

## An extended Lyman $\alpha$ outflow from a radio galaxy at $z=3.7$ ?

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### ABSTRACT

Spatially resolved observations of AGN host galaxies undergoing feedback processes are one of the most relevant avenues through which galactic evolution can be studied, given the long lasting effects AGN feedback has on gas reservoirs, star formation, and AGN environments at all scales. Within this context we report results from VLT/MUSE integral field optical spectroscopy of TN J1049-1258, one of the most powerful radio sources known, at a redshift of 3.7. We detected extended ( $\sim 18$  kpc) Ly $\alpha$  emission, spatially aligned with the radio axis, redshifted by  $2250 \pm 60$  km s<sup>-1</sup> with respect to the host galaxy systemic velocity, and co-spatial with UV continuum emission. This Ly $\alpha$  emission could arise from a companion galaxy, although there are arguments against this interpretation. Alternatively, it might correspond to an outflow of ionized gas stemming from the radio galaxy. The outflow would be the highest redshift spatially resolved ionized outflow to date. The enormous amount of energy injected, however, appears to be unable to quench the host galaxy's prodigious star formation, occurring at a rate of  $\sim 4500$  M<sub>⊙</sub>yr<sup>-1</sup>, estimated using its far infra-red luminosity. Within the field we also found two companion galaxies at projected distances of  $\sim 25$  kpc and  $\sim 60$  kpc from the host, which suggests the host galaxy is harbored within a protocluster.

### 1. INTRODUCTION

Outflows are believed to be the main driver through which active galactic nuclei (AGN) reg-

ulate the evolution of their host galaxies and affect the large scale environment (AGN feedback), operating mostly at the epoch of the peak of star formation between redshift 2 and 3 (Hopkins 2006; Miley & De Breuck 2008). Outflows, extended on kiloparsec scales, powered by the AGN radiative pressure (quasar feedback mode) and/or radio jets (radio feedback mode), have been resolved in ionised, atomic, and molecular gas both in local (e.g. Rupke & Veilleux 2011, Rupke & Veilleux 2013, Cicone et al. 2012, Rodriguez Zaurin et al. 2013, Speranza et al. 2021, Ramos Almeida et al. 2022) and in high redshift (radio loud and radio quiet) quasars (e.g. Alexander et al. 2010; Nesvadba et al. 2010; Maiolino et al. 2012; Cano-Díaz et al. 2012; Cresci et al. 2015). Characterizing and constraining the effects and extent to which AGN feedback affects star formation is one of the most important problems in modern astrophysics, as it has ties to the biggest questions in galaxy evolution and cosmology.

In this context, high redshift radio galaxies (HzRGs) are particularly relevant as the nuclear regions are obscured along our line of sight due to a circumnuclear dusty structure (e.g., Antonucci 1993) or by the host galaxy itself. The lower nuclear continuum luminosities allow us to reduce the observational uncertainties produced by the subtraction of a prominent Point Spread Function (PSF), probing the host and the gas structure down to the limit of the spatial resolution of only a few kpc.

We started a comprehensive program of observations with the Multi Unit Spectroscopic Explorer (MUSE) at the VLT to explore the properties of the ionized gas in HzRGs. In this Letter we present the results obtained for TN J1049-1258, an HzRG at redshift  $3.697 \pm 0.004$  (Borncini et al. 2007), located at  $\alpha=10:49:06.2$ ,  $\delta=-12:58:19$  (J2000). The scale factor at this redshift (assuming an  $H_0=69.6$  and  $\Omega_m=0.286$  cosmology) is  $7.3 \text{ kpc arcsec}^{-1}$ . Its radio flux

measured at 74 MHz by the NRAO VLA Sky Survey is  $5.18 \pm 0.56 \text{ Jy}$  while its radio spectral index is  $\alpha = 1.4$  (Condon et al. 1998).<sup>1</sup> The jet power,  $P_{\text{jet}}$ , can be obtained using the correlation between  $P_{\text{jet}}$  and radio luminosity, based on the measurement of the mechanical power of radio sources producing cavities in the surrounding hot gas as proposed by (Cavagnolo et al. 2010). The radio luminosity of TN J1049-1258 at (rest frame) 327 MHz can be estimated as  $P_{327} = 2.7 \times 10^{44} \text{ erg s}^{-1}$  which yields a jet power of  $P_{\text{jet}} \sim 6 \times 10^{45} \text{ erg s}^{-1}$ . A similar value is obtained by using the relation between radio and jet power from Willott et al. (1999). In the Karl G. Jansky VLA Sky Survey (VLASS, Lacy et al. 2020), performed at 3 GHz with a resolution of  $2''5$ , its radio emission is dominated by two radio components separated by  $11''$  (80 kpc) oriented at the position angle  $-80^\circ$ , with the host located close to their midpoint. Its radio luminosity at the rest frame frequency of 500 MHz is  $8.7 \times 10^{28} \text{ W Hz}^{-1}$ , making TN J1049-1258 one of the most luminous radio sources known (see Miley & De Breuck 2008). The flux at  $3.5 \mu\text{m}$ , corresponding to the optical emission in the rest frame, measured by Wide-field Infrared Survey Explorer (WISE, Wright et al. 2010) of TN J1049-1258 indicates a luminosity of  $2 \times 10^{11} L_\odot$ , suggesting that its host is already a very massive galaxy.

## 2. OBSERVATIONS AND DATA ANALYSIS

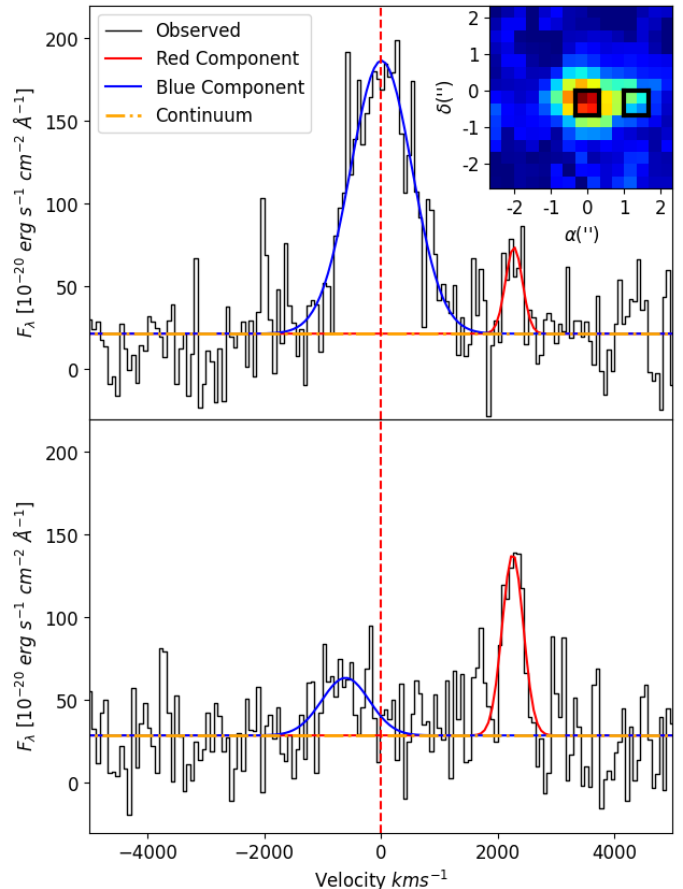
The observations were carried out with the Multi Unit Spectroscopic Explorer (MUSE) on the 29th of January of 2022 as part of the program ID:108.22FU. Four separate observations were performed, between which the telescope was rotated  $90^\circ$  to reject cosmic rays, for a total of 2800s of exposure time. The seeing measured from several point sources in the field of view was estimated as  $0''.71 \pm 0''.04$ . We used the ESO

<sup>1</sup> Spectral index,  $\alpha$ , is defined by flux density,  $S_\nu \propto \nu^{-\alpha}$ .

MUSE pipeline (version 2.8.7) to obtain a fully reduced and calibrated data cube (Weilbacher et al. 2020). We derived an absolute astrometric calibration of the MUSE data by cross-matching visible sources in both MUSE and Pan-STARRS (Panoramic Survey and Rapid Response System (Chambers et al. 2016)). Based on the scatter of the sources considered we estimate the derived astrometry to have an error of  $\sim 0''.2$ .

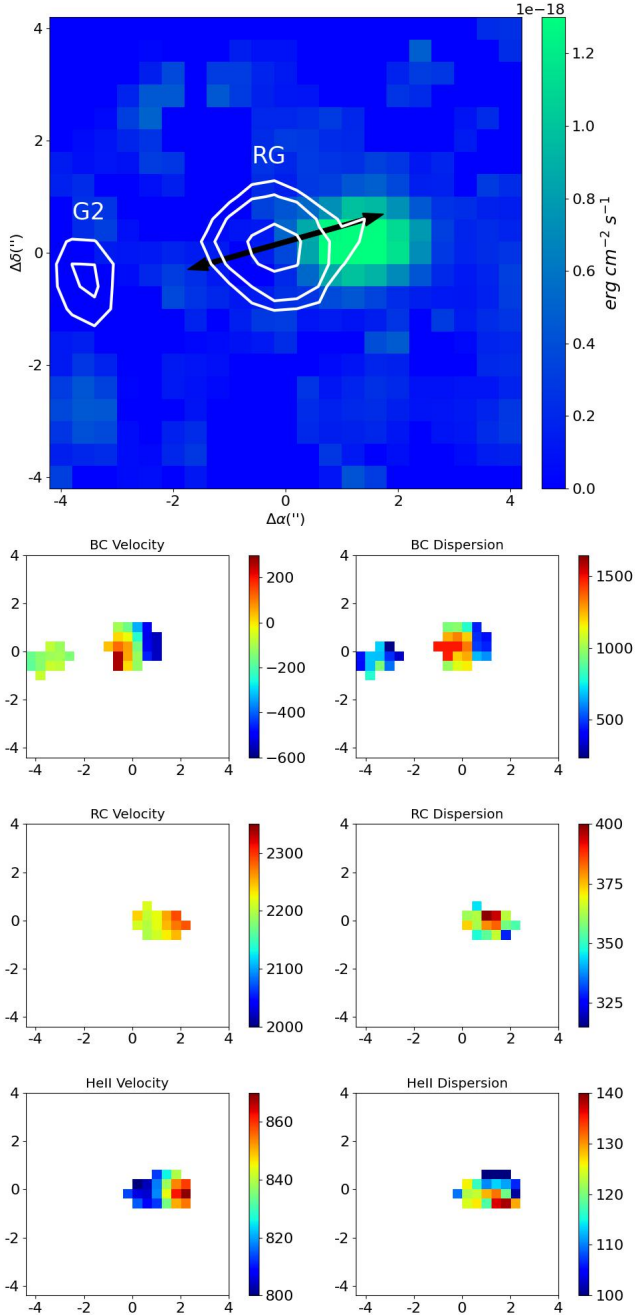
After first inspecting the nuclear spectrum on a large aperture around the peak of Ly $\alpha$  emission, we found two emission lines separated by about 50 Å. To fully characterize both the emission line morphology and the gas kinematics, we then performed a two Gaussian component fit in each spaxel (see Fig. 1 for the nuclear and off-nuclear spectra). Aiming to increase the signal to noise ratio per pixel, the cube was spatially rebinned into  $2 \times 2$  pixels, resulting in a  $0''.4 \times 0''.4$  area. The *modeling* tool from *Astropy*'s Python package (Astropy Collaboration et al. 2013, Astropy Collaboration et al. 2018, Price-Whelan et al. 2022) was used to model and fit both the continuum and emission lines. For the continuum we used a linear least squares fit, while the emission lines were modelled with a Gaussian using a Levenberg-Marquardt least squares fit.

From the fit we obtained the distribution of total flux and first and second moments for both components, while the uncertainties of the parameters were obtained through a Monte Carlo simulation. Each realization consists of adding flux variations to each spectral pixel, with the value of the variations taken at random from a normal distribution centered at the pixel noise. From each realization, a value for the relevant parameters (line center, amplitude, width) is obtained. After a set number of iterations (100 in our case) we obtain a value distribution for each of these parameters, from which we extract the mean (parameter value) and  $1\sigma$  standard deviation (parameter uncertainty). The pixel noise was calculated as the  $1\sigma$  standard deviation



**Figure 1. Top:** Nuclear spectrum of TN J1049-1258, extracted from a  $0''.8 \times 0''.8$  synthetic aperture, along with the modelled continuum and fitted lines. The spectrum shows the presence of two components, separated by  $\sim 2,250 \text{ km s}^{-1}$ . In the top right inset we show the total Ly $\alpha$  flux map, spatially rebinned into  $2 \times 2$  pixels, with the two synthetic apertures used to extract the nuclear and off-nuclear spectra superimposed. **Bottom:** Off-nuclear spectrum extracted  $1''.6$  west of the nucleus. The redshifted line becomes the brightest component. In both cases, the red dashed line marks the host galaxy's systemic velocity, corresponding to a redshift of 3.697. The X and Y scales are identical in both spectra.

of the flux in a 50 nm interval centered on Ly $\alpha$  from a section of empty sky of equal size as the cube extracted for the line analysis ( $80 \times 80$  pixels or  $16'' \times 16''$ ). All emission line maps have been obtained by setting a S/N threshold at 3.



**Figure 2.** **Top:** Ly $\alpha$  emission map of the red component, with superimposed iso-intensity contours of the Ly $\alpha$  from the host galaxy. The field of view is  $\sim 30 \text{ kpc} \times 30 \text{ kpc}$ . There is a clear offset between the center of the host galaxy emission and of the red component. A second source of Ly $\alpha$  emission (that we name G2) is seen  $\sim 3.5''$  east of the host galaxy. The arrow shows the radio axis. **Bottom:** maps of projected velocity and velocity dispersion of the host galaxy (top), of the red Ly $\alpha$  component (middle), and of the HeII line in  $\text{km s}^{-1}$ . In all cases, the S/N threshold on the line fit is set at  $3\sigma$ .

The continuum map was obtained using multiple intervals within MUSE’s wavelength range (475 nm to 935 nm, 101 nm to 198 nm rest frame) over a total range of 144 nm, selected to avoid both sky lines as well as object emission lines.

### 3. RESULTS

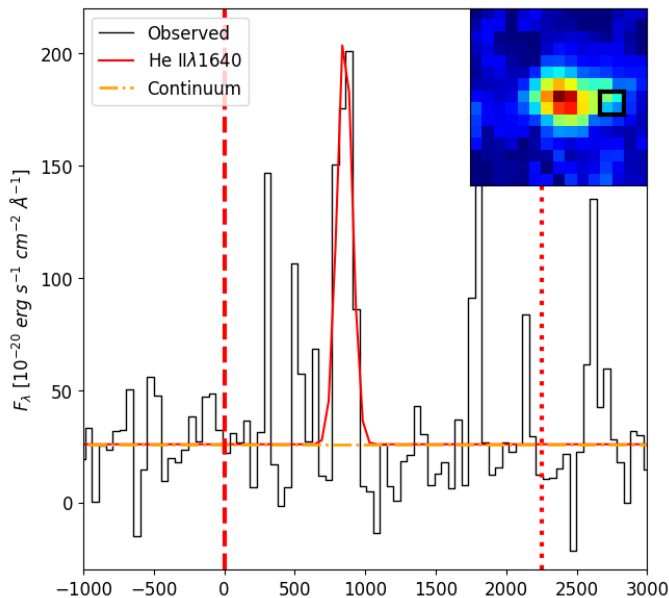
We estimated the redshift of TN J1049-1258 from the fit of the main component of Ly $\alpha$  line in the nuclear spectrum, which corresponded to  $z=3.697\pm 0.001$ , the same value obtained by [Bornancini et al. \(2007\)](#). However, it must be noted that the Ly $\alpha$  line peak has been observed to be redshifted from the systemic velocity (measured by the H $\alpha$  emission) by  $\sim 450 \text{ km s}^{-1}$  on average, and by as much as  $\sim 900 \text{ km s}^{-1}$  (e.g., [Steidel et al. 2010](#)).

The emission maps of both components and all relevant velocity and velocity dispersion maps are shown in Fig. 2. The velocity field estimated from the host galaxy’s Ly $\alpha$  emission presents a large gradient, with a velocity range of  $\sim 900 \text{ km s}^{-1}$  and a velocity dispersion ranging from 500 to  $1500 \text{ km s}^{-1}$ , that can be ascribed to either rotation or radial motion. While there is a clear spatial correlation between width and relative velocity, interpreting the absolute values of the velocity dispersion is made difficult by the resonant character of Ly $\alpha$  emission. Neutral hydrogen in the line of sight of the emitting gas will result in peaks and troughs, deeply modifying the original line profile depending on the density, distribution, and velocity of the neutral H gas.

The redshifted component, has an elongated structure extending  $\sim 2.5''$ , i.e.,  $\sim 18.5 \text{ kpc}$ , measured as the distance between the two farthest pixels with  $S/N > 3$ , along the same position angle of the radio structure. Its relative velocity, measured using the line profile in the off-nuclear spectrum (see Figure 1) is  $v = 2250 \pm 60 \text{ km s}^{-1}$  and it remains remarkably similar throughout the spatial extent of the

emission zone. In a similar fashion, the velocity dispersion values are within a range of  $\sim 70$   $\text{km s}^{-1}$ , with this red component being much narrower than the one in the host galaxy; its FWHM measured from the off-nuclear spectrum is  $\sigma_{\text{RC}} = 430 \pm 20$   $\text{km s}^{-1}$ .

In the region cospatial to the  $\text{Ly}\alpha$  red component we also detected emission from the  $\text{He II}\lambda 1640$  line (see Fig. 3). This line is blueshifted by  $1404 \pm 9$   $\text{km s}^{-1}$  (value obtained from the off-nuclear spectrum shown in Fig 1) with respect to the  $\text{Ly}\alpha$  emission. Its velocity and velocity dispersion maps are shown in Figure 2. From this spectrum we obtain a line ratio  $\text{Ly}\alpha/\text{He II}=1.5 \pm 0.3$ , which is at the very low end of the values measured in HzRGs (see Villar-Martín et al. 2007).



**Figure 3.** He II spectrum extracted using the same off-nuclear aperture as  $\text{Ly}\alpha$ . The dashed line marks the host systemic velocity, and the dotted line marks the extended emission velocity.

UV continuum emission, within the rest frame range  $\sim 1,000 - 2,000$   $\text{\AA}$ , (see Fig. 4) is observed on the west side of the host galaxy, cospatial with the  $\text{Ly}\alpha$  red component. The source observed  $\sim 5''$  south of the radio galaxy (RG

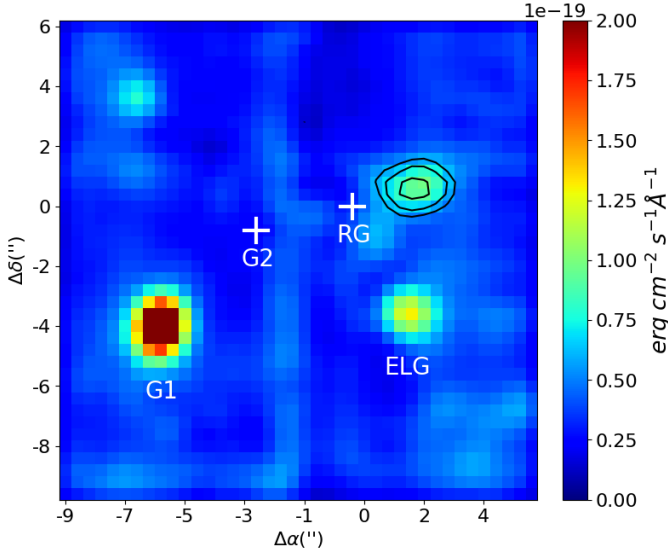
from hereon) is a galaxy at  $z=0.87$ , based on the identification of the  $[\text{O II}]\lambda\lambda 3726, 3729$  doublet as well as  $\text{H}\delta$  and  $\text{H}\gamma$  in its spectrum. The brightest UV source in the field is another potential companion we named G1, located  $\sim 8''$  ( $\sim 60$  kpc) southeast of the host galaxy. Its  $\text{Ly}\alpha$  emission was detected at a very similar redshift (see Fig. 5, top panel) and its high redshift identification is also supported by it being a g-dropout based on the photometry from Pan-STARRS (Chambers et al. 2016).<sup>2</sup>

To the east of the radio galaxy, another patch of  $\text{Ly}\alpha$  emission can be seen in the top panel of Fig. 2. This is associated with a companion galaxy we named G2, whose identification is possible thanks to the presence of  $\text{Ly}\alpha$  (see Fig. 5, bottom panel),  $\text{He II}\lambda 1640$ , and  $\text{CIV}\lambda\lambda 1548, 1550$  in its spectrum. Its redshift,  $z=3.685 \pm 0.002$ , was estimated from the  $\text{He II}\lambda 1640$  emission, which is a non resonant line. The presence of G1 and G2 suggests TN J1049-1258 may be harboured within a proto-cluster.

The main result of these observations is the presence of extended  $\text{Ly}\alpha$  emission redshifted by  $\sim 2,250$   $\text{km s}^{-1}$  and we can envisage two possibilities for its origin:

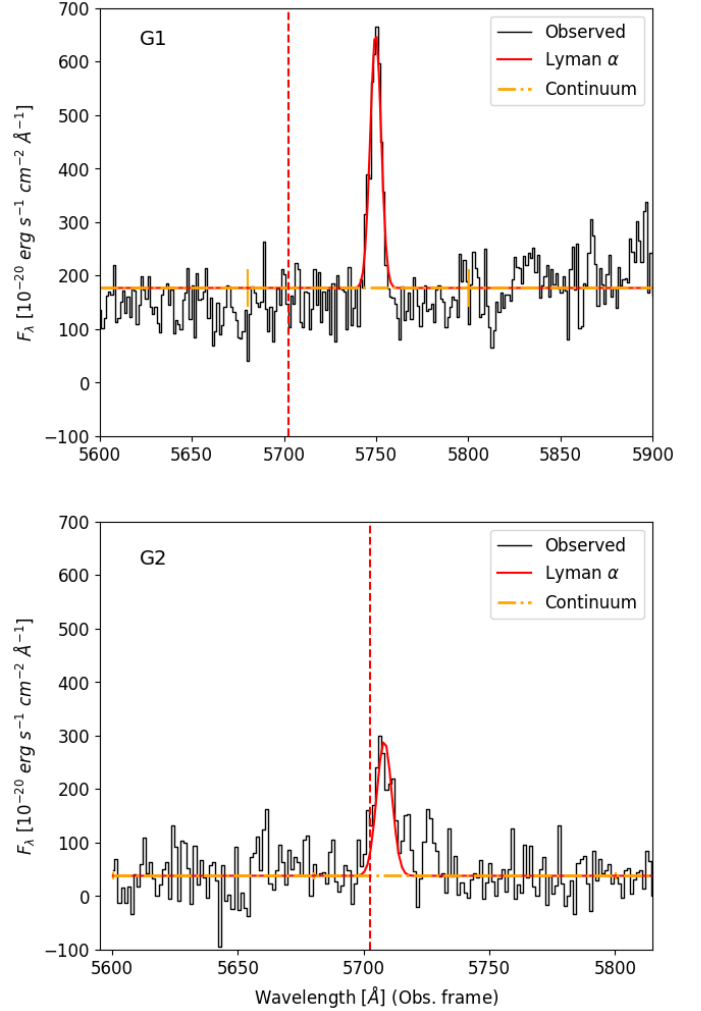
**1) A companion galaxy.** The presence of UV continuum emission co-spatial with the  $\text{Ly}\alpha$  region suggests that this might be associated with another galaxy at a similar redshift. This interpretation is also supported by the low velocity dispersion ( $\sigma = 430 \pm 20$   $\text{km s}^{-1}$ ) of the  $\text{Ly}\alpha$  emission. The  $\text{Ly}\alpha$  equivalent width (EW) of  $6.6 \pm 2.0$   $\text{\AA}$  does not provide useful information as to the origin of the line emission since the EW of Lyman Alpha Emitters spans a considerably wide range, with measurements ranging from a few  $\text{\AA}$  to several hundreds (see Kerutt

<sup>2</sup> The Pan-STARRS aperture magnitudes for G1 are  $g=24.08 \pm 0.02$ ,  $r=22.44 \pm 0.01$ , and  $i=22.39 \pm 0.01$ , respectively, then satisfying the g-dropout criteria by Pouliaxis et al. (2022).



**Figure 4.** Rest frame UV continuum flux density map of the field around TN J1049-1258. A source of continuum can be seen to the west of the Ly $\alpha$  emission center, cospatial with the red component (black contours at 0.5, 0.8 and 1.3  $10^{-18}$   $\text{erg cm}^{-2} \text{s}^{-1}$ ). The brightest UV source, G1, is located  $\sim 8''$  ( $\sim 60$  kpc) southeast of the host galaxy. At the location of G2 (marked with a "+" sign) an image defect (the vertical line at  $\Delta\alpha \sim -1''$ ) does not allow us to obtain a robust UV counterpart to the Ly $\alpha$  emission seen in Fig.2. Another source found  $5''$  to the south and slightly west is identified as a  $z=0.87$  emission line galaxy and labelled ELG

et al. 2022). However, studies of high redshift galaxies found that the typical radius of a  $z \sim 4$  galaxy is  $0''.2$ , that is, 1.5 kpc (Ferguson et al. 2004; Bouwens et al. 2004). The full width at half maximum of the Ly $\alpha$  red component emission is significantly larger: by fitting its spatial distribution with a Gaussian we obtain a size of  $1''.79 \pm 0''.14$ , equivalent to a deconvolved extension of  $R=12.0 \pm 0.9$  kpc. We must note, however, that G1 is also extended, with a deconvolved size of  $7.3 \pm 0.5$  kpc. The strongest argument against this interpretation is the blueshift of  $\sim 1,400$   $\text{km s}^{-1}$  which the He II $\lambda 1640$  line presents with respect to the extended Ly $\alpha$  line,



**Figure 5.** Spectra in the Ly $\alpha$  region for the two companion galaxies, G1 (top) and G2 (bottom). The red dashed line marks the RG's redshift measured with Ly $\alpha$ .

as it greatly exceeds the range of velocity offsets between Ly $\alpha$  and nebular lines observed in high redshift sources (Steidel et al. 2010), arguing in favour of a different origin.

**2) An outflow of ionized gas.** This hypothesis is supported by simulations of the theoretical profile of Ly $\alpha$  (Ahn 2004; Verhamme et al. 2006; Behrens et al. 2014). In particular, the models by Behrens et al. (2014) show that a bipolar expanding shell, a reasonable approximation for a nuclear outflow, can produce a secondary red-shifted peak. This unusual line profile is explained by the fact that Ly $\alpha$  is a

resonant emission line, subject to the effects of radiative transfer due to the large cross-section of interaction between Ly $\alpha$  photons and neutral hydrogen atoms. As a consequence, the ubiquity of neutral hydrogen not only within galaxies but in the intergalactic medium along our line of sight prevents most, if not all, light blueward of the central Ly $\alpha$  wavelength from emerging and, as is the case here, only the red side is observed. These effects can also explain the large velocity difference between the Ly $\alpha$  and the He II.

Additionally, the presence of UV continuum aligned with the radio axis is a common feature of HzRGs and has been explained as due to a combination of scattered nuclear light, fast shocks, and jet-induced star formation (see, e.g., McCarthy 1993 or Miley & De Breuck 2008), which also points to an in-host origin for the emission.

With all of these considerations taken into account, we conclude that the outflow hypothesis is the more plausible scenario.

We estimated the mass of the outflow,  $M_{out}$ , from the H $\beta$  luminosity,  $L_{H\beta}$ , using the relation derived by Osterbrock (1989):

$$M_{out} = 7.5 \times 10^{-3} \left( \frac{10^4 L_{H\beta}}{n_e L_{\odot}} \right) M_{\odot} \quad (1)$$

assuming the ratios Ly $\alpha$ /H $\alpha$ =8.7 (Sobral & Matthee 2019) and H $\alpha$ /H $\beta$ =2.86, and a gas density of  $n_e = 200 \text{ cm}^{-3}$ , which enables a comparison between the energetics of this object and a sample of low redshift radio loud AGN that utilizes this same value (Speranza et al. 2022).

We derived the following physical quantities (see e.g. Fiore et al. 2017) for the outflow:

- mass outflow rate:  $\dot{M} = 3v_{out} \frac{M_{H\beta}}{R_{out}} = 3.4M_{\odot} \text{ yr}^{-1}$
- kinetic energy:  $E_{kin} = \frac{1}{2}M_{H\beta}v_{out}^2 = 10^{56.6} \text{ erg}$

- kinetic power  $\dot{E}_{kin} = \frac{1}{2}\dot{M}v_{out}^2 = 10^{42.4} \text{ erg s}^{-1}$

where  $v_{out} = 2,250 \text{ km s}^{-1}$  and  $R_{out} = 18.5 \text{ kpc}$  are the outflow velocity and full extension respectively.

These estimates should be considered as lower limits for two reasons: first of all we cannot correct for the effects of projection and, in addition, there might be a contribution of collisional excitation to the production of emission lines.

#### 4. DISCUSSION AND CONCLUSIONS

The star formation rate (SFR) can be estimated from the ultraviolet (UV) and far infrared (FIR) emission of the source. While the UV light directly traces the young stellar population, the FIR emission is produced by warm dust heated by the ultraviolet photons emitted by young stars. We found the UV contribution to the SFR estimate for the RG to be negligible; it's also worth mentioning that UV continuum is subject to various effects that make it a more unreliable tracer at higher redshift (Wilkins et al. 2012).

Conversely, TN J1049-1258 has been detected by pointed far infrared observations from the Spectral and Photometric Imaging Receiver (SPIRE), onboard the Herschel satellite (Poglitsch et al. 2010; Griffin et al. 2010). At 250  $\mu\text{m}$  and 350  $\mu\text{m}$ , the observed fluxes are  $63.0 \pm 9.5 \text{ mJy}$  and  $67.7 \pm 11.8 \text{ mJy}$ , respectively. These frequencies correspond to  $\sim 50$  and  $70 \mu\text{m}$  in rest frame, which coincides with the peak of the warm dust emission, heated by star forming regions. Measurements from WISE were obtained on the W1 and W2 passbands (W3 and W4 were upper limits). From a fit to the spectral energy distribution (SED) using the WISE and Herschel measurements (see Balmaverde et al. 2016 for details) the derived total FIR luminosity is  $L_{(8-1000\mu\text{m})} = 3.2 \times 10^{13} L_{\odot}$ , which makes this source a hyper luminous infrared galaxy. It corresponds to a prodigious

star formation rate of  $\text{SFR} \sim 4,500 \text{ M}_{\odot} \text{ yr}^{-1}$ , among the highest values ever estimated (see the compilation of SFRs from [Lagache et al. 2018](#)). However, due to the low spatial resolving power of Herschel (the Herschel data reduction guide suggested to adopt an aperture photometry radius for the  $250 \mu\text{m}$  and  $350 \mu\text{m}$  images of  $22''$  and  $30''$ , respectively), both companions and the  $z=0.87$  galaxy are included within the aperture and thus there may be contamination from these sources. Additionally, there might be an AGN contribution. However, the torus emission peaks at  $20 \mu\text{m}$  to  $30 \mu\text{m}$  ([Nenkova et al. 2008](#)) and in order to reproduce the FIR fluxes this component would exceed the WISE measurements. We conclude that the AGN can only be a minor contributor to the Herschel fluxes.

Regarding the energetics of the system, the jet power is  $\sim 4$  orders of magnitude larger than the outflow kinetic power. [Speranza et al. \(2021\)](#) estimated the outflow kinetic power of a sample of radio galaxies at low redshift: while they measured values of  $E_{kin}$  similar to that of TN J1049-1258, they found significantly larger values of  $E_{kin}/P_{jet}$ , typically  $\sim 0.1$ . This suggests that the relativistic jets and the outflow may be decoupled and produced by two independent acceleration mechanisms: the jets might be driven by the magnetic field in the innermost regions of the accretion disk while the outflow is powered by radiation pressure. Alternatively, the ionized gas traced by the  $\text{Ly}\alpha$  emission might just be the tip of the iceberg of a larger amount of outflowing gas, because this line is a poor tracer in terms of representative masses and luminosities. Dedicated radiative transfer simulations with geometries and parameters resembling those observed in TN 1049-1258 will be essential to fully model the results of these observations. To obtain a comprehensive view of the energetic of the outflow, we should also consider other ionized gas tracers, e.g. the optical [O III] line, unaffected by resonance and ab-

sorption effects. Moreover, a relevant contribution to the energetic of the outflow is expected to originate from the molecular gas component, which can be studied using lines produced by the CO molecule transitions or the [C II]  $158 \mu\text{m}$  emission.

Another source of uncertainty in the outflow estimates are the ionization mechanisms in play, as collisional excitation may contribute. In fact, the low  $\text{Ly}\alpha/\text{He II}$  ratio observed in the extended emission is not dissimilar to the one obtained in [Scarlata et al. 2009](#) where, within another high redshift ( $z=2.38$ ) system of Lyman Alpha Emitters, one of the sources displayed a comparably low ratio ( $\text{Ly } \alpha/\text{He II}=2.2 \pm 1.0$ ) and the mechanism of origin was suspected to be accretion of cold gas onto the dark matter halo, resulting in collisional excitation.

Despite the large amount of energy injected into the host galaxy by the ionized outflow and the relativistic jets, TN J1049-1258 displays globally a very large star formation rate, although the companion galaxies might also contribute to the Herschel-measured flux. Rest frame FIR observations at high spatial resolution, e.g., those that can be produced by facilities such as the Atacama Large Millimeter/submillimeter Array (ALMA), will be fundamental to obtain a more accurate estimate of the SFR in the host galaxy, and if possible, resolve its spatial distribution in a bid to better understand the interplay between jets, outflows, and gas reservoirs.

The possibility of outflows and more generally, jets, inducing star formation has been studied in both galactic (see [Bicknell et al. 2000](#); [Miley & De Breuck 2008](#); [Capetti et al. 2022](#)) and cosmological scales ([Gilli et al. 2019](#)), as well as simulations ([Gaibler et al. 2012](#)). As such, it is possible that the observed star formation from the RG is either induced or enhanced by the jet. Spatially resolved observations of cold molecular gas could answer questions such as where the



star formation is occurring: whether jet-driven shocks are pressurizing gas bubbles in the intergalactic medium and causing fragmentation, or the expanding cocoon of the jet is inducing similar phenomena within the host galaxy itself. The co-spatiality between UV continuum and outflow suggests the former, though a mixture of both is not out of the question.

The presence of TN J1049-1258 inside a protocluster is within expectations. Studies have found that up to 75% of powerful HzRGs reside within these structures (Venemans et al. 2006). This is understandable considering that the powerful AGN responsible for this radio emission require massive galaxies, which in turn inhabit rich, dense environments that act as fer-

tile soil for cluster-like structures (Hatch et al. 2014). There is a great deal of relevance in protoclusters within the  $\Lambda$ CDM model, as they represent massive structures hypothetically linked by large dark matter halos: understanding how high- $z$  and nearby galaxy clusters relate to each other is essential in unveiling the behaviour and characteristics of dark matter. Their evolution through the cosmic ages as well as their relationship to powerful RLAGN which likely become the brightest cluster galaxies are ongoing and widely studied topics within cosmology and high-redshift astrophysics.

*Facilities:* VLT(MUSE), Herschel, VLA, WISE

## REFERENCES

- Ahn, S.-H. 2004, *The Astrophysical Journal*, 601, L25, doi: [10.1086/381750](https://doi.org/10.1086/381750)
- Alexander, D. M., Swinbank, A. M., Smail, I., McDermid, R., & Nesvadba, N. P. H. 2010, *Monthly Notices of the Royal Astronomical Society*, 402, 2211, doi: [10.1111/j.1365-2966.2009.16046.x](https://doi.org/10.1111/j.1365-2966.2009.16046.x)
- Antonucci, R. 1993, *ARA&A*, 31, 473, doi: [10.1146/annurev.aa.31.090193.002353](https://doi.org/10.1146/annurev.aa.31.090193.002353)
- Astropy Collaboration, Robitaille, T. P., Tollerud, E. J., et al. 2013, *A&A*, 558, A33, doi: [10.1051/0004-6361/201322068](https://doi.org/10.1051/0004-6361/201322068)
- Astropy Collaboration, Price-Whelan, A. M., Sip6ocz, B. M., et al. 2018, *AJ*, 156, 123, doi: [10.3847/1538-3881/aabc4f](https://doi.org/10.3847/1538-3881/aabc4f)
- Balmaverde, B., Marconi, A., Brusa, M., et al. 2016, *A&A*, 585, A148, doi: [10.1051/0004-6361/201526694](https://doi.org/10.1051/0004-6361/201526694)
- Behrens, C., Dijkstra, M., & Niemeyer, J. C. 2014, *Astronomy & Astrophysics*, 563, A77, doi: [10.1051/0004-6361/201322949](https://doi.org/10.1051/0004-6361/201322949)
- Bicknell, G. V., Sutherland, R. S., van Breugel, W. J. M., et al. 2000, *The Astrophysical Journal*, 540, 678, doi: [10.1086/309343](https://doi.org/10.1086/309343)
- Bornancini, C. G., Breuck, C. D., Vries, W. D., et al. 2007, *Monthly Notices of the Royal Astronomical Society*, 378, 551, doi: [10.1111/j.1365-2966.2007.11769.x](https://doi.org/10.1111/j.1365-2966.2007.11769.x)
- Bouwens, R. J., Illingworth, G. D., Blakeslee, J. P., Broadhurst, T. J., & Franx, M. 2004, *ApJL*, 611, L1, doi: [10.1086/423786](https://doi.org/10.1086/423786)
- Cano-Díaz, M., Maiolino, R., Marconi, A., et al. 2012, *Astronomy & Astrophysics*, 537, L8, doi: [10.1051/0004-6361/201118358](https://doi.org/10.1051/0004-6361/201118358)
- Capetti, A., Balmaverde, B., Tadhunter, C., et al. 2022, *Astronomy & Astrophysics*, 657, A114, doi: [10.1051/0004-6361/202141965](https://doi.org/10.1051/0004-6361/202141965)
- Cavagnolo, K. W., McNamara, B. R., Nulsen, P. E. J., et al. 2010, *ApJ*, 720, 1066, doi: [10.1088/0004-637X/720/2/1066](https://doi.org/10.1088/0004-637X/720/2/1066)
- Chambers, K. C., Magnier, E. A., Metcalfe, N., et al. 2016, *arXiv e-prints*, arXiv:1612.05560, doi: [10.48550/arXiv.1612.05560](https://doi.org/10.48550/arXiv.1612.05560)
- Cicone, C., Feruglio, C., Maiolino, R., et al. 2012, *Astronomy & Astrophysics*, 543, A99, doi: [10.1051/0004-6361/201218793](https://doi.org/10.1051/0004-6361/201218793)
- Condon, J. J., Cotton, W. D., Greisen, E. W., et al. 1998, *Astronomical Journal*, 115, 1693, doi: [10.1086/300337](https://doi.org/10.1086/300337)
- Cresci, G., Mainieri, V., Brusa, M., et al. 2015, *The Astrophysical Journal*, 799, 82, doi: [10.1088/0004-637x/799/1/82](https://doi.org/10.1088/0004-637x/799/1/82)
- Ferguson, H. C., Dickinson, M., Giavalisco, M., et al. 2004, *ApJL*, 600, L107, doi: [10.1086/378578](https://doi.org/10.1086/378578)

- Fiore, F., Feruglio, C., Shankar, F., et al. 2017, *Astronomy & Astrophysics*, 601, A143, doi: [10.1051/0004-6361/201629478](https://doi.org/10.1051/0004-6361/201629478)
- Gaibler, V., Khochfar, S., Krause, M., & Silk, J. 2012, *Monthly Notices of the Royal Astronomical Society*, 425, 438, doi: [10.1111/j.1365-2966.2012.21479.x](https://doi.org/10.1111/j.1365-2966.2012.21479.x)
- Gilli, R., Mignoli, M., Peca, A., et al. 2019, *Astronomy & Astrophysics*, 632, A26, doi: [10.1051/0004-6361/201936121](https://doi.org/10.1051/0004-6361/201936121)
- Griffin, M. J., Abergel, A., Abreu, A., et al. 2010, *A&A*, 518, L3, doi: [10.1051/0004-6361/201014519](https://doi.org/10.1051/0004-6361/201014519)
- Hatch, N. A., Wylezalek, D., Kurk, J. D., et al. 2014, *Monthly Notices of the Royal Astronomical Society*, 445, 280, doi: [10.1093/mnras/stt1725](https://doi.org/10.1093/mnras/stt1725)
- Hopkins, A. M. 2006, *The Star Formation History of the Universe*. <https://arxiv.org/abs/astro-ph/0611283>
- Kerutt, J., Wisotzki, L., Verhamme, A., et al. 2022, *A&A*, 659, A183, doi: [10.1051/0004-6361/202141900](https://doi.org/10.1051/0004-6361/202141900)
- Lacy, M., Baum, S. A., Chandler, C. J., et al. 2020, *Publications of the Astronomical Society of the Pacific*, 132, 035001, doi: [10.1088/1538-3873/ab63eb](https://doi.org/10.1088/1538-3873/ab63eb)
- Lagache, G., Cousin, M., & Chatzikos, M. 2018, *A&A*, 609, A130, doi: [10.1051/0004-6361/201732019](https://doi.org/10.1051/0004-6361/201732019)
- Maiolino, R., Gallerani, S., Neri, R., et al. 2012, *Monthly Notices of the Royal Astronomical Society: Letters*, 425, L66, doi: [10.1111/j.1745-3933.2012.01303.x](https://doi.org/10.1111/j.1745-3933.2012.01303.x)
- McCarthy, P. J. 1993, *ARA&A*, 31, 639, doi: [10.1146/annurev.aa.31.090193.003231](https://doi.org/10.1146/annurev.aa.31.090193.003231)
- Miley, G., & De Breuck, C. 2008, *Astronomy and Astrophysics Review*, 15, 67, doi: [10.1007/s00159-007-0008-z](https://doi.org/10.1007/s00159-007-0008-z)
- Nenkova, M., Sirocky, M. M., Nikutta, R., Ivezić, Ž., & Elitzur, M. 2008, *ApJ*, 685, 160, doi: [10.1086/590483](https://doi.org/10.1086/590483)
- Nesvadba, N. P. H., Breuck, C. D., Lehnert, M. D., et al. 2010, *Astronomy & Astrophysics*, 525, A43, doi: [10.1051/0004-6361/201014960](https://doi.org/10.1051/0004-6361/201014960)
- Osterbrock, D. E. 1989, *Astrophysics of gaseous nebulae and active galactic nuclei* (University Science Books)
- Poglitsch, A., Waelkens, C., Geis, N., et al. 2010, *A&A*, 518, L2, doi: [10.1051/0004-6361/201014535](https://doi.org/10.1051/0004-6361/201014535)
- Pouliasis, E., Georgantopoulos, I., Ruiz, A., et al. 2022, *A&A*, 658, A175, doi: [10.1051/0004-6361/202142059](https://doi.org/10.1051/0004-6361/202142059)
- Price-Whelan, A. M., Lim, P. L., Earl, N., et al. 2022, *The Astrophysical Journal*, 935, 167, doi: [10.3847/1538-4357/ac7c74](https://doi.org/10.3847/1538-4357/ac7c74)
- Ramos Almeida, C., Bischetti, M., García-Burillo, S., et al. 2022, *Astronomy & Astrophysics*, 658, A155, doi: [10.1051/0004-6361/202141906](https://doi.org/10.1051/0004-6361/202141906)
- Rodriguez Zaurin, J., Tadhunter, C. N., Rose, M., & Holt, J. 2013, *Monthly Notices of the Royal Astronomical Society*, 432, 138, doi: [10.1093/mnras/stt423](https://doi.org/10.1093/mnras/stt423)
- Rupke, D. S. N., & Veilleux, S. 2011, *The Astrophysical Journal*, 729, L27, doi: [10.1088/2041-8205/729/2/127](https://doi.org/10.1088/2041-8205/729/2/127)
- . 2013, *The Astrophysical Journal*, 775, L15, doi: [10.1088/2041-8205/775/1/L15](https://doi.org/10.1088/2041-8205/775/1/L15)
- Scarlata, C., Colbert, J., Teplitz, H. I., et al. 2009, *The Astrophysical Journal*, 706, 1241, doi: [10.1088/0004-637x/706/2/1241](https://doi.org/10.1088/0004-637x/706/2/1241)
- Sobral, D., & Matthee, J. 2019, *Astronomy & Astrophysics*, 623, A157, doi: [10.1051/0004-6361/201833075](https://doi.org/10.1051/0004-6361/201833075)
- Speranza, G., Balmaverde, B., Capetti, A., et al. 2021, *Astronomy & Astrophysics*, 653, A150, doi: [10.1051/0004-6361/202140686](https://doi.org/10.1051/0004-6361/202140686)
- Speranza, G., Ramos Almeida, C., Acosta-Pulido, J. A., et al. 2022, *A&A*, 665, A55, doi: [10.1051/0004-6361/202243585](https://doi.org/10.1051/0004-6361/202243585)
- Steidel, C. C., Erb, D. K., Shapley, A. E., et al. 2010, *ApJ*, 717, 289, doi: [10.1088/0004-637X/717/1/289](https://doi.org/10.1088/0004-637X/717/1/289)
- Venemans, B. P., Röttgering, H. J. A., Miley, G. K., et al. 2006, *Astronomy & Astrophysics*, 461, 823, doi: [10.1051/0004-6361:20053941](https://doi.org/10.1051/0004-6361:20053941)
- Verhamme, A., Schaerer, D., & Maselli, A. 2006, *Astronomy & Astrophysics*, 460, 397, doi: [10.1051/0004-6361:20065554](https://doi.org/10.1051/0004-6361:20065554)
- Villar-Martín, M., Humphrey, A., Breuck, C. D., et al. 2007, *Monthly Notices of the Royal Astronomical Society*, 375, 1299, doi: [10.1111/j.1365-2966.2006.11371.x](https://doi.org/10.1111/j.1365-2966.2006.11371.x)
- Weilbacher, P. M., Palsa, R., Streicher, O., et al. 2020, *A&A*, 641, A28, doi: [10.1051/0004-6361/202037855](https://doi.org/10.1051/0004-6361/202037855)

Wilkins, S. M., Gonzalez-Perez, V., Lacey, C. G., & Baugh, C. M. 2012, *Monthly Notices of the Royal Astronomical Society*, 427, 1490, doi: [10.1111/j.1365-2966.2012.22092.x](https://doi.org/10.1111/j.1365-2966.2012.22092.x)

Willott, C. J., Rawlings, S., Blundell, K. M., & Lacy, M. 1999, *MNRAS*, 309, 1017, doi: [10.1046/j.1365-8711.1999.02907.x](https://doi.org/10.1046/j.1365-8711.1999.02907.x)

Wright, E. L., Eisenhardt, P. R. M., Mainzer, A. K., et al. 2010, *Astronomical Journal*, 140, 1868, doi: [10.1088/0004-6256/140/6/1868](https://doi.org/10.1088/0004-6256/140/6/1868)