derived from the MC fit for the narrow components is $F_{\text{H}\alpha}/F_{\text{H}\beta} = 4.108 \pm 0.079$, corresponding to $E(B - V)_{\text{NLR}} = 0.271 \pm 0.019 \text{ mag}$ for the AGN intrinsic ratio $I_{\text{H}\alpha}/I_{\text{H}\beta} = 3.1$ (Veilleux & Osterbrock 1987) and the SMC extinction by Pei (1992), whereas a ratio $F_{\text{H}\alpha}/F_{\text{H}\beta} = 3.80 \pm 0.34$ for the outflow yields $E(B - V)_{\text{out}} = 0.195 \pm 0.091$. Finally, we apply the extinctions derived in this way to deredden the corresponding emission-line amplitudes. The best fit of the reddened I20210S spectrum is shown in Fig. 6, whereas the best-fit parameters of both its NLR and outflow emission are reported in Tab. 1. On average, the I20210S wind has an outflow velocity $\Delta v = 330 \pm 170 \text{ km s}^{-1}$ and $\text{FWHM} = 2000 \pm 390 \text{ km s}^{-1}$: such values are a factor of ~ 2.4 higher than the corresponding mean parameters found by Arribas et al. (2014) in ULIRGs hosting AGN (see their table 2), and more in line with those found by Rodríguez Zaurín et al. (2013) for ionized outflows in nearby ULIRGs (see their table 2) and by Zakamska et al. (2016) in

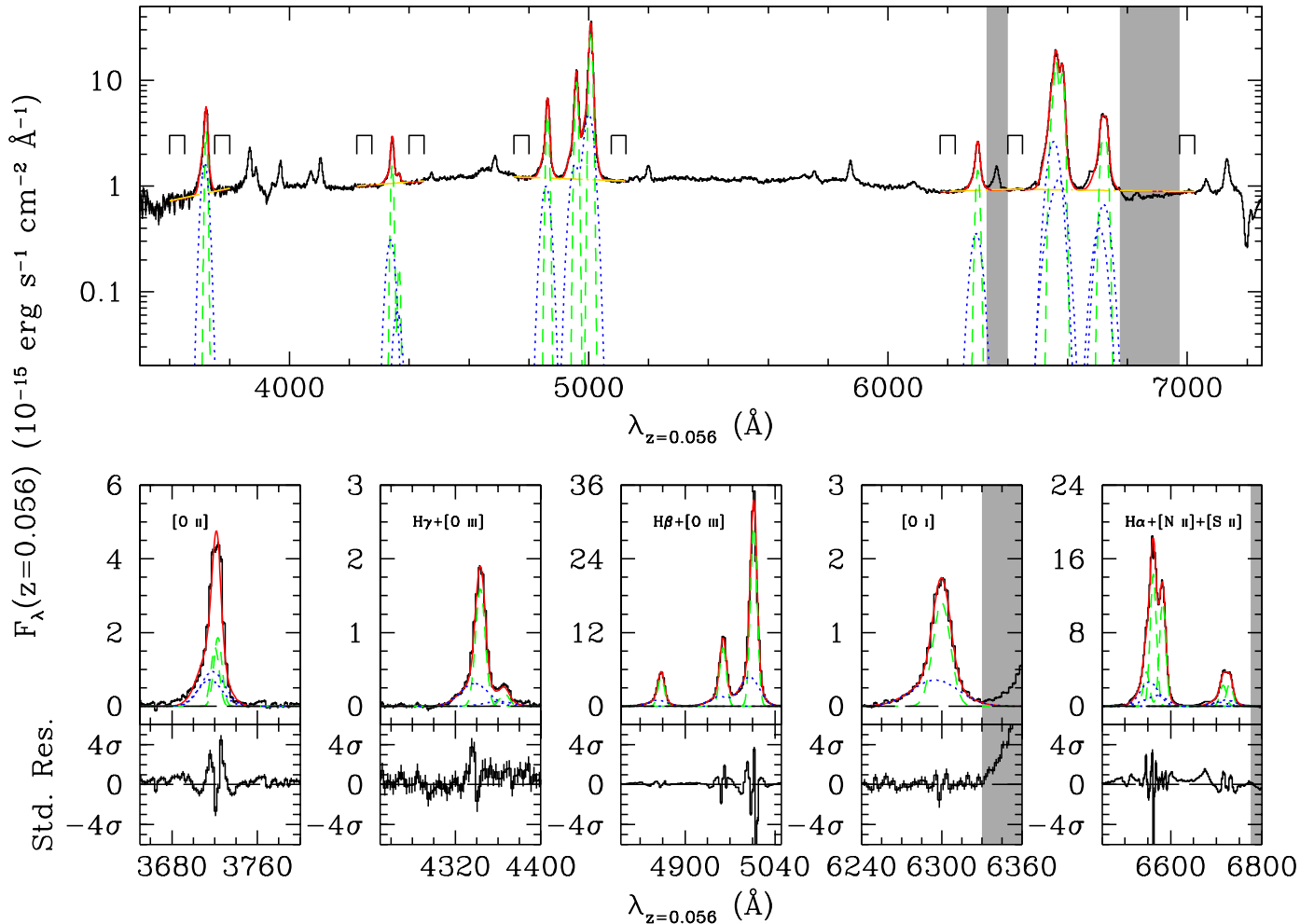


Fig. 6. *Upper panel:* rest-frame spectrum of I20210S (black solid line), along with the local power laws adopted for continuum subtraction (yellow dot-dashed lines), the best-fit reddened narrow (green short-dashed lines) and broad emission profiles (blue dotted lines) and the global spectral model (red solid line) shown superimposed to the data. The intervals used for the local continuum subtraction under the emission lines are marked, along with the masks applied to the [O I]+[Fe X] emission and to the telluric H₂O absorption line (grey bands). *Lower panel:* zoom on the continuum-subtracted emission lines (black solid line), shown along with the global best fit (red solid line) and the best-fit narrow (green short-dashed lines) and broad components (blue dotted lines). The standardized residuals after the best-fit subtraction are also shown in separate windows below each spectral region. In all panels, the zero-level flux (black long-dashed line) is indicated; in the panel with [O I], the mask applied to the [O I]+[Fe X] emission line is shown superimposed to the data, as well as the one for the H₂O telluric absorption in the panel with Hα+[N II] and [S II] (grey bands).

high- z reddened quasars (see their table 1) where the emission-line profiles are modeled using multiple Gaussians.

4.2. The South Nebula spectrum

We now estimate the intrinsic reddening of the South Nebula. The detection of both H β and H α narrow transitions allows us to apply the MC line-fitting procedure described in the case of I20210N (see Fig. 7 and Sect. 3.1). Lacking any trace of an underlying continuum which could have been used in the determination of the reddening law, we adopt the SMC extinction by Pei (1992) as in the case of the parent nucleus. This in turn yields $F_{\text{H}\alpha}/F_{\text{H}\beta} = 2.05 \pm 0.10$, which is lower than the intrinsic ratio of 2.85 valid for [H II] regions. Also, the low associated error of measurement potentially indicates a poorly determined estimate

of the Balmer ratio, likely due to the uncertainties in extracting a continuumless spectrum that is ~ 100 times less intense than the I20210S nuclear emission. Therefore, the issue of determining the South Nebula intrinsic reddening is clearly matter for future more sensitive, spatially-resolved spectroscopic data; in the following, we consider it compatible with $E(B - V) \sim 0$.

We then fit the parameters of the five emission features clearly identified, namely H β , [O III], H α , [N II] and [S II], without applying any dereddening (see Tab. 1). Interestingly, the South Nebula exhibits a blueshift of $550 \pm 150 \text{ km s}^{-1}$ with respect to the systemic redshift and a FWHM of $710 \pm 330 \text{ km s}^{-1}$. Such features are a clear indication of highly disrupted gas (Bellocchi et al. 2013), similar to that found by Ramos Almeida et al. (2017) in the Teacup Galaxy ($L_{[\text{O III}]}$ $\sim 5 \times 10^{42} \text{ erg s}^{-1}$ according to Reyes et al. 2008, to be compared with $L_{[\text{O III}]}$ $\sim 6.5 \times 10^{42} \text{ erg$

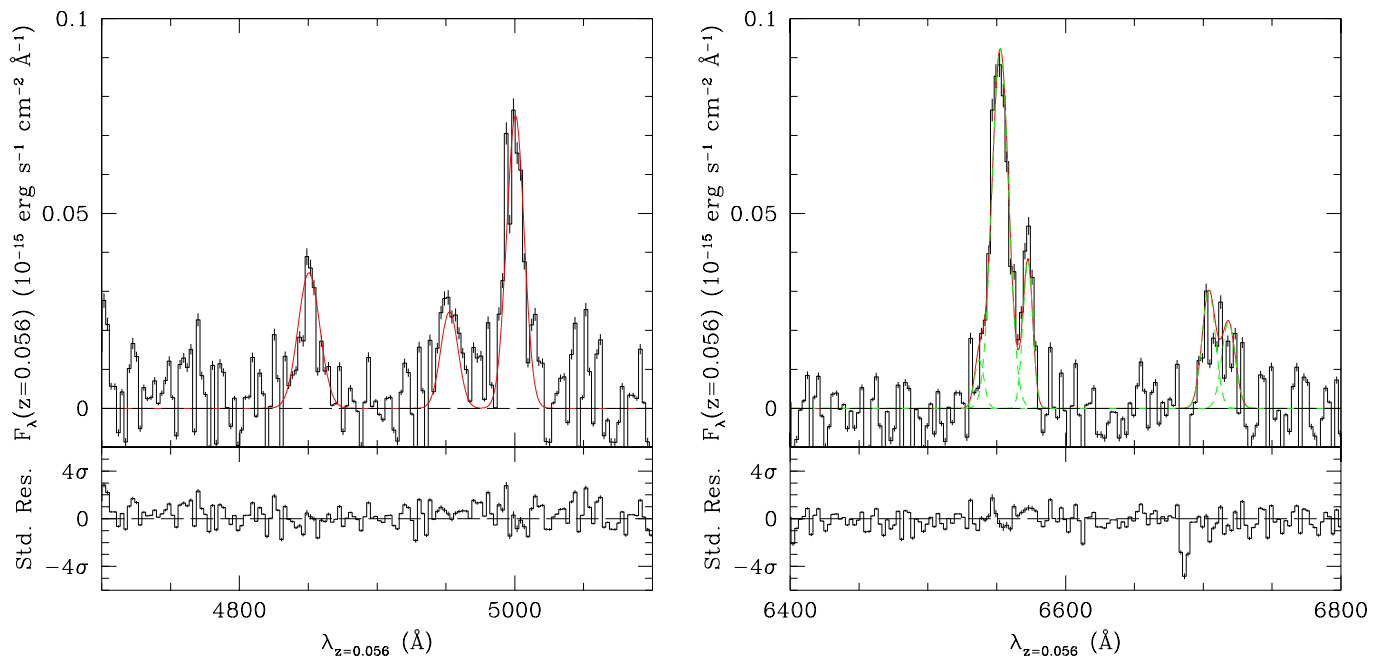


Fig. 7. *Left panel:* best-fit profile of the $H\beta$ + $[O\text{ III}]$ emission lines from the South Nebula in the rest frame. *Right panel:* best-fit profile of the $H\alpha$ + $[N\text{ II}]$ and $[S\text{ II}]$ emission lines from the South Nebula in the rest frame. In both panels, the global emission profile (red solid line) is shown superimposed to the line spectrum (black histogram) along with the single components of the $H\alpha$ + $[N\text{ II}]$ and $[S\text{ II}]$ blended profiles (green short-dashed lines), and the zero-flux level (black long-dashed line) is indicated. The standardized residuals after the best-fit subtraction are also shown in separate windows below each spectral region.

s^{-1} for I20210S) at comparable distances (~ 5.6 kpc) from the central engine.

4.3. I20210S classification

As done in Sect. 3.2 for I20210N, we compute the line ratios for all the regions decomposed from the spectrum of I20210S, namely the NLR, the outflow and the South Nebula, and place them in the relevant BPT diagrams in order to obtain a first discrimination between an AGN or star-formation powered emission. A visual inspection confirms that the NLR properties are fully consistent with their AGN nature, with also the outflow emission falling well inside the AGN region in agreement with the scenario of an ionized wind driven by the nuclear activity.

The South Nebula sits close enough to the boundary between AGN and star-forming galaxies to prevent its straightforward inclusion among the AGN-powered processes. However, the kinematic properties of this region (velocity blueshift of ~ 500 km s^{-1} , FWHM of ~ 700 km s^{-1}) may be actually interpreted as due to the I20210S outflow, which has stripped out ionized gas from the I20210S nucleus. The possibility that the South Nebula is an extended NLR component blown out of the central engine by radiation pressure is in principle supported by studies that ubiquitously find NLRs extended over ~ 10 kpc from the central engine in both Type 1 and Type 2 quasars (see e.g. Husemann et al. 2013, and refs. therein); in the case of I20210S, it is less likely than extended, non-outflowing gas associated with extreme mergers, due to the fact that blown-out NLRs typically show $\text{FWHM} < 250$ km s^{-1} (Bellocchi et al. 2013).

Our data do not allow a deeper exploration of the spectral properties of the South Nebula. Therefore, we point out that, due to the intermediate values of its diagnostic line ratios between AGN and star-forming galaxies, this detached emitting region is a very interesting environment in which the effects of AGN

feedback may be at work in pushing the gas outside the central region of the host galaxy (negative feedback) while also triggering some amount of star formation into it (positive feedback; see e.g. Maiolino et al. 2017), and is therefore worth of further investigation with high-quality spatially resolved spectroscopy.

5. Physical properties of the outflow in I20210S

We are now able to characterize the physics of the I20210S ionized wind. To this aim, we estimate the outflowing mass M_{out} and the mass loss rate \dot{M} of the ionized gas following the method presented in Kakkad et al. (2016) and Bischetti et al. (2017). Under the assumptions that:

- the AGN wind is free, spherically or biconically symmetric, and mass-conserving;
- the AGN wind has mass-outflow rate and velocity independent on the outflow radius (Rupke et al. 2002, 2005);
- most of the oxygen consists of O^{2+} ions,

we can use the relation by Carniani et al. (2015):

$$\log\left(\frac{M_{\text{out}}}{M_{\odot}}\right) = 7.6 + \log\left(\frac{C}{10^{(O/H)-10}/M_{\odot}}\right) + \log\left(\frac{L_{[\text{O III}] \lambda 5007}^{\text{out}}}{10^{44} \text{ erg s}^{-1}}\right) - \log\left(\frac{\langle n_e \rangle}{10^3 \text{ cm}^{-3}}\right) \quad (3)$$

where $C = \langle n_e \rangle^2 / \langle n_e^2 \rangle$, $[O/H] - [O/H]_{\odot}$ is the gas metallicity relative to the solar value, $L_{[\text{O III}] \lambda 5007}^{\text{out}}$ is the outflowing $[O\text{ III}] \lambda 5007$ luminosity and $\langle n_e \rangle$ is the average electron density. The latter is in turn related to the electron temperature T_e , which can be derived from the line ratios $(I_{4959} + I_{5007}) / I_{4363}$ of the outflow emission (Osterbrock & Ferland 2006):

$$\frac{I_{4959} + I_{5007}}{I_{4363}} \approx 7.90 \cdot \exp\left(\frac{32,900 \text{ K}}{T_e}\right) \quad (4)$$

From the decomposition of the emission lines in the spectrum of I20210S through the fit with multiple Gaussian components described in Sect. 3 (see Fig. 6), we compute $(I_{4959} + I_{5007})/I_{4363} = 100 \pm 70$, which corresponds to $T_e = 12,900 \pm 3600$ K and is in agreement with the value of $\sim 10^4$ K generally assumed for AGN outflows (see e.g. Perna et al. 2017, and refs. therein).

The electron density $\langle n_e \rangle$ is then related to the ratio I_{6717}/I_{6731} between the components of the [S II] doublet through:

$$\frac{I_{6717}}{I_{6731}} = 1.49 \cdot \frac{1 + 3.77x}{1 + 12.8x} \quad (5)$$

with $x = 10^{-2} \langle n_e \rangle T_e^{-1/2}$ (Weedman 1968; Osterbrock & Ferland 2006; Sanders et al. 2016). However, we compute a ratio $I_{6717}/I_{6731} = 0.60 \pm 0.25$ for the ionized wind which is on the saturating side of Eq. 5, and only allows us to establish a lower limit at 95% probability of $\langle n_e \rangle \gtrsim 4000 \text{ cm}^{-3}$ to the outflow electron density. This might either be an indication of a high electron density or just a consequence of the severe blending that affects the [S II] region at the DOLoRes low spectral resolution, preventing us from deriving a reliable estimate of $\langle n_e \rangle$. The same issue holds for the [O II] doublet, which could have been used in place of the [S II] for such a measurement (Osterbrock & Ferland 2006) but is even more blended because of its peak separation of ~ 3 Å only.

As an alternative possibility to derive solid estimates of $\langle n_e \rangle$ for the I20210S outflow, we also consider the application of the trans-auroral ratio (TR) method by Rose et al. (2018). This method, based on the evaluation of the line ratios $[S II]_{4068,4076}/[S II]_{6717,6731}$ and $[O II]_{3726,3729}/[O II]_{7319,7331}$ allows us to obtain at once both $\langle n_e \rangle$ and the intrinsic reddening $E(B-V)$ of the emitting gas. We thus fit the trans-auroral doublets [S II] $\lambda\lambda 4068,4076$ and [O II] $\lambda\lambda 7319,7331$ through the MC procedure with two narrow and two broad components each, fixing their widths to the corresponding values for the [O III] $\lambda\lambda 4959,5007$ (see Tab. 1). This yields $TR([S II])_{\text{out}} = 0.192 \pm 0.051$ and $TR([O II])_{\text{out}} = 1.72 \pm 0.94$. Having derived for the outflow a ionization parameters $\log U_{\text{out}} = -3.09 \pm 0.47$ from its relation to the [O III]/H β and [N II]/H α ratios (Baron & Ménard 2019, BM19 hereafter), we finally compare its TRs to the simulations presented in Davies et al. (2020, see their figure 7), obtaining $\langle n_e \rangle_{\text{out}} = 10,400^{+4000}_{-5200} \text{ cm}^{-3}$ and $E(B-V)_{\text{out}} = 0.34^{+0.24}_{-0.15}$.

As pointed out in the literature (Rose et al. 2018; Spence et al. 2018; Davies et al. 2020), the TR method allows to probe denser gas with respect to the use of the “traditional” [S II] doublet, whose emission is likely produced at the ionization front where the electron density significantly decreases. This issue is probably at the base of the high values of ionized gas mass and mass outflow rate recently found in AGN winds (e.g., Carniani et al. 2015; Kakkad et al. 2016; Bischetti et al. 2017; Perna et al. 2017), for which values of $10^2 \text{ cm}^{-3} \lesssim \langle n_e \rangle \lesssim 10^3 \text{ cm}^{-3}$ are usually assumed. Such an assumption is justified from measurements of the outflow electron density based on the [S II] method: for example, Arribas et al. (2014) get $\langle n_e \rangle \sim 400 \text{ cm}^{-3}$ for the outflowing emission in ULIRGs, whereas Perna et al. (2020) find $\sim 200 \text{ cm}^{-3}$ in the archetypal ULIRG Arp 220. For comparison, the values of $\langle n_e \rangle$ found by Rose et al. (2018) for AGN-driven outflows in ULIRGs fall in the range $3000 \div 56,000 \text{ cm}^{-3}$, with a median value of $\sim 5000 \text{ cm}^{-3}$. Also, Kakkad et al. (2018) obtain spatially-resolved values of $\langle n_e \rangle$ up to $\sim 2000 \text{ cm}^{-3}$ for ionized winds in nearby radio-selected Seyfert galaxies. Due to the limited DOLoRes spectral resolution and the severe blending that affects the I20210S trans-auroral emission lines with nearby features (e.g., the H δ close to the [S II] $\lambda\lambda 4068,4076$, the He I blue-

Quantity	Value		Units
	Outflow	South Nebula	
T_e	$12,900 \pm 3600$	$\sim 10,000$	K
$\langle n_e \rangle$	$\gtrsim 5000$	~ 100	cm^{-3}
$L_{[\text{O III}]}^{\text{out}}$	$(2.44 \pm 0.74) \times 10^{42}$	$(9.05 \pm 0.55) \times 10^{39}$	erg s^{-1}
v_{max}	2160 ± 380	1100 ± 430	km s^{-1}
R_{out}	2.20 ± 0.14	6.52 ± 0.43	kpc
t_{dyn}	0.99 ± 0.27	5.8 ± 2.6	Myr
M_{out}	$(1.94^{+0.69}_{-0.51}) \times 10^5$	$\sim 3 \times 10^4$	M_{\odot}
\dot{M}	$0.59^{+0.46}_{-0.26}$	$\sim 6 \times 10^{-3}$	$M_{\odot} \text{ yr}^{-1}$
\dot{E}_{kin}	$(0.86^{+1.27}_{-0.54}) \times 10^{42}$	$\sim 2 \times 10^{39}$	erg s^{-1}
\dot{P}_{out}	$(0.80^{+0.88}_{-0.44}) \times 10^{34}$	$\sim 4 \times 10^{31}$	erg cm^{-1}

Table 2. Summary of the relevant physical properties of the ionized outflow discovered in the I20210S optical spectrum and of the South Nebula. *Upper section:* quantities independent of the electron density. *Lower section:* quantities depending on the electron density, for which a value of $\langle n_e \rangle = 5000 \text{ cm}^{-3}$ (Rose et al. 2018) is assumed in the case of the outflow. Note that the quoted errors of measurement are only indicative of the magnitude of the statistical uncertainties, not the systematics, affecting the computed values (see Sect. 5).

ward and the [Ni II] redward of the [O II] $\lambda\lambda 7319,7331$), we cannot draw any firm conclusion on the reliability of the I20210S outflow electron density derived with the TR method. Therefore, in the following discussion of the physical properties of the I20210S ionized wind we adopt $\langle n_e \rangle \sim 5000 \text{ cm}^{-3}$ (Rose et al. 2018) as our main reference when computing all the related quantities.

With $L_{[\text{O III}]}^{\text{out}} = (2.44 \pm 0.74) \times 10^{42} \text{ erg s}^{-1}$ obtained from the outflow [O III] flux reported in Tab. 1, and the further assumptions of $C \approx 1$ and $[\text{O}/\text{H}] \sim [\text{O}/\text{H}]_{\odot}$ (i.e. solar metallicity), Eq. 3 yields $M_{\text{out}} = (1.94^{+0.69}_{-0.51}) \times 10^5 M_{\odot}$. Clearly, this value and those based on it are affected by the assumption on $\langle n_e \rangle$. We then derive the expression of the outflowing mass rate \dot{M} from the fluid-field continuity equation as done in Bischetti et al. (2017), in order to provide a local estimate of this quantity at the outflow termination radius R_{out} (e.g., Feruglio et al. 2015):

$$\dot{M} = 3 \frac{M_{\text{out}} v_{\text{max}}}{R_{\text{out}}} \quad (6)$$

In order to estimate the spatial extension of the outflow, we perform a series of adjacent, 1-px wide (i.e. ~ 0.27 kpc, owing to the DOLoRes angular scale of $0.252 \text{ arcsec px}^{-1}$ and the scale distance of $1.087 \text{ kpc arcsec}^{-1}$ at $z = 0.056$) extractions of the I20210S spectrum along its 2D trace in the high-S/N H β + [O III] region.

First, we fit a Gaussian function to the trace profile at the [O III] peak to get the trace width $\sigma = 1.67 \pm 0.01 \text{ px}$ (i.e. ~ 0.46 kpc). Then, we extract 1D off-axis spectra in both the Northern and Southern direction offset by 3 to 12 px (~ 0.8 to ~ 3.3 kpc) from the aperture center (Perna et al. 2015; Bischetti et al. 2017), in order to both exclude the signal enclosed in the instrumental

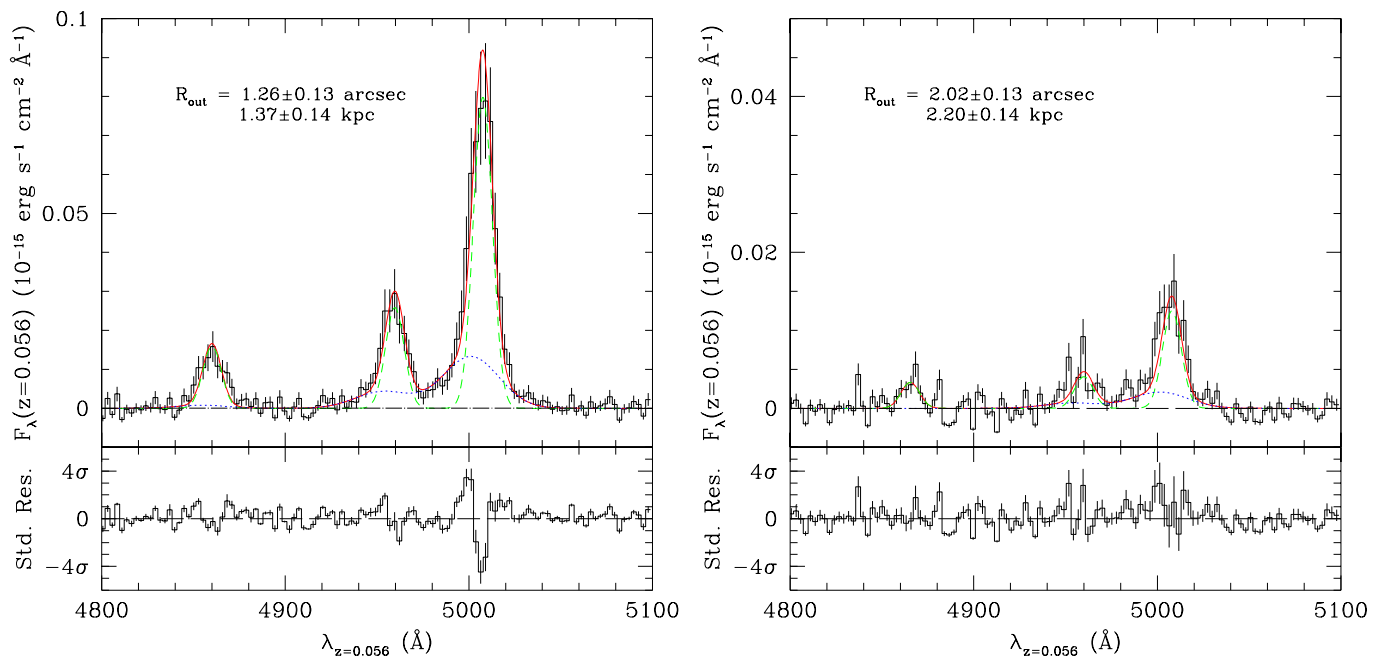


Fig. 8. Continuum-subtracted I20210S off-axis H β + [O III] spectra extracted at a 5-px offset in the Northern direction (*left panel*) and at an 8-px offset in the Southern direction (*right panel*). In each panel, the best fit model is shown (*red solid line*) along with the profiles of the narrow (*green short-dashed line*) and broad components (*blue dotted line*), and the zero-flux level is indicated (*black long-dashed line*). The standardized residuals after the best-fit subtraction are also shown in separate windows below each spectral region. The fit to the H β emission is not accounted for the calculations of χ^2 and p_F , which are performed on the [O III] doublet only (see text), and is shown here for visual purposes only.

PSF of $0''.85$ ($0''.43$ in each direction, i.e. 2.5 px) and avoid overlap with either I20210N or the South Nebula. Finally, we calibrate such spectra with the same wavelength dispersion and sensitivity function applied to the I20210S average spectrum, and apply the MC fitting procedure to the continuum-subtracted [O III] emission only letting the line amplitudes free to vary – narrow FWHM, broad FWHM and blueshifts are fixed to the values reported in Tab. 1. We produce the MC fit for both the two-component model and a comparison single-component model of the emission lines, in which the broad emission from the outflow is neglected. In this way, we are able to identify through a statistical F -test the transition region where the outflow signal becomes negligible with respect to the NLR emission: specifically, we define the significance threshold of the outflow by requesting an F -test probability $p_F > 0.90$.

The statistical analysis yields significant emission associated with the outflow up to 5 px (~ 1.3 kpc) in the Northern direction, with $p_F \geq 0.93$ ($\chi^2/\nu_{d.o.f} = 76/73$); at larger distances along this direction, the signal emitted from the ionized wind quickly becomes undistinguishable from the noise, and has therefore no impact on the best fit ($p_F = 0$). Instead, the outflow emission in the Southern direction remains significant out to 8 px (~ 2 kpc, $p_F \geq 0.93$, $\chi^2/\nu_{d.o.f} = 78/73$). The “terminal” line spectra extracted at 5-px Northward and 8-px Southward offset are shown in Fig. 8. From this point on, we therefore adopt the distance $R_{\text{out}} = 2.20 \pm 0.14$ kpc (i.e. 8.0 ± 0.5 px) as our fiducial value for the termination radius of the I20210S ionized wind within the applicability limits of the F -test statistics. This choice is motivated by the observation that an 8-px offset produces an extraction lying beyond 3σ pixels from the trace center, and thus fiducially outside the 2.5-px PSF radius.

Then, we calculate $v_{\text{max}} = |\Delta v|_{[\text{O III}]}^{\text{out}} + 2\sigma_{[\text{O III}]}^{\text{out}} = 2160 \pm 380$ km s $^{-1}$ (see Bischetti et al. 2017, and refs. therein) from the out-

flowing [O III] $\lambda 5007$ parameters reported in Tab. 1. In this way, from Eq. 6 we obtain $\dot{M} = 0.59^{+0.46}_{-0.26}$ M_{\odot} yr $^{-1}$. Finally, we derive the outflow kinetic power \dot{E}_{kin} , the dynamical time scale t_{dyn} and the outflow momentum rate \dot{P}_{out} as:

$$\dot{E}_{\text{kin}} = \frac{1}{2} \dot{M} v_{\text{max}}^2 \quad (7)$$

$$t_{\text{dyn}} \approx \frac{R_{\text{out}}}{v_{\text{max}}} \quad (8)$$

$$\dot{P}_{\text{out}} = \dot{M} v_{\text{max}} \quad (9)$$

which yield $\dot{E}_{\text{kin}} = (0.86^{+1.27}_{-0.54}) \times 10^{42}$ erg s $^{-1}$, $t_{\text{dyn}} = 0.99 \pm 0.27$ Myr and $\dot{P}_{\text{out}} = (0.80^{+0.88}_{-0.44}) \times 10^{34}$ erg cm $^{-1}$, respectively. We report all these quantities in Tab. 2, highlighting that, given the high reference value of ~ 5000 cm $^{-3}$ adopted for the outflow (n_e), the electron-density dependent parameters might be underestimated by a factor of $\sim 10 \div 50$.

It is to be noted that, although the quantities presented in Tab. 2 are given with errors, they should actually be interpreted as rough estimates, since the systematic uncertainties insisting on Eq. 3 are biasing them by $1 \div 2$ orders of magnitude (Bischetti et al. 2017). Given this caveat, the values of M_{out} , \dot{M} and \dot{E}_{kin} obtained for the outflow of I20210S are in line with those found by Rupke & Veilleux (2013) for a sample of nearby galaxy mergers (see also Rodríguez Zaurín et al. 2013; Spence et al. 2018). In order to definitely assess the AGN nature of the I20210S outflow, we compare its kinetic power to the expected ejected mass rate \dot{M}_{SN} , energy output \dot{E}_{SN} and momentum injection \dot{P}_{SN} of starbursts associated with supernova (SN) explosions (Brusa et al. 2015). According to Veilleux et al. (2005), such quantities are related to the host-galaxy star formation rate (SFR) by:

$$\dot{M}_{\text{SN}} \lesssim 0.26 \left(\frac{\text{SFR}}{M_{\odot} \text{ yr}^{-1}} \right) \quad (10)$$

$$\dot{E}_{\text{SN}} \lesssim 7 \times 10^{41} \left(\frac{\text{SFR}}{\text{M}_{\odot} \text{ yr}^{-1}} \right) \quad (1)$$

$$\dot{P}_{\text{SN}} \lesssim 5 \times 10^{33} \left(\frac{\text{SFR}}{\text{M}_{\odot} \text{ yr}^{-1}} \right) \quad (1)$$

whereas the SFR is linked to the host-galaxy IR ($8 \div 1000 \mu\text{m}$) luminosity L_{IR}^* by (Kennicutt 1998; Kennicutt & Evans 2011; Kennicutt & De Los Reyes 2021):

$$\frac{\text{SFR}}{\text{M}_{\odot} \text{ yr}^{-1}} = 3.9 \times 10^{-44} \left(\frac{L_{\text{IR}}^*}{\text{erg s}^{-1}} \right) \quad (1)$$

We estimate $L_{\text{IR}}^* \sim 3.4 \times 10^{44} \text{ erg s}^{-1}$ for I20210S from the values for the total IR luminosity of the I20210 system and the AGN IR luminosity for the single members presented in Imanishi & Saito (2014), who estimated the contributions to the total galaxy IR emission coming from the active nucleus through photometric aperture size at high spatial resolution (see their table 1, 3 and 5). This in turn yields $\text{SFR} \sim 13 \text{ M}_{\odot} \text{ yr}^{-1}$, and hence $\dot{M}_{\text{SN}} \lesssim 3.4 \text{ M}_{\odot} \text{ yr}^{-1}$, $\dot{E}_{\text{SN}} \lesssim 9 \times 10^{42} \text{ erg s}^{-1}$ and $\dot{P}_{\text{SN}} \lesssim 6.5 \times 10^{33} \text{ erg cm}^{-1}$. Such values are ~ 6 to ~ 10 times higher than those listed in Tab. 2. Therefore, a starburst at work in I20210S is potentially able to produce the observed ionized outflow. However, Veilleux et al. (2005) note that Eqs. 10 to 12 give the limit values for a thermalization efficiency of 100% – i.e. when none of the starburst-injected energy is radiated away. Since typical starburst thermalization efficiencies are of the order of $\sim 10\%$ (see Veilleux et al. 2005, and refs. therein), the actual energy output from SNe is expected to be at most in line with the values listed in Tab. 2. This, in combination with an IR emission powered by the AGN (Imanishi & Saito 2014), leads us to conclude that the wind in I20210S is likely AGN-driven, although a non-negligible contribution from star formation cannot be ruled out. Establishing the main driving mechanism of the I20210S outflow is even further challenged by the uncertainty in its electron density, which biases the derivation of the physical properties that can be directly compared with the expected SN energetics; future spatially resolved observations of I20210S will therefore be of paramount importance to also precisely assess the nature of its ionized wind.

Broadened emission lines in ULIRGs hosting Seyfert nuclei are a common feature. Rodríguez Zaurín et al. (2013) report that up to 94% of nearby objects of this kind ($z < 0.175$) show strongly blueshifted ($\Delta v > 150 \text{ km s}^{-1}$) [O III] broad emission components ($\text{FWHM} > 500 \text{ km s}^{-1}$) that are emitted by near-nuclear ($R_{\text{out}} \lesssim 3.5 \text{ kpc}$) warm ionized outflows, albeit – while well detectable in optical and UV spectra – usually such outflows are not able to inject enough power into the AGN surrounding environment to effectively affect the host-galaxy ISM and star formation. In the face of a required \dot{E}_{kin} of the order of 0.5% to 5% of the total AGN radiant energy (Di Matteo et al. 2005; Hopkins & Elvis 2010), Fiore et al. (2017) show in fact that the majority of near-nuclear warm outflows clusters around $\dot{E}_{\text{kin}} \sim 0.001 L_{\text{bol}}$ (see their figure 1). The I20210S outflow appears to be consistent with this scenario, given its $\dot{E}_{\text{kin}}/L_{\text{bol}}$ ratio of ~ 0.002 ; however, its power can still be sufficient to *locally* affect the star formation rate in some host-galaxy regions, as demonstrated by the anti-correlation found between the distribution of star-forming clouds and wind zones in AGN hosts over resolved spatial regions $\sim 3 \div 7$ -kpc wide (Cano-Díaz et al. 2012; Carniani et al. 2015; Cresci et al. 2015). Given its proximity, brightness and spatial structure, I20210 therefore configures as

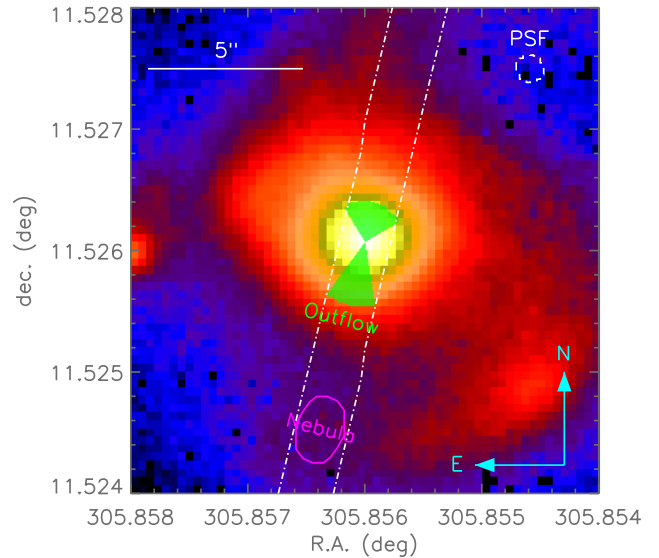


Fig. 9. Schematic representation of the emitting components observed in the spectrum of I20210S: the high-velocity outflow (green shaded area) – assumed biconical for simplicity, since only one slit position is available and the PSF size prevents us from resolving the morphology of such a small structure – and the South Nebula (magenta solid ellipse), plotted on top of the PS1 *grizy* image of the galaxy. The slit position and orientation (white dot-dashed lines) are reported; in the upper left corner, the diameter of $0''.85$ of the DOLoRes PSF (white dashed circle) is also indicated.

an extremely peculiar laboratory in which the impact and interplay of ongoing galaxy merging on both star formation and AGN activity can be investigated in great detail.

6. Physical properties of the South Nebula

In this section we briefly discuss the properties of the South Nebula. As previously described in Sect. 4.2, this region exhibits in fact interesting intermediate ionization properties between AGN-powered ($\text{FWHM} \sim 700 \text{ km s}^{-1}$, velocity blueshift of $\sim 500 \text{ km s}^{-1}$) and star-forming gas clouds. We present the schematic structure of I20210S in Fig. 9, overlapping the position and extension – within the spectrograph slit – of both the outflow and the South Nebula to the PS1 image of the galaxy. From this picture, it is evident that the outflow propagating Southwards extends outside enough of the innermost nuclear region to touch the inner regions of the disk structure, potentially interacting with the I20210S baryonic reservoir. In addition, the South Nebula is interestingly located on both the extensions of the outflow and the I20210S Western spiral arm.

According to Heckman et al. (2000), a galaxy with $L_{\text{IR}} \sim 3 \times 10^{45} \text{ erg s}^{-1}$ as I20210S (Piconcelli et al. 2010) has an average rotational velocity (v_{rot}) of $\sim 250 \text{ km s}^{-1}$; at a distance of $\sim 2.2 \text{ kpc}$ from the central engine, this translates to an escape velocity $v_{\text{esc}} \sim 550 \text{ km s}^{-1}$ assuming a galactic radius of $\sim 6.5 \text{ kpc}$ (i.e. the distance of the South Nebula). For comparison, the wind has a maximum velocity ~ 4 times higher (see Tab. 2), and the South Nebula itself has $v_{\text{max}} = 1100 \pm 430 \text{ km s}^{-1}$ (~ 2 times higher). Therefore, this implies that the nebula is being ejected outside the host galaxy by the ionized outflow, that also triggers possible star formation activity via quasar feedback as suggested

by the placement of the South Nebula line ratios on the AGN-star formation boundary (see Sect. 4.3).

It is also interesting to characterize its physical properties as done in Sect. 5 for the main outflow. To this end, we assume $T_e \sim 10^4$ K, which according to Eq. 5 translates to $\langle n_e \rangle \sim 100$ cm $^{-3}$ for the South Nebula [S II] ratio $I_{6717}/I_{6731} = 1.36 \pm 0.10$ (see Tab 1). This in turn yields $M_{\text{out}} \sim 3 \times 10^4 M_{\odot}$, given $L_{[\text{O III}]}$ = $(9.05 \pm 0.55) \times 10^{39}$ erg s $^{-1}$ from the [O III] flux listed in Tab. 1, and finally $\dot{M} = M_{\text{out}} v_{\text{max}}/R_{\text{out}} \sim 6 \times 10^{-3} M_{\odot} \text{ yr}^{-1}$ (Bischetti et al. 2017) for $R_{\text{out}} = 6.52 \pm 0.43$ kpc (see Fig. 2). We report all of these quantities in Tab. 2 along with the corresponding \dot{t}_{dyn} , \dot{E}_{kin} and \dot{P}_{out} , to allow a direct comparison with the values holding for the outflow.

7. I20210 SMBH mass estimates

In order to evaluate the SMBH mass M_{BH} in both objects, we use the approach detailed in BM19 for Type II AGN, in which obscuration prevents from detecting the broad components of permitted emission lines. In this case, the following single-epoch relation linking M_{BH} to a BLR virial shape factor ε that summarizes the uncertainties on the real BLR geometry, the monochromatic AGN luminosity $\lambda L_{\lambda}(5100 \text{ \AA})$ at 5100 Å and the broad H α FWHM holds:

$$\log\left(\frac{M_{\text{BH}}}{M_{\odot}}\right) = \log \varepsilon + 6.90 + 0.54 \cdot \log\left[\frac{\lambda L_{\lambda}(5100 \text{ \AA})}{10^{44} \text{ erg s}^{-1}}\right] + 2.06 \cdot \log\left[\frac{\text{FWHM}_{\text{H}\alpha}^{(\text{BLR})}}{10^3 \text{ km s}^{-1}}\right] \quad (14)$$

The validity of this relation holds as long $\text{FWHM}_{\text{H}\alpha}^{(\text{BLR})}$ and the [O III]/H β ratio are measured for AGN-dominated systems, and are therefore related by the following logarithmic linear relation:

$$\log\left(\frac{[\text{O III}]}{\text{H}\beta}\right) = (0.58 \pm 0.07) \cdot \log\left[\frac{\text{FWHM}_{\text{H}\alpha}^{(\text{BLR})}}{\text{km s}^{-1}}\right] - (1.38 \pm 0.38) \quad (15)$$

According to figures 3 and 4 of BM19, this happens for AGN-dominated systems with [O III]/H $\beta \gtrsim 0.55$ in which the line intensities are not contaminated by star formation in the host galaxy; since from Tab. 1 we have $\log([\text{O III}]/\text{H}\beta) \sim 0.7$ for I20210N and ~ 0.8 for I20210S respectively, we can apply Eqs. 14 and 15 to both I20210 members. We do not quote any errors for the following estimations of physical quantities involved in the determination of M_{BH} , since the measurement method is indirect and is therefore subject to uncertainties of at least ~ 0.5 dex (see BM19 and refs. therein).

For I20210N, we derive $\lambda L_{\lambda}(5100 \text{ \AA})$ using its absorption-corrected hard X-ray luminosity $L_{2-10 \text{ keV}} = 4.7 \times 10^{42}$ erg s $^{-1}$ (Piconcelli et al. 2010) through equation 5 by Maiolino et al. (2007):

$$\log L_{2-10 \text{ keV}} = 0.721 \cdot \log[\lambda L_{\lambda}(5100 \text{ \AA})] + 11.78 \quad (16)$$

This yields $\lambda L_{\lambda}(5100 \text{ \AA}) \simeq 7.0 \times 10^{42}$ erg s $^{-1}$. The FWHM of the invisible H α broad component is estimated from Eq. 15; in this way, we find $\text{FWHM}_{\text{H}\alpha}^{(\text{BLR})} \simeq 3650$ km s $^{-1}$, corresponding to $M_{\text{BH}} \simeq 2.9 \times 10^7 M_{\odot}$ if the value of $\varepsilon = 1.075$ by Reines & Volonteri (2015), representative of the mean virial factor $\langle f \rangle = 4\varepsilon = 4.3$ derived by Grier et al. (2013) measuring the stellar velocity dispersion in the host galaxies of powerful nearby quasars, is assumed.

For I20210S, we decide to do not directly obtain $\lambda L_{\lambda}(5100 \text{ \AA})$ from the observed spectrum due to the uncertainty on its intrinsic reddening; similarly, we do not apply Eq. 16 to indirectly compute it, since only a lower limit to $L_{2-10 \text{ keV}} \gtrsim$

Quantity	I20210N	I20210S	Units
$L_{\text{H}\beta}^{(\text{NLR})}$	$(1.622 \pm 0.075) \times 10^{40}$	$(1.040 \pm 0.072) \times 10^{42}$	erg s $^{-1}$
L_{bol}	5.2×10^{43}	3.6×10^{45}	erg s $^{-1}$
$\lambda L_{\lambda}(5100 \text{ \AA})$	7.0×10^{42}	2.7×10^{44}	erg s $^{-1}$
$\text{FWHM}_{\text{H}\alpha}^{(\text{BLR})}$	3650	5700	km s $^{-1}$
M_{BH}	2.9×10^7	5.2×10^8	M_{\odot}
λ_{Edd}	0.01	0.05	—
σ_v^*	390 ± 50	—	km s $^{-1}$
M_*	$\lesssim 1.5 \times 10^{12}$	—	M_{\odot}

Table 3. Physical parameters of the AGN hosted in the I20210 members. Quantities derived from the application of proportionality relations (Eq. 14 to Eq. 18) are reported without errors due to the uncertainties of at least ~ 0.5 dex affecting them (see BM19 and refs. therein).

5×10^{43} erg s $^{-1}$ is reported in Piconcelli et al. (2010). Instead, we rely upon equation 6 by BM19:

$$\log L_{\text{bol}} = \log L_{\text{H}\beta}^{(\text{NLR})} + 3.48 + \max\{0, 0.31 \cdot [\log([\text{O III}]/\text{H}\beta) - 0.6]\} \quad (17)$$

and equation 6 by Netzer (2009):

$$\log[\lambda L_{\lambda}(5100 \text{ \AA})] = 1.09 \cdot \log L_{\text{bol}} - 5.23 \quad (18)$$

From an I20210S $L_{\text{H}\beta}^{(\text{NLR})} = (1.040 \pm 0.072) \times 10^{42}$ erg s $^{-1}$ computed from the H β flux listed in Tab. 1, we are thus able to derive $L_{\text{bol}} \simeq 3.6 \times 10^{45}$ erg s $^{-1}$ and $\lambda L_{\lambda}(5100 \text{ \AA}) \simeq 2.7 \times 10^{44}$ erg s $^{-1}$. The L_{bol} estimated in this way is fully compatible with the value of $\sim 3 \times 10^{45}$ erg s $^{-1}$ assumed by Piconcelli et al. (2010) on the basis of the I20210S infrared luminosity in the range $10 \div 100$ μm (Sargsyan et al. 2011). Finally, Eq. 15 yields $\text{FWHM}_{\text{H}\alpha}^{(\text{BLR})} \simeq 5700$ km s $^{-1}$, which translates into $M_{\text{BH}} \simeq 5.2 \times 10^8 M_{\odot}$.

These SMBH masses imply an Eddington luminosity $L_{\text{Edd}} \sim 3.7 \times 10^{45}$ erg s $^{-1}$ for I20210N and $\sim 6.6 \times 10^{46}$ erg s $^{-1}$ for I20210S, respectively. Combining them with the AGN bolometric luminosities of $\sim 5.2 \times 10^{43}$ erg s $^{-1}$ for I20210N and $\sim 3.6 \times 10^{45}$ for I20210S, we obtain Eddington ratios $\lambda_{\text{Edd}} = L_{\text{bol}}/L_{\text{Edd}} \sim 0.01$ for I20210N and ~ 0.05 for I20210S. Though affected by large uncertainties, such ratios are consistent with the galaxy classification from the BPT diagrams, with I20210S clearly falling in the Seyfert region and I20210N exhibiting more mixed properties. In addition, we combine the equation system 9 by BM19 to infer the the stellar mass M_* of the I20210N host galaxy from its σ_v^* of 390 km s $^{-1}$ derived in Sect. 3.1, finding $M_* \sim 1.5 \times 10^{12} M_{\odot}$. We summarize all these parameters in Tab. 3.

The value of M_* derived for I20210N is ~ 50 times larger than that typically expected from the $M_{\text{bulge-to-}M_{\text{BH}}}$ relation ($M_{\text{bulge}} \simeq 10^3 M_{\text{BH}}$; e.g., Magorrian et al. 1998; Häring & Rix 2004; Gültekin et al. 2009); however, we highlight that the I20210N kinematics is likely altered by its gravitational interaction with I20210S, thus making its measured σ_v^* unreliable for the purposes of estimating the stellar mass. Therefore, the value of M_* derived in this way should only be treated as an (overestimated) upper limit to the I20210N baryonic content.

8. Summary and conclusions

In this article, we have presented the optical spectroscopic analysis of the AGN pair hosted in the interacting system IRAS 20210+1121 (I20210; P90; Heisler & Vader 1995; Burston et al.

2001; Davies et al. 2002; Arribas et al. 2004; Piconcelli et al. 2010) at $z = 0.056$. This study is based on spectroscopy through a slit aligned along the nuclei of the two interacting galaxies. The high-quality data taken at the *Telescopio Nazionale Galileo* allowed us to perform a detailed study of the light emission from both components and from their surrounding environment in the rest-frame wavelength range 3500 – 7300 Å, with the possibility of a comprehensive characterization of this interacting galaxy pair. Here we summarize our main findings:

- I20210N, the Northern member of the I20210 system, can be definitely classified as a Seyfert 2 galaxy with an exceptional stellar velocity dispersion $\sigma_v^* \sim 400 \text{ km s}^{-1}$, hosting an AGN powered by a black hole with $M_{\text{BH}} \sim 3 \times 10^7 M_{\odot}$ that radiates at 1% of its Eddington limit;
- the environment around I20210S, the Southern component – a powerful Type II quasar with $M_{\text{BH}} \sim 5 \times 10^8 M_{\odot}$ radiating at 5% of its Eddington limit – reveals to be highly structured, with an ionized outflow and a detached gaseous nebula (the South Nebula) alongside the nuclear emission;
- the physical properties of the ionized outflow derived from the analysis of the broad emission-line components ($T_e \sim 10^4 \text{ K}$, $\langle n_e \rangle \gtrsim 5000 \text{ cm}^{-3}$, $v_{\text{max}} \sim 2000 \text{ km s}^{-1}$, $R_{\text{out}} \sim 2 \text{ kpc}$, $M_{\text{out}} \sim 2 \times 10^5 M_{\odot}$, $\dot{M} \sim 0.6 M_{\odot} \text{ yr}^{-1}$), in line with those found in other powerful AGN hosted in ULIRGs (Rodríguez Zaurín et al. 2013; Rupke & Veilleux 2013; Kakkad et al. 2018; Spence et al. 2018), suggest that the I20210S AGN activity has potentially a direct impact on the host-galaxy environment through quasar feedback; however, they need to be further investigated with higher-resolution spectral observations, in order to constrain the value of the wind electron density and thus allow to better characterize the feedback effect to the I20210S star formation activity (e.g., Carniani et al. 2015; Fiore et al. 2017);
- the South Nebula exhibits dynamical properties consistent with those of highly-disrupted gas stripped out of the I20210S nucleus (velocity blueshift of $\sim 500 \text{ km s}^{-1}$, FWHM of $\sim 700 \text{ km s}^{-1}$ – similar to the case of the Teacup Galaxy (Ramos Almeida et al. 2017) – coupled to intermediate ionization properties between AGN-powered and star-forming gas. Such features qualify this region as a very interesting target for a deeper investigation of the potential feedback processes – either triggered by AGN activity or by the galaxy merger – at work in I20210S.

Thanks to the above properties, the I20210 system characterizes as a very interesting target in the local Universe to be investigated with dedicated multi-wavelength follow-ups aimed at studying in detail the effects of AGN feedback coupled to host-galaxy interaction on the AGN surrounding environment. In particular, higher-resolution spectra ($\lambda/\Delta\lambda \gtrsim 1500$) are crucial to obtain in order to improve the emission-line diagnostics of the I20210S components (nucleus, outflow, South Nebula) and to allow a precise evaluation of the I20210S outflow physical conditions. Furthermore, integral-field spectroscopic observations are required to accurately constrain both morphology and interplay of outflow and off-nuclear emitting region in I20210S.

Acknowledgements. We thank our anonymous referee for their helpful comments. GV, EP, CV, MB, CF and FF acknowledge support from PRIN MIUR project “Black Hole winds and the Baryon Life Cycle of Galaxies: the stone-guest at the galaxy evolution supper”, contract #2017PH3WAT. GV also acknowledges financial support from Premiale 2015 MITic (PI: B. Garilli). CRA acknowledges financial support from the Spanish Ministry of Science, Innovation and Universities (MCIU) under grant with reference RYC-2014-15779, from the European Union’s Horizon 2020 research and innovation programme under Marie Skłodowska-Curie grant agreement No 860744 (Bid4BEST), from the

State Research Agency (AEI-MCINN) of the Spanish MCIU under grants “Feeding and feedback in active galaxies” with reference PID2019-106027GB-C42, “Feeding, feedback and obscuration in active galaxies” with reference AYA2016-76682-C3-2-P, and “Quantifying the impact of quasar feedback on galaxy evolution (QSOFEED)” with reference EUR2020-112266. CRA also acknowledges support from the Consejería de Economía, Conocimiento y Empleo del Gobierno de Canarias and the European Regional Development Fund (ERDF) under grant with reference ProID2020010105 and from IAC project P/301404, financed by the Ministry of Science and Innovation, through the State Budget and by the Canary Islands Department of Economy, Knowledge and Employment, through the Regional Budget of the Autonomous Community. Based on observations made with the Italian *Telescopio Nazionale Galileo* (TNG) operated on the island of La Palma by the Fundación Galileo Galilei of the INAF (Istituto Nazionale di Astrofisica) at the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofísica de Canarias. IRAF is distributed by the National Optical Astronomy Observatories, which is operated by the Association of Universities for Research in Astronomy, Inc. (AURA) under cooperative agreement with the National Science Foundation. Reproduced with permission from Astronomy & Astrophysics, © ESO.

References

- Arribas, S., Bushouse, H., Lucas, R. A., Colina, L., & Borne, K. D. 2004, *AJ*, 127, 2522
- Arribas, S., Colina, L., Bellocchi, E., Maiolino, R., & Villar-Martín, M. 2014, *A&A*, 568, A14
- Baldwin, J. A., Phillips, M. M., & Terlevich, R. 1981, *PASP*, 93, 5
- Baron, D. & Ménard, B. 2019, *MNRAS*, 487, 3404
- Bellocchi, E., Arribas, S., Colina, L., & Miralles-Caballero, D. 2013, *A&A*, 557, A59
- Bischetti, M., Piconcelli, E., Vietri, G., et al. 2017, *A&A*, 598, A122
- Bisogni, S., Marconi, A., & Risaliti, G. 2017, *MNRAS*, 464, 385
- Brusa, M., Bongiorno, A., Cresci, G., et al. 2015, *MNRAS*, 446, 2394
- Burston, A. J., Ward, M. J., & Davies, R. I. 2001, *MNRAS*, 326, 403
- Cano-Díaz, M., Maiolino, R., Marconi, A., et al. 2012, *A&A*, 537, L8
- Cappellari, M. 2012, pPXF: Penalized Pixel-Fitting stellar kinematics extraction
- Cappellari, M. 2017, *MNRAS*, 466, 798
- Cappellari, M. & Emsellem, E. 2004, *PASP*, 116, 138
- Carniani, S., Marconi, A., Maiolino, R., et al. 2015, *A&A*, 580, A102
- Cattaneo, A., Faber, S. M., Binney, J., et al. 2009, *Nature*, 460, 213
- Chambers, K. C. & Pan-STARRS Team. 2016, in *American Astronomical Society Meeting Abstracts*, Vol. 227, American Astronomical Society Meeting Abstracts #227, 324.07
- Cicone, C., Brusa, M., Ramos Almeida, C., et al. 2018, *Nature Astronomy*, 2, 176
- Cid Fernandes, R., Stasińska, G., Schlickmann, M. S., et al. 2010, *MNRAS*, 403, 1036
- Cook, R. D. & Weisberg, S. 1982, *Residuals and Influence in Regression*
- Cresci, G., Mainieri, V., Brusa, M., et al. 2015, *ApJ*, 799, 82
- Croton, D. J., Springel, V., White, S. D. M., et al. 2006, *MNRAS*, 365, 11
- Davies, R., Baron, D., Shimizu, T., et al. 2020, *MNRAS*, 498, 4150
- Davies, R. I., Burston, A., & Ward, M. J. 2002, *MNRAS*, 329, 367
- Davis, M., Efstathiou, G., Frenk, C. S., & White, S. D. M. 1985, *ApJ*, 292, 371
- De Rosa, A., Vignali, C., Husemann, B., et al. 2018, *MNRAS*, 480, 1639
- Di Matteo, T., Springel, V., & Hernquist, L. 2005, *Nature*, 433, 604
- Doyon, R., Wells, M., Wright, G. S., et al. 1994, *ApJ*, 437, L23
- Ellison, S. L., Viswanathan, A., Patton, D. R., et al. 2019, *MNRAS*, 487, 2491
- Elvis, M. 2000, *ApJ*, 545, 63
- Fabian, A. C. 2012, *ARA&A*, 50, 455
- Falcón-Barroso, J., Lyubenova, M., van de Ven, G., et al. 2017, *A&A*, 597, A48
- Farrah, D., Bernard-Salas, J., Spoon, H. W. W., et al. 2007, *ApJ*, 667, 149
- Feruglio, C., Fiore, F., Carniani, S., et al. 2015, *A&A*, 583, A99
- Fiore, F., Feruglio, C., Shankar, F., et al. 2017, *A&A*, 601, A143
- Forbes, D. A. & Ponman, T. J. 1999, *MNRAS*, 309, 623
- Gaskell, C. M. & Ferland, G. J. 1984, *PASP*, 96, 393
- Grier, C. J., Martini, P., Watson, L. C., et al. 2013, *ApJ*, 773, 90
- Gültekin, K., Richstone, D. O., Gebhardt, K., et al. 2009, *ApJ*, 698, 198
- Häring, N. & Rix, H.-W. 2004, *ApJ*, 604, L89
- Heckman, T. M. 1980, *A&A*, 87, 152
- Heckman, T. M. & Best, P. N. 2014, *ARA&A*, 52, 589
- Heckman, T. M., Lehnert, M. D., Strickland, D. K., & Armus, L. 2000, *ApJS*, 129, 493
- Heisler, C. A. & Vader, J. P. 1995, *AJ*, 110, 87
- Hernquist, L. 1989, *Nature*, 340, 687
- Ho, L. C., Filippenko, A. V., & Sargent, W. L. 1995, *ApJS*, 98, 477
- Ho, L. C., Filippenko, A. V., & Sargent, W. L. W. 1993, *ApJ*, 417, 63
- Hopkins, P. F. & Elvis, M. 2010, *MNRAS*, 401, 7
- Hopkins, P. F., Hernquist, L., Cox, T. J., et al. 2006, *ApJS*, 163, 1

- Husemann, B., Wisotzki, L., Sánchez, S. F., & Jahnke, K. 2013, *A&A*, 549, A43
- Imanishi, M. & Saito, Y. 2014, *ApJ*, 780, 106
- Jones, D. H., Read, M. A., Saunders, W., et al. 2009, *MNRAS*, 399, 683
- Kakkad, D., Groves, B., Dopita, M., et al. 2018, *A&A*, 618, A6
- Kakkad, D., Mainieri, V., Padovani, P., et al. 2016, *A&A*, 592, A148
- Kauffmann, G., Heckman, T. M., Tremonti, C., et al. 2003, *MNRAS*, 346, 1055
- Kennicutt, Robert C., J. 1998, *ApJ*, 498, 541
- Kennicutt, Robert C., J. & De Los Reyes, M. A. C. 2021, *ApJ*, 908, 61
- Kennicutt, R. C. & Evans, N. J. 2012, *ARA&A*, 50, 531
- Kewley, L. J., Dopita, M. A., Leitherer, C., et al. 2013a, *ApJ*, 774, 100
- Kewley, L. J., Dopita, M. A., Sutherland, R. S., Heisler, C. A., & Trevena, J. 2001, *ApJ*, 556, 121
- Kewley, L. J., Groves, B., Kauffmann, G., & Heckman, T. 2006, *MNRAS*, 372, 961
- Kewley, L. J., Maier, C., Yabe, K., et al. 2013b, *ApJ*, 774, L10
- Kormendy, J. & Ho, L. C. 2013, *ARA&A*, 51, 511
- Lamareille, F. 2010, *A&A*, 509, A53
- Lynden-Bell, D. 1969, *Nature*, 223, 690
- Magorrian, J., Tremaine, S., Richstone, D., et al. 1998, *AJ*, 115, 2285
- Maiolino, R., Russell, H. R., Fabian, A. C., et al. 2017, *Nature*, 544, 202
- Maiolino, R., Shemmer, O., Imanishi, M., et al. 2007, *A&A*, 468, 979
- Markwardt, C. B. 2009, in *Astronomical Society of the Pacific Conference Series*, Vol. 411, *Astronomical Data Analysis Software and Systems XVIII*, ed. D. A. Bohlender, D. Durand, & P. Dowler, 251
- Miller, J. S. & Mathews, W. G. 1972, *ApJ*, 172, 593
- Netzer, H. 2009, *MNRAS*, 399, 1907
- Oh, K., Sarzi, M., Schawinski, K., & Yi, S. K. 2011, *ApJS*, 195, 13
- Oh, K., Yi, S. K., Schawinski, K., et al. 2015, *ApJS*, 219, 1
- Osterbrock, D. E. 1981, *ApJ*, 249, 462
- Osterbrock, D. E. & Ferland, G. J. 2006, *Astrophysics of gaseous nebulae and active galactic nuclei*
- Pei, Y. C. 1992, *ApJ*, 395, 130
- Perez, E., Manchado, A., Garcia-Lario, P., & Pottasch, S. R. 1990, *A&A*, 227, 407
- Perna, M., Arribas, S., Catalán-Torrecilla, C., et al. 2020, *A&A*, 643, A139
- Perna, M., Arribas, S., Pereira Santaella, M., et al. 2021, *A&A*, 646, A101
- Perna, M., Brusa, M., Cresci, G., et al. 2015, *A&A*, 574, A82
- Perna, M., Lanzuisi, G., Brusa, M., Cresci, G., & Mignoli, M. 2017, *A&A*, 606, A96
- Piconcelli, E., Vignali, C., Bianchi, S., et al. 2010, *ApJ*, 722, L147
- Ramos Almeida, C., Piqueras López, J., Villar-Martín, M., & Bessiere, P. S. 2017, *MNRAS*, 470, 964
- Reines, A. E. & Volonteri, M. 2015, *ApJ*, 813, 82
- Reyes, R., Zakamska, N. L., Strauss, M. A., et al. 2008, *AJ*, 136, 2373
- Rodríguez Zaurín, J., Tadhunter, C. N., Rose, M., & Holt, J. 2013, *MNRAS*, 432, 138
- Rosales-Ortega, F. F., Arribas, S., & Colina, L. 2012, *A&A*, 539, A73
- Rose, M., Tadhunter, C., Ramos Almeida, C., et al. 2018, *MNRAS*, 474, 128
- Rupke, D. S., Veilleux, S., & Sanders, D. B. 2002, *ApJ*, 570, 588
- Rupke, D. S., Veilleux, S., & Sanders, D. B. 2005, *ApJS*, 160, 115
- Rupke, D. S. N. & Veilleux, S. 2013, *ApJ*, 768, 75
- Sanders, D. B. & Mirabel, I. F. 1996, *ARA&A*, 34, 749
- Sanders, D. B., Soifer, B. T., Elias, J. H., et al. 1988, *ApJ*, 325, 74
- Sanders, R. L., Shapley, A. E., Kriek, M., et al. 2016, *ApJ*, 816, 23
- Sargsyan, L., Weedman, D., Leboutteiller, V., et al. 2011, *ApJ*, 730, 19
- Scannapieco, E. & Oh, S. P. 2004, *ApJ*, 608, 62
- Spence, R. A. W., Tadhunter, C. N., Rose, M., & Rodríguez Zaurín, J. 2018, *MNRAS*, 478, 2438
- Springel, V., White, S. D. M., Jenkins, A., et al. 2005, *Nature*, 435, 629
- Vazdekis, A., Sánchez-Blázquez, P., Falcón-Barroso, J., et al. 2010, *MNRAS*, 404, 1639
- Veilleux, S., Cecil, G., & Bland-Hawthorn, J. 2005, *ARA&A*, 43, 769
- Veilleux, S., Kim, D. C., & Sanders, D. B. 2002, *ApJS*, 143, 315
- Veilleux, S. & Osterbrock, D. E. 1987, *ApJS*, 63, 295
- Weedman, D. W. 1968, *PASP*, 80, 314
- Zakamska, N. L., Hamann, F., Pâris, I., et al. 2016, *MNRAS*, 459, 3144